

A High Temperature Gas Radiation Module for
Compressible Flow CFD:
Theory and User's Guide

Daniel F. Potter

11 January 2013

Contents

1	List of Symbols	vii
	Preface	xii
2	Introduction	1
	2.1 Structure of the report	1
I	Theory	2
3	The Photaura spectral model	3
	3.1 Spectral radiation coefficients	3
	3.1.1 Monatomic bound-bound transitions	4
	3.1.2 Diatomic bound-bound transitions	14
	3.1.3 Continuum transitions	28
	3.1.4 Uncertainty of the radiation calculation	30
	3.2 Collisional-radiative modelling	31
	3.2.1 Collisional-radiative mechanisms	32
	3.2.2 Master equation formulation and solution	38
	3.2.3 Collisional-radiative model for N ₂ -O ₂ mixtures	40
	3.2.4 Collisional-radiative model for CO ₂ -N ₂ -Ar mixtures	52
	3.2.5 A note on the selection of data sources	57
4	The equilibrium air model	59
5	Other spectral models	60
II	User's Guide	61
6	Installation	62
7	Creating Input Files	63
	7.1 Creating an input file for the baseline modelling	63
	7.2 Creating an input file for the advanced modelling of monatomic plasma	65
	7.3 Examples input file interfacing with PARADE	67
	7.4 Examples input file interfacing with SPRADIAN07	68
8	Examples	69
	8.1 VKI minitorch	69
	8.2 EAST CO ₂ -N ₂ shock tube	69
	8.3 Hayabusa	69

8.4 Rutowski hemisphere 69

References 75

A Diatomic collisional-radiative models 76

A.1	Collisional-radiative model for C_2	77
A.2	Collisional-radiative model for CN	78
A.3	Collisional-radiative model for CO	79
A.4	Collisional-radiative model for N_2	80
A.5	Collisional-radiative model for N_2^+	81

List of Figures

3.1	Components of the equilibrium vacuum ultraviolet absorption coefficient spectra for a 10 km/s shock through 0.1 Torr air.	4
3.2	Comparison of electronic partition function Q_{el} for the monatomic radiators using various levels sets.	7
3.1	<i>(Continued)</i> Comparison of electronic partition function Q_{el} for the monatomic radiators using various levels sets.	8
3.2	Gaussian, Lorentzian and Voigt profiles as a function of the normalised frequency. The Gaussian and Lorentzian profiles have the same half-widths.	9
3.3	Monatomic line half-widths at half-maximum for typical Lunar return peak heating conditions (Fire II $t = 1642.66$ s).	12
3.4	Monatomic line half-widths at half-maximum for a hypothetical high-speed Mars entry trajectory point with a freestream pressure and velocity of 18 Pa and 8 km/s respectively.	13
3.5	Sensitivity of atomic bound-bound emissive power density and intensity for a 10 cm slab of equilibrium air to the atomic line cut-off limit ($p = 1$ atm).	14
3.6	Comparison between SPRADIAN07 and the present work for the spectra of atomic oxygen lines in the range $128 \leq \lambda \leq 133$ nm.	15
3.7	Diagrammatic representations of the (a), (b), (c) and (d) Hund's coupling cases describing the limiting angular momentum interactions for rovibronic transitions.	16
3.8	Absorption coefficient for the CN Violet 0-0 band modelled via Hund's case (a) and Hund's case (b).	18
3.9	Comparison of intensity spectra for the N_2^+ First Negative 0-0 band head calculated with and without $L_{e,J}$	20
3.10	Sensitivity of diatomic bound-bound emissive power density and intensity for a 10 cm slab of equilibrium air to the diatomic line cut-off limit ($p = 1$ atm).	26
3.11	Comparison between SPRADIAN07 and the present work for the spectra of the CN Violet 0-0 band-head in the range $387 \leq \lambda \leq 388.5$ nm.	27
3.12	Post-shock species number density and radiative emission profiles along the stagnation streamline of the Fire II $t = 1634$ s condition ($p_\infty = 2$ Pa, $T_\infty = 195$ K, $u_\infty = 11,360$ m/s).	41
3.13	Comparison of electron impact excitation rate coefficients for atomic nitrogen.	44
3.12	<i>(Continued)</i> Comparison of electron impact excitation rate coefficients for various transitions of N.	45

3.13	Comparison of electron impact ionisation rate coefficients for atomic nitrogen.	46
3.14	Comparison of electron impact excitation rate coefficients for atomic oxygen.	47
3.13	<i>(Continued)</i> Comparison of electron impact excitation rate coefficients for atomic oxygen.	48
3.14	Comparison of electron impact ionisation rate coefficients for atomic oxygen.	50
3.15	Comparison of electron impact excitation rate coefficients for transitions to the $B^3\Pi_g$ state of N_2	50
3.16	Comparison of electron impact excitation rate coefficients for transitions to the $B^2\Sigma_u^+$ state of N_2^+	51
3.17	Post-shock species number density and radiative emission profiles along the stagnation streamline for a hypothetical Mars aerocapture entry condition ($p_\infty = 6.2$ Pa, $T_\infty = 161$ K, $u_\infty = 9,440$ m/s).	53
3.18	Comparison of electron impact excitation rate coefficients for atomic carbon.	55
3.17	<i>(Continued)</i> Comparison of electron impact excitation rate coefficients for atomic carbon.	56
3.18	Comparison of electron impact ionisation rate coefficients for atomic carbon.	57
3.19	Comparison of excitation rate coefficients obtained from Park [56, 57] and Zalogin [66].	58

List of Tables

3.1	Summary of monatomic electronic levels from the NIST Atomic Spectra Database [10] implemented in the present work.	6
3.2	Summary of atomic electronic levels and lines from the NIST Atomic Spectra Database [10] implemented in the present work.	8
3.3	Comparison of atomic bound-bound model from the present work with the SPRADIAN07 code [3].	14
3.4	Diatomic systems considered in the present work.	23
3.5	Hönl–London factors for Hund’s case (a).	24
3.6	Hönl–London factors for ${}^2\Sigma\text{--}{}^2\Sigma$ transitions belonging to Hund’s case (b).	24
3.7	Hönl–London factors for ${}^2\Pi\text{--}{}^2\Sigma$ transitions belonging to Hund’s intermediate (a)-(b) case.	25
3.8	Comparison of diatomic bound-bound model from the present work with the SPRADIAN07 code [3].	26
3.9	Comparison of integrated emission and intensity using the transition moments of Hyun [3] and of Babou <i>et al.</i> [27].	28
3.10	Comparison of integrated emission and intensity using the transition moments from Hyun [3] and Chauveau <i>et al.</i> [26].	29
3.11	Summary of the collisional-radiative mechanisms implemented for N and O.	42
3.12	Summary of the collisional-radiative mechanisms implemented for Ar, C, N and O in $\text{CO}_2\text{--N}_2\text{--Ar}$ mixtures.	54
A.1	Implemented collisional-radiative model for C_2	77
A.2	Implemented collisional-radiative model for CN.	78
A.3	Implemented collisional-radiative model for CO.	79
A.4	Implemented collisional-radiative model for N_2	80
A.5	Implemented collisional-radiative model for N_2^+	81

List of Symbols

Roman characters:

A_{ul}, B_{ul}, B_{lu}	Radiative transition probabilities (1/s)
a	speed of sound (m/s)
$b(\nu)$	spectral distribution function
C, n, T_a	generalised Arrhenius parameters (moles/m ³ -s, ND and K)
c	speed of light (2.99792458×10^8 m/s)
c_v	specific heat at constant volume (J/kg-K)
c_p	specific heat at constant pressure (J/kg-K)
D_e	dissociation energy (J/particle)
\tilde{D}_i	mixture average diffusion coefficients (m ² /s)
e	internal energy (J/kg)
E	total energy (J/kg), or radiative power (W)
F_J	rotational energy for level J (J/particle)
f	mass-fraction
f_{lu}	oscillator strength
G_v	vibrational energy for level v (J/particle)
G_i	gaunt factor for electronic level i , or Gibbs free energy for species i (J/kg)
g	degeneracy
h	enthalpy (J/kg), or Planck constant (6.626076×10^{-34} J-s)
I	radiative intensity (W/cm ² -sr) or ionisation potential (J/particle)
I_H	Ionisation potential of the hydrogen atom ($2.1787113 \times 10^{-18}$ J/particle)

I_ν	spectral intensity ($\text{W}/\text{cm}^2\text{-}\mu\text{m}\text{-sr}$)
J	diffusion flux ($\text{kg}/\text{m}^2\text{-s}$), radiative emissive power density (W/cm^3) or rotational quantum number
j	total emissivity ($\text{W}/\text{m}^3\text{-sr}$)
j_ν	spectral emission coefficient ($\text{W}/\text{m}^3\text{-sr}\text{-Hz}$)
K	conductivity ($\text{W}/\text{m}^2\text{-K}$) or collisional-radiative rate coefficient ($\text{particles}/\text{m}^3\text{-s}$)
K_c	Chemical kinetic equilibrium constant based on concentrations
K_p	Chemical kinetic equilibrium constant based on partial pressures
k_f, k_b	Forward and backwards chemical kinetic rate coefficient ($\text{moles}/\text{m}^3\text{-s}$)
$L_{e,J}$	Line alternation factor
M	molecular weight (kg/mol)
m	particle mass (kg)
m_e	electron mass ($9.109390 \times 10^{-28} \text{ kg}$)
N	number density (m^{-3}) or general integer quantity
p_{atm}	atmospheric pressure ($101,325 \text{ Pa}$)
p	pressure (Pa)
q	heat-flux (W/m^2)
Q	partition function (m^{-3})
R	specific gas constant ($\text{J}/\text{kg}\text{-K}$)
R_e	electronic transition moment
S	Spin angular momentum quantum number
$S_{J_l}^{J_u}$	Hönl-London factor for rotational transition $J_u \rightleftharpoons J_l$
s	entropy ($\text{J}/\text{kg}\text{-K}$)
t	time (s)
T_i	Electronic term energy for level i ($\text{J}/\text{particle}$)
T	temperature (K)
U	Treanor-Marrone pseudo-temperature (K)
u, v	velocities in the x and y directions (m/s)

\vec{u}	velocity vector (m/s)
V	volume (m ³)
v	vibrational quantum number
X	mole-fraction
x, y, z	x, y, z cartesian coordinates (m)
Z	Particle charge or nonequilibrium rate coefficient correction factor

Greek characters:

$\alpha_e, B_e, B_v, D_e, D_v$	Rotational coupling constants (J)
α_i, β_i	stoichiometric coefficients
ϵ	quantum state energy (J/particle)
δ_0	Kronecker Delta function
γ	Ratio of specific heats, or Gaussian half-width at half-maximum (Hz)
κ_ν	spectral absorption coefficient (1/m)
Λ	Total angular momentum quantum number, or radiative escape factor
λ	wavelength (nm)
μ	viscosity coefficient
ν	frequency (Hz)
$\Omega_{i,j}^{(l,m)}$	collision integral (m ³ /s)
$\Omega_{i,j}^{(l,m)*}$	reduced collision integral
Ω_{VT}	vibration-translation energy exchange source term (W/m ³)
Ω_{VE}	vibration-electron energy exchange source term (W/m ³)
Ω_{VC}	vibration-chemistry energy exchange source term (W/m ³)
Ω_{EC}	electron-chemistry energy exchange source term (W/m ³)
ω	solid angle (steradians)
$\dot{\omega}$	mass production source term (kg/m ³ -s)
$\omega_e, \omega_e x_e, \omega_e y_e, \omega_e z_e$	Klein-Dunham coefficients
ρ	density (kg/m ³)

σ	rigid sphere collision diameter (m), or homonuclear symmetry factor
$\sigma^2 \Omega_{i,j}^{(l,m)*}$	collision cross section (m ²)
τ	characteristic time scale (s) or relaxation time constant (s ⁻¹)
τ	viscous stress tensor (kg/m-s)
Θ	characteristic temperature (K)

Subscripts and acronyms:

CFD	computational fluid dynamics
D	dissociation
e	free electrons
el	(bound) electronic thermal energy mode
eq	equilibrium thermochemistry
ESA	European Space Agency
froz	Frozen thermochemistry
HO	harmonic oscillator
hp	heavy particles
HWHM	half-width at half-maximum
int	internal energy modes
ion	ionised particles
IR	infrared
LTE	local thermodynamic equilibrium
NASA	National Aeronautics and Space Administration
ND	non-dimensional
m	thermal mode
rad	radiation
reac	reactive thermochemistry
rot	rotation thermal energy mode
s	chemical species
TPS	thermal protection system

tr translation-rotation thermal energy mode
trans translation thermal energy mode
UV ultraviolet
v (grouped) vibration thermal energy mode
vib (individual) vibration thermal energy mode
ve vibration-electron-electronic thermal energy mode

Preface

This document is intended to serve as a reference for the high temperature gas radiation module that is part of the University of Queensland's Compressible Flow CFD group's current code collection, [1]. The radiation module began its life in 2004 as a means to compute equilibrium air radiation with the gray gas approximation. In 2006 work began to expand the radiation module to treat high temperature gases in a spectrally resolved manner. Presently the radiation module can implement a number of spectral models, including the original equilibrium air model and an in-house line-by-line model called Photaura. If the user has access to Fluid Gravity's Parade code [2] or KAIST's Spradian07 code [3], an interfacing framework exists so that these programs can also be implemented as the spectral model within code collection.

Acknowledgements

Many thanks to Peter Jacobs, Rowan Gollan and the other Compressible-Flow CFD developers for constructing and maintaining the code collection; without the main CFD codes and supporting libraries the radiation module would not exist!

2

Introduction

The general differential form of the radiative transfer equation in a participating medium can be written as [4]:

$$\underbrace{\frac{1}{c} \frac{\partial I_\nu}{\partial t}}_{\text{temporal var.}} + \underbrace{\frac{\partial I_\nu}{\partial s}}_{\text{spatial var.}} = \underbrace{j_\nu}_{\text{emission}} - \underbrace{\kappa_\nu I_\nu}_{\text{absorption}} - \underbrace{\sigma_{s,\nu} I_\nu + \frac{\sigma_{s,\nu}}{4\pi} \int I_\nu(\hat{s}_i) \phi(\hat{s}_i, \hat{s}) d\Omega_i}_{\text{scattering}} \quad (2.1)$$

This radiation module allows the spectral emission and absorption coefficients, j_ν and κ_ν , to be calculated for high temperatures gases. The radiation transport models within the Eilmer3 and Poshax3 programs make use of this radiation module when solving for the radiative divergence and radiative heat fluxes within a given computational domain. A number of tools also are provided within the radiation module to solve the radiative transfer equation along a line-of-sight for a non-scattering medium. The spectral models and line-of-sight tools can be used via the provided Python programs, or in a user-created Python script by loading the `radpy` module.

2.1 Structure of the report

Part 1 presents the theoretical formulation of the `photaura` spectral model for high temperature gases. In § 3.1 the calculation method for the spectral coefficients is described, while in § 3.2 the collisional-radiative model is described. Part 2 presents a user guide for the radiation module software.

Part I

Theory

The Photaura spectral model

Photaura is spectral model for high temperature gases that can presently treat bound-bound transitions of monatomic and diatomic species in a line-by-line manner and monatomic continuum processes via hydrogenic approximations or tabulated cross-sections. The `photaura` model is described and implemented in the PhD thesis of Potter [5] where good agreement with air and $\text{CO}_2\text{-N}_2$ shock tube spectroscopy measurements in the NASA Ames EAST facility was found. The `photaura` model was also found by Sobbia et al [6] to accurately reproduce intensity spectra measured in an ICP facility with a $\text{N}_2\text{-CH}_4$ test gas. This chapter is comprised on two sections; spectral modelling in § 3.1, and collisional-radiative modelling in § 3.2.

3.1 Spectral radiation coefficients

The calculation of the spectral radiation coefficients, namely the emission j_ν and absorption κ_ν coefficients, are required when solving the radiation transfer equation presented in Equation 2.1. For an ionised gas, there are three types of radiative mechanisms that make contributions to the bulk spectral coefficients:

1. Bound-bound transitions,
2. Bound-free transitions, and
3. Free-free transitions.

Bound-bound radiative transitions occur between two bound electronic states, whilst bound-free and free-free radiative transitions involve a free electron state. Figure 3.1 presents a sample vacuum ultraviolet absorption coefficient spectra with the contributions from bound-bound, bound-free and free-free transitions identified. As bound electronic states are quantised, the spectrum of a bound-bound transition is distributed about a discrete wavelength characterised by the energy gap between the upper and lower states. In contrast, the energy spectrum of bound-free and free-free radiative transitions are distributed into a continuum due to the arbitrary free electron energy. The spectrum of bound-free transitions are further characterised by a limiting wavelength corresponding to the ionisation threshold.

At the most fundamental level, bound-bound transitions in both atoms and molecules occur between two Zeeman states of a hyperfine line due to the change in nuclear spin. In the present work bound-bound transitions are described by a line-by-line model that considers the hyperfine structure where necessary. Continuum transitions are described

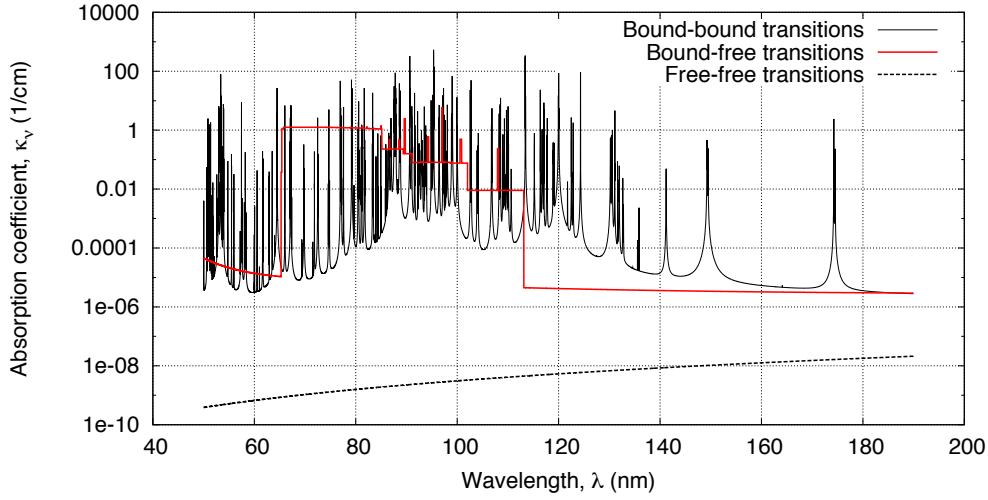


Figure 3.1: Components of the equilibrium vacuum ultraviolet absorption coefficient spectra for a 10 km/s shock through 0.1 Torr air.

by step models presented in the literature or hydrogenic approximations when unavailable. For an indepth discussion of the theory behind the models implemented here, see the texts of Zel'dovich and Razier [7], Huber and Herzberg [8] and Kovács [9].

3.1.1 Monatomic bound-bound transitions

The spectral emission and absorption coefficients for an atomic or molecular bound-bound transition with energy $h\nu_{ul}$ are:

$$j_{\nu,ul} = \frac{N_u h\nu_{ul} A_{ul}}{4\pi} b_{ul}(\nu), \quad (3.1)$$

and,

$$\kappa_{\nu,lu} = (N_l B_{lu} - N_u B_{ul}) h\nu_{ul} b_{ul}(\nu), \quad (3.2)$$

where l and u denote the lower and upper energy levels, N is the level number density, A_{ul} , B_{lu} and B_{ul} are the Einstein coefficients for spontaneous emission, absorption and induced emission, and $b_{ul}(\nu)$ is the spectral distribution function. The absorption and induced emission B_{ul} Einstein coefficients B_{lu} and B_{ul} can be related to the spontaneous emission Einstein coefficient A_{ul} via the principal of detailed balancing [7]. Equation 3.2 is then expressed as:

$$\kappa_{lu} = \left(N_l \frac{g_u}{g_l} - N_u \right) \frac{c^2}{8\pi\nu_{ul}^2} A_{ul} b_{ul}(\nu) \quad (3.3)$$

Level populations

For monatomic species, the electronic level populations are bound by two limiting distributions:

1. Boltzmann thermal equilibrium distribution, and
2. Saha-Boltzmann ionisation equilibrium distribution.

At thermal equilibrium conditions the electronic levels are populated according to the Boltzmann distribution, where the number density of level i is expressed as:

$$N_i = N_{\text{atom}} \frac{Q_{\text{el-}i}}{Q_{\text{int-atom}}} = N_{\text{atom}} \frac{g_i \exp\left(\frac{-E_i}{kT_{\text{el}}}\right)}{\sum_j^{j_{\text{max}}} g_j \exp\left(\frac{-E_j}{kT_{\text{el}}}\right)}, \quad (3.4)$$

where N_{atom} is the total number density of the atom, E_i is the electronic energy of level i , T_{el} is the electronic temperature and $Q_{\text{int-atom}}$ is the total internal (electronic) partition function¹. Another constraint is imposed by considering chemical equilibrium between the electronic level, ions and free electrons. The Saha equation relates the number densities of an atom, its ion and free electrons via the principle of detailed balancing:

$$\frac{N_{\text{atom}}}{N_{\text{ion}}N_e} = \frac{Q_{\text{atom}}}{Q_{\text{ion}}Q_e} \exp\left(\frac{I_{\text{atom}}}{k_B T_e}\right), \quad (3.5)$$

where T_e is the free electron translation temperature, Q and N are respectively the *total* partition function and total number density of the denoted species and I_{atom} is the ionisation potential of the atom. By substituting the Boltzmann equation for an electronic level, Equation 3.4, into the Saha equation for an atomic species, Equation 3.6, the Saha-Boltzmann equation is obtained:

$$N_i = N_{\text{ion}}N_e \frac{Q_{\text{atom}}}{Q_{\text{ion}}Q_e} \exp\left(\frac{I_{\text{atom}}}{k_B T_e}\right) \frac{g_i \exp\left(\frac{-E_i}{kT_{\text{el}}}\right)}{Q_{\text{int-atom}}} \quad (3.6)$$

In compression flows, the Saha-Boltzmann distribution forms the lower bound and the Boltzmann distribution the upper bound, whilst in expanding flows they are reversed. As thermochemical equilibration occurs, the atom number density approaches that predicted by the Saha equation and the Saha-Boltzmann and Boltzmann distributions converge to the same result.

To model the level populations in nonequilibrium, the rate of all transitions affecting the level must be considered. As all transitions can be grouped into those occurring due to particle collisions and those due to radiative transitions, the nonequilibrium modelling of quantum levels is often referred to as ‘collisional-radiative modelling’. In the present work we consider the electronic levels of neutral atoms to possess nonequilibrium populations, whilst the electronic levels of atomic ions are assumed to be in Boltzmann equilibrium. The collisional-radiative framework is described in Section 3.2.

Electronic level energies and degeneracies

The critical data for calculating monatomic partition functions are the energies and degeneracies of the electronic levels. In the present work these parameters are obtained from the NIST Atomic Spectra Database [10], with data for high lying states of neutral atoms taken from Park [11]. Table 3.1 summarises the total, individual and grouped electronic levels and lines considered for monatomic species in the present work. Following the recommendations of Johnston [12], the majority of levels are included as individual multiplets for maximum precision in the collisional-radiative modelling. For the neutral monatomic species C, N and O levels up to energies of $84,000 \text{ cm}^{-1}$, $108,000 \text{ cm}^{-1}$ and

¹Whereas only the first few electronic levels were retained when calculating the partition function for determining thermodynamic properties, all the electronic levels up to the ionisation limit are included for the spectral coefficient calculations. This is necessary as transitions originating from near the ionisation limit are often very strong, and their populations need to be determined to a high degree of accuracy.

106,000 cm^{-1} respectively are treated individually, with the remaining levels included via the groupings proposed by Park [11]. For neutral Ar levels with energy 120,000 cm^{-1} and less are treated individually, with the remaining grouped according to energy proximity. For the ionic monatomic species significantly less levels are required as only the first few excited states can be excited at the conditions of present interest; the levels for Ar^+ , C^+ , N^+ and O^+ are truncated at energies of 150,000 cm^{-1} , 160,000 cm^{-1} , 160,000 cm^{-1} and 200,000 cm^{-1} respectively. Figure 3.1 compares the electronic partition function for the monatomic radiators using the electronic levels from NIST [10], Park [11] and the present work. For the neutral monatomic species good agreement between all three level sets is achieved at temperatures less than 14,000 K, with the Park and present work level sets rising above the NIST results at higher temperatures. This is due to the Park level sets including super-ionised levels, whereas the NIST level sets have been truncated at the ionisation limit. For the ionic monatomic species the NIST and present work level sets agree for the whole temperature range, indicating the chosen truncated energies are adequate. While the C^+ Park and NIST level sets show good agreement, those for N^+ and O^+ do not. These discrepancies have been found to be due to anomalies in the tabulated level data presented by Park [11] for N^+ and O^+ .

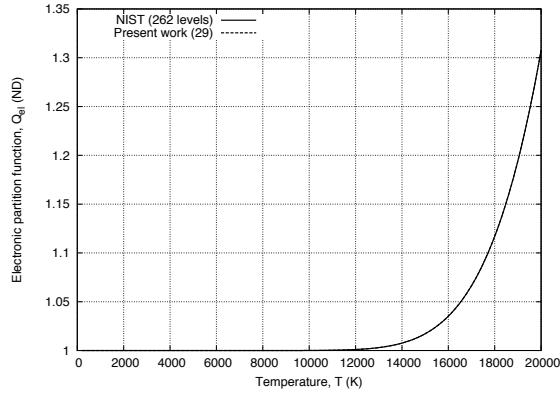
Table 3.1: Summary of monatomic electronic levels from the NIST Atomic Spectra Database [10] implemented in the present work.

Species	Total number of levels	Individual levels	Grouped levels
Ar	29	1 - 17	17 - 29
Ar^+	10	1 - 10	-
C	43	1 - 34	35 - 43
C^+	11	1 - 11	-
N	37	1 - 27	28 - 37
N^+	17	1 - 10	-
O	32	1 - 27	28 - 32
O^+	8	1 - 8	-

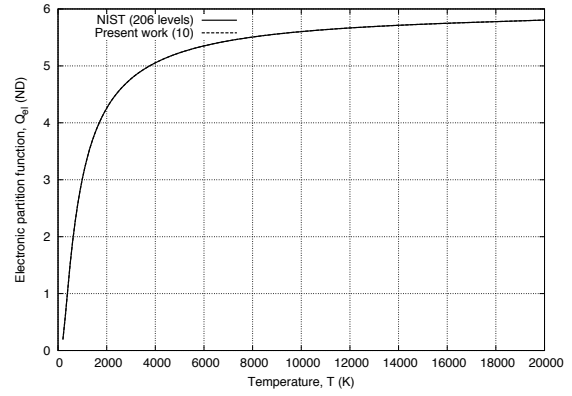
Electronic transitions

Table 3.2 summarises the lines considered for monatomic species in the present work. Following the recommendations of Johnston [12], when performing radiatively coupled Navier–Stokes simulations transitions with energy less than 6 eV are modelled as multiplet lines whilst higher energy transitions are modelled as individual lines. This line selection strategy was shown in Reference [12] to enable the radiant energy to be accurately captured whilst optimising the efficiency of the calculation. It should be noted that the multiplet treatment of spectral lines inevitably leads to some error in the transport calculation, and future work should seek to treat all lines individually if sufficient computational resources are available to make the calculations possible. Also, when performing comparisons with experimental spectra in the present work, all lines are treated individually to best represent the observed spectra. This is possible as the single line-of-sight calculations required for spectra comparisons are not computationally intensive.

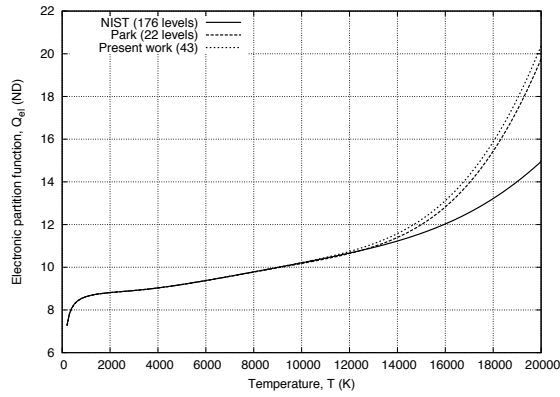
As the electronic level data for each atomic species used for partition function calculations consists of multiplet and grouped levels, a mapping strategy is required for calculating the upper and lower line state populations. This is achieved by assuming Boltzmann equilibrium with the associated multiplet or grouped electronic level. For an upper state of a line denoted by * with associated grouped electronic level i , for example,



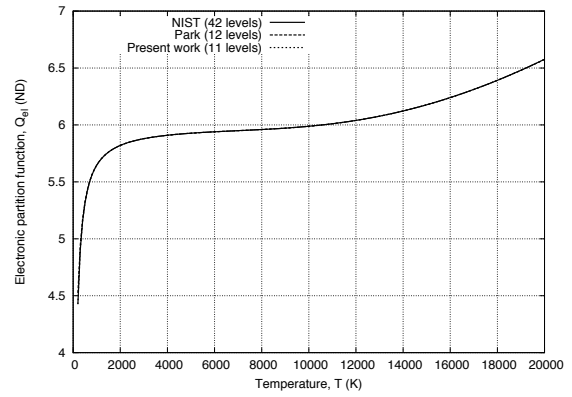
(a) Atomic argon, Ar



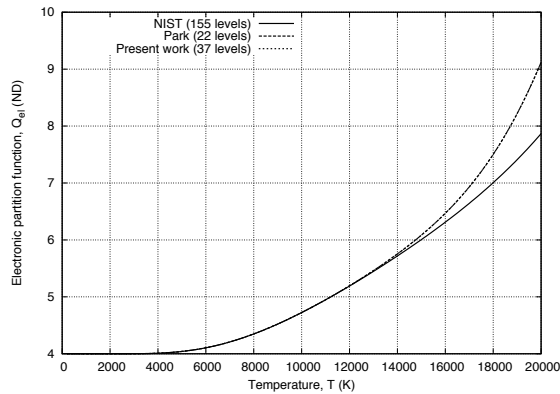
(b) Atomic argon cation, Ar⁺



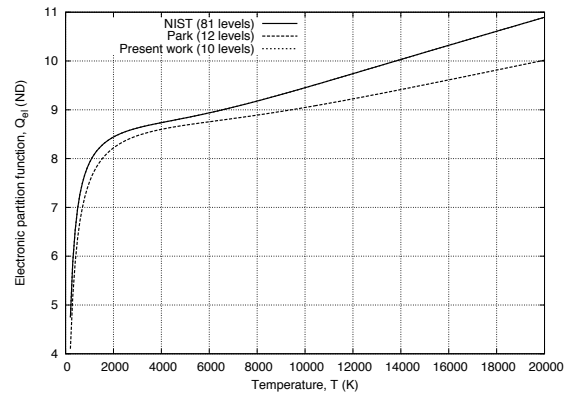
(c) Atomic carbon, C



(d) Atomic carbon cation, C⁺

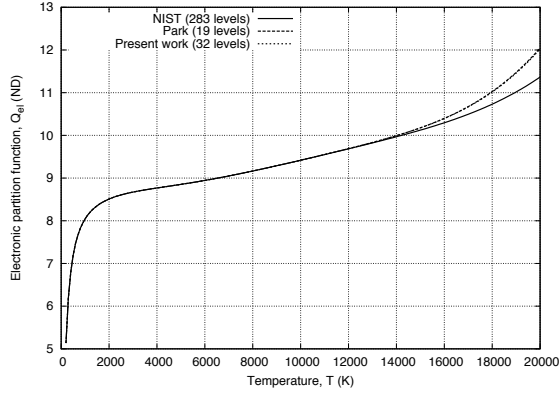


(e) Atomic nitrogen, N

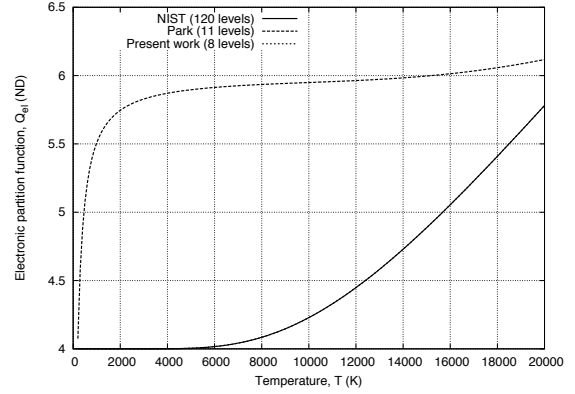


(f) Atomic nitrogen cation, N⁺

Figure 3.2: Comparison of electronic partition function Q_{el} for the monatomic radiators using various levels sets.



(g) Atomic oxygen, O



(h) Atomic oxygen cation, O⁺

Figure 3.1: (*Continued*) Comparison of electronic partition function Q_{el} for the monatomic radiators using various levels sets.

Table 3.2: Summary of atomic electronic levels and lines from the NIST Atomic Spectra Database [10] implemented in the present work.

Species	Number of individual lines		Number of multiplet lines	
	$\Delta E \leq 6 \text{ eV}$	$\Delta E > 6 \text{ eV}$	$\Delta E \leq 6 \text{ eV}$	$\Delta E > 6 \text{ eV}$
Ar	422	6	204	6
Ar ⁺	297	10	98	3
C	1141	157	390	56
C ⁺	358	278	89	69
N	970	129	223	44
N ⁺	481	241	71	89
O	691	163	125	55
O ⁺	617	259	175	77

the upper state population is calculated as:

$$N^* = N_i \frac{g^*}{g_i} \exp \left[\frac{-(E^* - E_i)}{kT_{el}} \right] \quad (3.7)$$

where the associated grouped electronic levels for each line are determined from the NIST tabulations upon initialisation.

Spectral distribution function

The spectral distribution function $b(\nu)$ in Equations 3.1 and 3.2 describes the spectral distribution of the emission and absorption coefficients of a line transition. Although the energy gap characterising a transition is discrete, the energy spectrum of the resulting photon is smeared over a finite range due to various broadening mechanisms. These broadening mechanisms can be classified into two types: those described by a Lorentzian distribution, and those described by a Gaussian distribution. The Lorentzian broadening mechanisms considered in the present work for monatomic radiators are:

- Resonant pressure broadening
- Van der Waals broadening

- Stark broadening
- Natural broadening

The only Gaussian broadening mechanism considered is Doppler broadening. The resultant spectral distribution function is therefore modelled as a Voigt profile which is a convolution of a Lorentzian and a Gaussian distribution. A Gaussian and Lorentzian profile with equal half-widths and the convolved Voigt profile are shown in Figure 3.2. The Gaussian profile exhibits a rapid rise to the central peak, whilst the Lorentzian profile is characterised by slowly decaying ‘wings’.

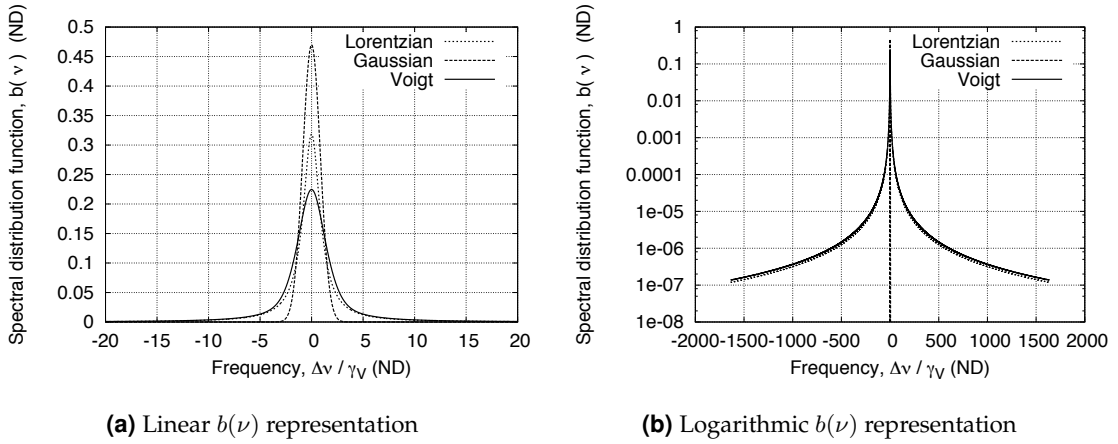


Figure 3.2: Gaussian, Lorentzian and Voigt profiles as a function of the normalised frequency. The Gaussian and Lorentzian profiles have the same half-widths.

In the present work the Voigt profile approximation proposed by Whiting [13] is implemented:

$$b(\nu) = \frac{(1 - R_D) \exp(-2.772R_L^2) + \frac{R_D}{1+4R_L^2} + 0.016(1 - R_D)R_D \exp\left(\frac{-0.4R_L^{2.25}-10}{10+R_L^{2.25}}\right)}{2\gamma_V (1.065 + 0.447R_D + 0.058R_D^2)} \quad (3.8)$$

where R_D and R_L are defined as:

$$R_D = \frac{\gamma_L}{\gamma_V}, \text{ and } R_L = \frac{\nu_{ul}}{2\gamma_V}, \quad (3.9)$$

and γ_L , γ_D and γ_V are respectively the Lorentzian, Doppler and Voigt half-widths at half-maximum (HWHM) in frequency units. The Voigt half-width is a function of the Lorentzian and Doppler (Gaussian) half-widths, and is calculated by the following approximation of Olivero and Longbothum [14]:

$$\gamma_V = (1 - 0.18121(1 - d^2) - [0.023665 \exp(0.6d) + 0.0418 \exp(-1.9d) \sin(\pi d)]) (\gamma_L + \gamma_D) \quad (3.10)$$

where d is defined as:

$$d = \frac{\gamma_L - \gamma_D}{\gamma_L + \gamma_D}. \quad (3.11)$$

The Lorentzian half-width γ_L is the sum of the contributions from the Lorentzian broadening mechanisms:

$$\gamma_L = \gamma_R + \gamma_{VW} + \gamma_S + \gamma_N \quad (3.12)$$

where γ_R , γ_{VW} , γ_S and γ_N are the resonance, Van der Waals, Stark and natural broadening half-widths respectively. Resonant pressure broadening is modelled via the expression of Nicolet [15]:

$$\gamma_R = 3\pi \sqrt{\frac{g_l}{g_u}} \left[\frac{e^2 f_{lu}}{2\pi m \nu_{ul}} \right] N_a \quad (3.13)$$

where f_{lu} is the transition oscillator strength and N_a is the number density of perturbing atoms. In the present work N_a is set to the number density of the lower state. Van der Waals broadening accounts for pressure broadening due to non-resonant interactions, and is modelled by the expression given by Traving [16]:

$$\gamma_{VW} = 1.95 \times 10^{-28} \sqrt{\frac{2T}{M_{av}}} N_{hp} \nu_{ul}^2 \quad (3.14)$$

where M_{av} is the average molecular weight of the mixture and N_{hp} is the heavy particle number density. Although accurate Stark widths for some atomic species are tabulated in the literature (*e.g.* Reference [17]), in the present work Stark broadening is modelled via the following approximate expression observed by Page [18]:

$$\gamma_S = \gamma_S^0 \left(\frac{T_e}{T_e^0} \right)^{\alpha_S} \left(\frac{N_e}{N_e^0} \right) \quad (3.15)$$

where α_S is a fitting constant and γ_S^0 is a reference Stark half-width per electron at electron temperature T_e^0 and electron number density N_e^0 . The reference half-widths are approximated by the following curve-fit proposed by Johnston [12]²:

$$\gamma_S^0 = \frac{8.45 \times 10^9}{(I - E_u)^{2.623}} \quad (3.16)$$

where the reference electron temperature T_e^0 and number density N_e^0 are 10,000 K and $1 \times 10^{16} \text{ cm}^{-3}$ respectively, and the fitting constant α_S is set to 0.33. This curve-fit is shown in Reference [12] to be a good approximation of the accurate N and O Stark widths presented by Griem [17] and others. Natural line broadening is modelled using the following classical expression [19]:

$$\gamma_N = \frac{2\pi e^2 \nu_{ul}^2}{3\epsilon_0 m c^3} \quad (3.17)$$

and Doppler broadening is modelled by the half-width expression given by Nicolet [15]:

$$\gamma_D = \frac{\nu_{ul}}{c} \sqrt{\frac{2k_B T_{tr} \ln(2)}{m_s}} \quad (3.18)$$

where m_s is the species mass per particle.

²Note that the original expression presented by Johnston [12] is for the full-width in wavelength units, whereas that presented here is for the half-width in frequency units.

Figures 3.3a and 3.3b compare the monatomic half-widths calculated at conditions characteristic of typical lunar return peak heating gas states in the boundary and shock layers respectively (Fire II $t = 1642.66$ s). The line widths in the boundary layer are dominated by Doppler broadening with Van der Waals broadening becoming significant at the higher wavelengths, whilst those in the shock layer are largely dominated by Stark broadening. Furthermore the Stark widths in the shock layer are on average approximately 10^3 times greater than in the boundary layer. This is explained by the much higher free electron temperature and density in the shock layer, and that the Stark width is proportional to $T_e^{1/3} N_e$. In both the boundary and shock layers the natural and resonance broadening contributions are negligible for the spectral range considered.

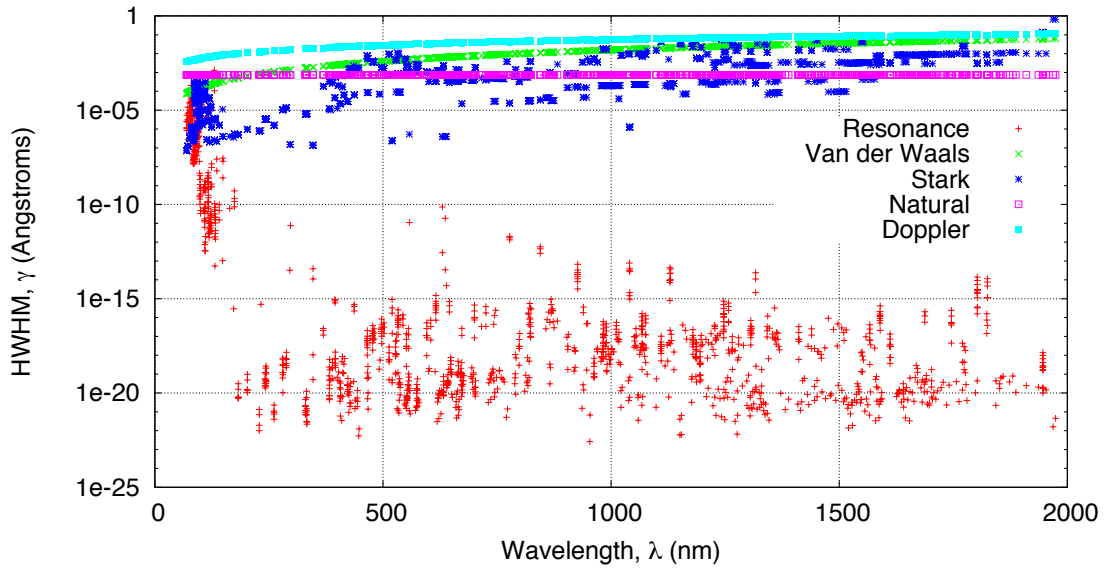
Figures 3.4a and 3.4b compare the monatomic half-widths calculated at conditions characteristic of a hypothetical high-speed Mars entry trajectory point with a freestream pressure and velocity of 18 Pa and 8 km/s respectively. Here the line widths in the boundary layer are also dominated by Doppler broadening, with both Stark and Van der Waals broadening making minor contributions especially at the higher wavelengths. In contrast to the lunar return case, both Stark and Doppler broadening make approximately equal contributions to the line widths for the Mars entry shock layer. This is due to the significantly lower free electron number density and temperature for the Mars entry case. From these results it is evident that both natural and resonance broadening can be omitted without significant loss of line width accuracy for the thermodynamic regimes of present interest.

Finally, an appropriate cut-off limit for each line must be determined. Although the wings of the Voigt profile are many orders of magnitude weaker than the central peak (see Figure 3.2), the wings extend far beyond the line-centre and the rate of decay is low. Figures 3.5a and 3.5b compare the sensitivity of atomic bound-bound emissive power density and intensity for a 10 cm slab of atmospheric pressure equilibrium air to the atomic line cut-off limit. The line cut-off limit $\Delta\nu_{\text{limit}}$ has been normalised by the voigt HWHM γ_V , and the emissive power density and intensity are normalised by the respective values at $\Delta\nu_{\text{limit}} = 10,000\gamma_V$. While the emissive power density is reasonably well described with $\Delta\nu_{\text{limit}}/\gamma_V \geq 10$, the intensity is much more sensitive, requiring $\Delta\nu_{\text{limit}}/\gamma_V \geq 1000$. To optimise the efficiency of the calculation, it is desirable to use the minimum cut-off limit; therefore in the present work the atomic line cut-off limit is set to $\Delta\nu_{\text{limit}} = 1000\gamma_V$.

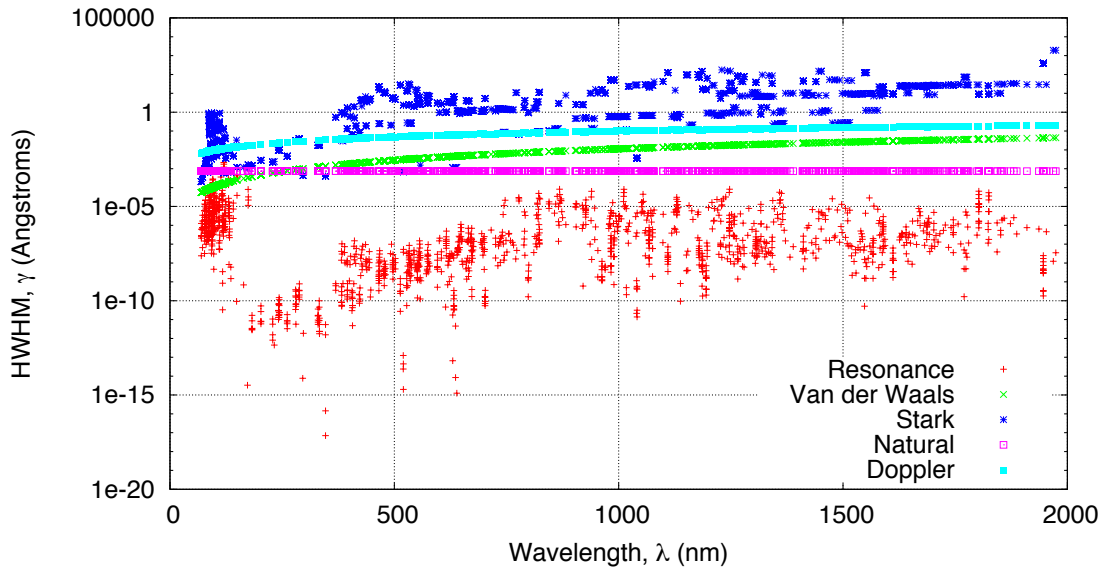
Comparison with SPRADIAN07

The Structured Package for Radiation Analysis 2007 (SPRADIAN07) program has been recently developed by the Japanese Aerospace Exploration Agency (JAXA) and Korea Advanced Institute of Science and Technology (KAIST). The theory and implementation of SPRADIAN07 is described in the PhD thesis of Hyun [3]. Both SPRADIAN07 and the model developed in the present work implement the spectroscopic data from the NIST Atomic Spectra Database [10]. Comparisons with the SPRADIAN07 code [3] have therefore been made in order to verify the calculation of atomic bound-bound spectral coefficients. The test case consists of a 10 cm slab of gas with temperature $T = 10,000$ K and pressure $p = 1$ atm. The number density of each radiator is $1 \times 10^{16} \text{ cm}^{-3}$, the electron number density is also $1 \times 10^{16} \text{ cm}^{-3}$ and total number density is $2.2 \times 10^{17} \text{ cm}^{-3}$. The bound-bound transitions of Ar, Ar⁺, C⁺, N⁺, O⁺ are not included in the comparison as the SPRADIAN07 code does not consider them. For each radiator, the emissive power density J (W/cm^3) and intensity I (W/cm^2) is calculated in the spectral range $50 \leq \lambda \leq 2,000$ nm with 1,950,000 equidistant frequency intervals.

Table 3.3 summarises the comparison between the SPRADIAN07 code [3] and the

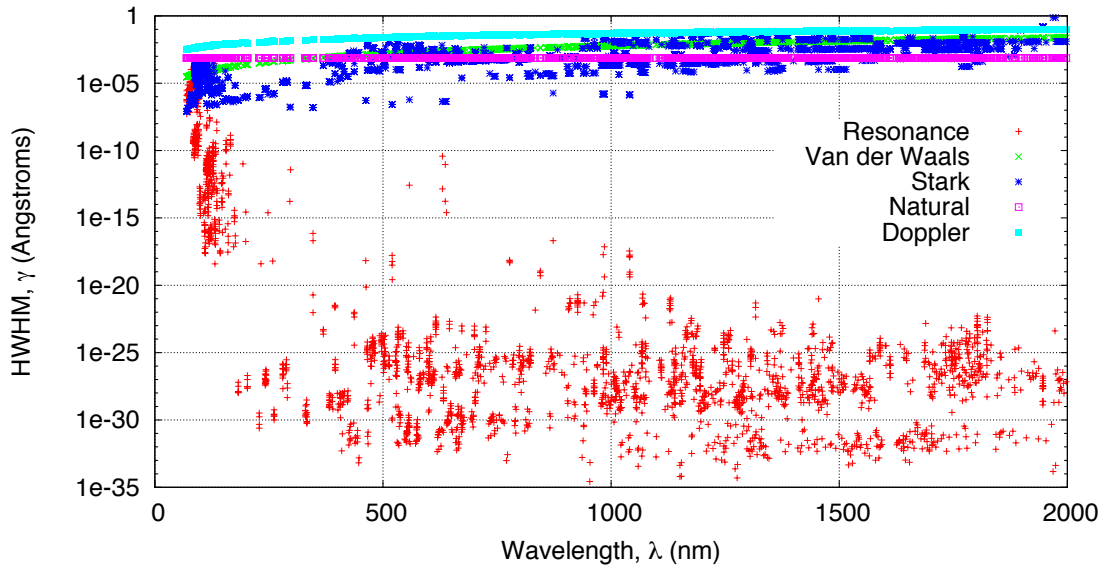


(a) Boundary layer conditions ($T_{tr} = 3,510$ K, $T_{ve} = 3,870$ K, $p = 77,900$ Pa and $N_e = 1.8 \times 10^{13}$ cm $^{-1}$)

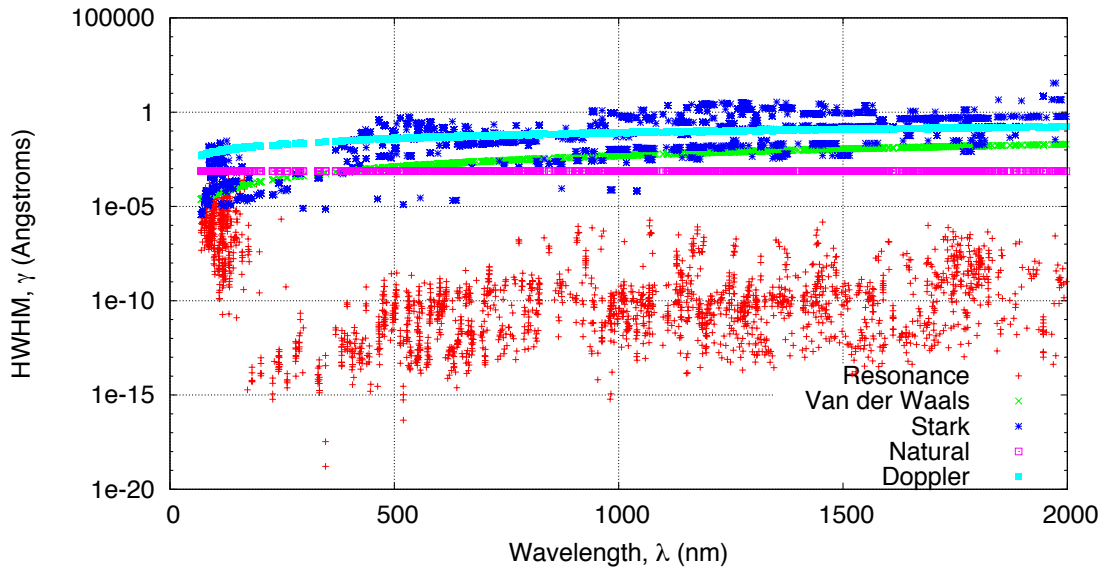


(b) Shock layer conditions ($T_{tr} = 11,220$ K, $T_{ve} = 11,200$ K, $p = 79,000$ Pa and $N_e = 3.6 \times 10^{16}$ cm $^{-1}$)

Figure 3.3: Monatomic line half-widths at half-maximum for typical Lunar return peak heating conditions (Fire II $t = 1642.66$ s).



(a) Boundary layer conditions ($T_{tr} = 2,970$ K, $T_{ve} = 2,550$ K, $p = 29,500$ Pa and $N_e = 2.3 \times 10^{13}$ cm $^{-1}$)



(b) Shock layer conditions ($T_{tr} = 6,930$ K, $T_{ve} = 6,930$ K, $p = 29,500$ Pa and $N_e = 7.7 \times 10^{14}$ cm $^{-1}$)

Figure 3.4: Monatomic line half-widths at half-maximum for a hypothetical high-speed Mars entry trajectory point with a freestream pressure and velocity of 18 Pa and 8 km/s respectively.

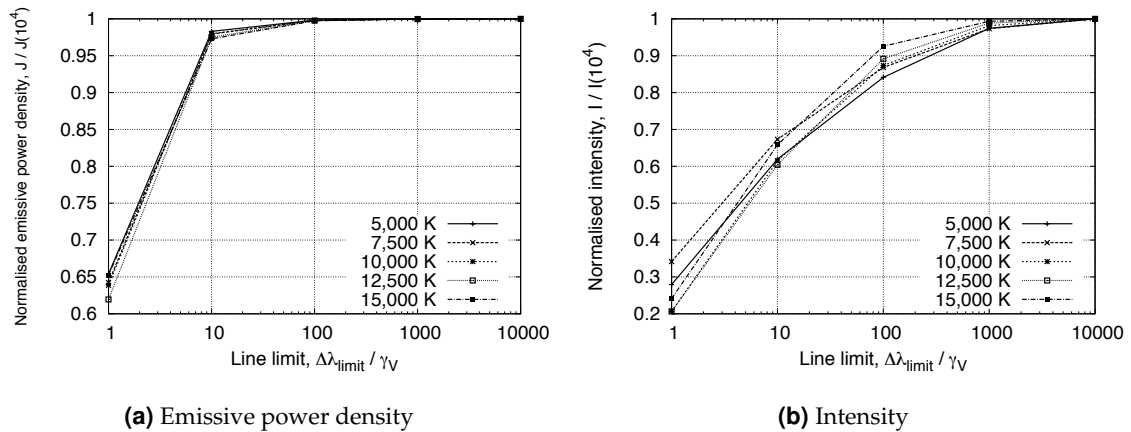


Figure 3.5: Sensitivity of atomic bound-bound emissive power density and intensity for a 10 cm slab of equilibrium air to the atomic line cut-off limit ($p = 1$ atm).

present work for atomic bound-bound transitions. While the agreement for emissive power density is within 1% for all the key atomic radiators, the SPRADIAN07 predicts between 16 and 28% lower intensity through the 10 cm slab. The difference in intensity can be attributed to slight discrepancies in the line half-widths. Figures 3.6a, 3.6b and 3.6c presents the intensity, absorption coefficient and emission coefficient spectra for the atomic oxygen lines in the spectral range $128 \leq \lambda \leq 133$ nm. The SPRADIAN07 emission and absorption spectra peaks higher and drops lower than that from the present work, indicating the SPRADIAN07 line-widths for these transitions are slightly lower. The resultant cumulative intensity is almost 25% lower, however, indicating the high sensitivity of intensity to the line half-widths. While the SPRADIAN07 code uses experimentally determined Stark broadening parameters, the present model uses an approximate curve-fit. Unfortunately the approximate method for calculating the Stark width is a limitation of the spectral model developed for the present work.

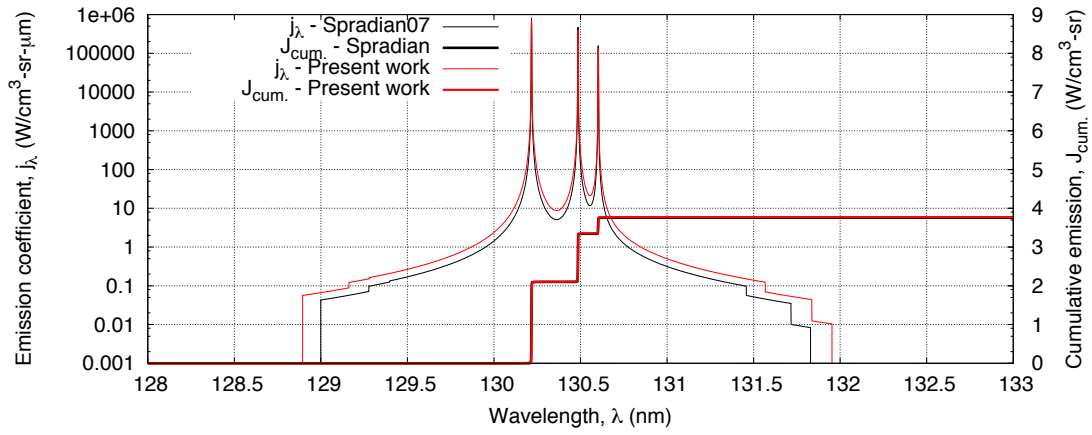
Table 3.3: Comparison of atomic bound-bound model from the present work with the SPRADIAN07 code [3].

Species	Emissive power density, J (W/cm^3)			Intensity, I (W/cm^2)		
	SPRADIAN07	Present work	Difference (%)	SPRADIAN07	Present work	Difference (%)
C	1265	1269	0.34	20.58	25.76	20.12
N	179.8	181.0	0.65	3.77	4.54	16.92
O	59.85	60.28	0.72	1.55	2.13	27.35

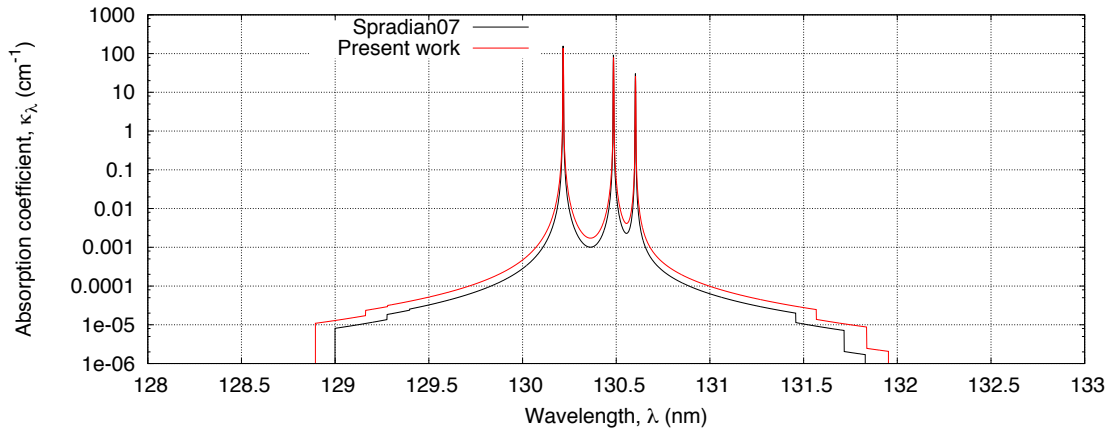
3.1.2 Diatomic bound-bound transitions

Diatomic bound-bound transitions occur between individual rovibronic³ states of the molecule. The resulting rotational lines are clustered into bands and systems corresponding to individual vibrational and electronic transition groups. The spectral emission and absorption coefficients for an individual diatomic bound-bound transition are the same as for monatomic transitions:

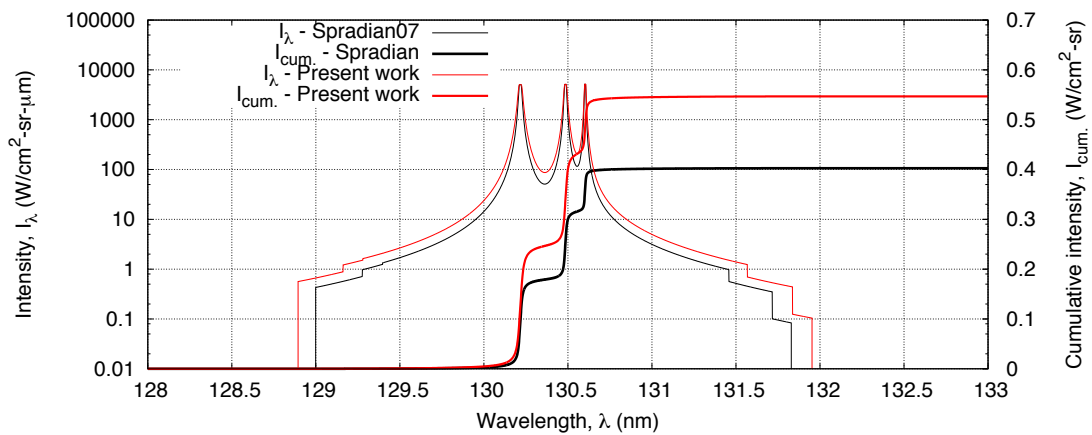
³A rovibronic state is a molecular configuration with a complete set of rotational, vibrational and electronic quantum numbers.



(a) Emission



(b) Absorption



(c) Spectral and cumulative intensity

Figure 3.6: Comparison between SPRADIAN07 and the present work for the spectra of atomic oxygen lines in the range $128 \leq \lambda \leq 133$ nm.

$$j_{\nu,ul} = \frac{n_u h \nu_{ul} A_{ul}}{4\pi} b_{ul}(\nu) \quad (3.19)$$

$$\kappa_{\nu,lu} = \left(n_l \frac{g_u}{g_l} - n_u \right) \frac{c^2}{8\pi \nu_{ul}^2} A_{ul} b_{ul}(\nu) \quad (3.20)$$

where l and u denote the lower and upper energy levels, n is the level number density, A_{ul} is the Einstein coefficient for spontaneous emission and $b_{ul}(\nu)$ is the spectral distribution function.

Rovibronic transitions

The determination of allowed transitions, their energies and probabilities is dependent on the coupling between electronic orbital \vec{L} , electron spin \vec{S} and nuclear rotation \vec{N} angular momentum vectors for the upper and lower rovibronic states. The Hund's coupling cases⁴ (a), (b), (c) and (d) illustrated in Figure 3.7 define idealised limiting cases of angular momentum couplings [8]. In Hund's case (a) nuclear rotation is completely decoupled from electronic motion, whilst electronic motion is strongly coupled to the internuclear axis. In Hund's case (b) electron spin decouples from the internuclear axis due to strong coupling with the rotational motion. When the interaction between the electronic orbital and electron spin angular-momentum is very strong we have Hund's case (c), and when the electronic orbital is strongly coupled to the axis of rotation we have Hund's case (d).

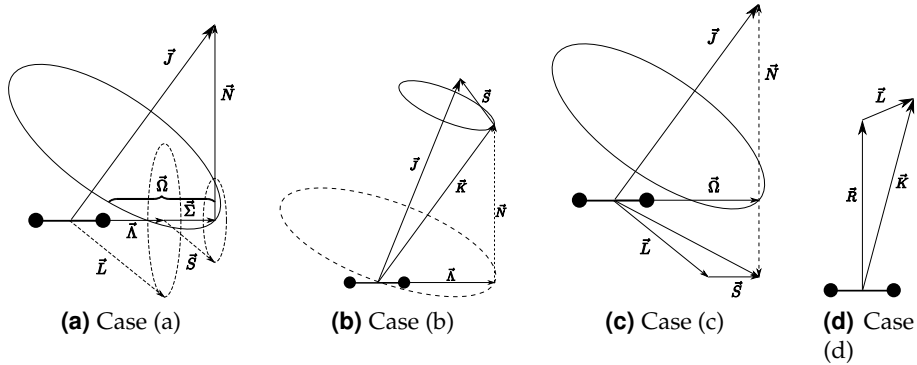


Figure 3.7: Diagrammatic representations of the (a), (b), (c) and (d) Hund's coupling cases describing the limiting angular momentum interactions for rovibronic transitions.

In the present work we consider cases (a) and (b) and an intermediate (a)-(b) case. This selection is a good compromise between speed and accuracy, as spin splitting is captured when important whilst Λ -type doubling which involves much finer perturbations is neglected. Here we will present a brief overview of the selection rules for these three transitions; for a complete discussion the reader is directed to the texts of Huber and Herzberg [8] and Kovács [9]. The energy and transition probability expressions for each case will be presented in the following sections.

For all coupling cases the selection rule for the total angular momentum quantum number J is:

$$J_u - J_l = \Delta J = 0, \pm 1, \text{ and } J_u = J_l \neq 0 \quad (3.21)$$

⁴A fifth coupling case (e) theoretically exists where \vec{L} and \vec{S} are strongly coupled, however such behaviour has not been observed for any species [8].

For Hund's case (a) the electronic angular momentum $\vec{\Omega}$ and rotational angular momentum \vec{N} couple to form the resultant angular momentum vector \vec{J} . Therefore the total angular momentum quantum number J cannot be smaller than the electronic component $\Omega = |\Lambda + \Sigma|$:

$$J = \Omega, \Omega + 1, \Omega + 2, \dots \quad (3.22)$$

An additional restriction for $\Sigma - -\Sigma$ transitions belonging to Hund's case (a) is that $\Delta J = 0$ transitions are universally prohibited. Therefore Hund's case (a) transitions have three branches P , Q and R corresponding to $\Delta J = +1, 0$ and -1 respectively, where only P and R branches exist for $\Sigma - -\Sigma$ transitions. In the present work all singlet ($X^1 - X^1$ where $X = \Sigma, \Pi, \Delta, \dots$) and multiplet parallel ($X^n - Y^n$ where $n > 1$ and $|L_m - L_n| = 1$) transitions except from the CN Violet system are assumed to belong to Hund's case (a).

For Hund's case (b) a total angular momentum quantum number *apart from spin* K is defined with one-to-one correspondence with J as defined for Hund's case (a). Therefore the selection rules pertaining to J outlined above are now applied to K . As $\vec{\Omega}$ and \vec{N} are coupled in Hund's case (b), the permitted range for K is:

$$K = \Lambda, \Lambda + 1, \Lambda + 2, \dots \quad (3.23)$$

As total angular momentum \vec{J} is the resultant of \vec{K} and the spin angular momentum \vec{S} , the possible values of J are:

$$J = (K + S), (K + S - 1), (K + S - 2), \dots, |K - S|. \quad (3.24)$$

Therefore Hund's case (b) considers three ΔJ branches that each consist of $2S + 1$ spin split components. In the present work only the CN Violet $\Sigma^2 - \Sigma^2$ transition is described by Hund's case (b) coupling; for this transition, we have a total of 6 branches with designations $R_1, R_2, {}^R Q_{21}, {}^P Q_{12}, P_2$ and P_1 . Figure 3.8 compares the absorption coefficient for the CN Violet 0-0 band modelled via Hund's case (a) and Hund's case (b). Although the two coupling cases produce similar results for branches close to the band head at 388.45 nm, the effect of spin splitting becomes more pronounced with increasing J .

The remaining transitions, which are the parallel doublets, are described by an intermediate (a)-(b) coupling case, where \vec{S} is strongly coupled to the internuclear axis for low J and becomes coupled with rotation with increasing J – hence this case is referred to as *spin uncoupling*. The quantum numbers K, S and J and their previous defined selection rules are all applicable to the intermediate (a)-(b) case. Intermediate (a)-(b) coupling transitions have three ΔJ branches each with $(2S + 1)^2$ spin split components. A common transition that is modelled via the intermediate (a)-(b) coupling case is a perpendicular doublet such as ${}^2\Pi - {}^2\Sigma$. These transitions have 12 branches with designations $P_1, P_2, P_{12}, P_{21}, Q_1, Q_2, Q_{12}, Q_{21}, R_1, R_2, R_{12}$ and R_{21} .

Level populations

The electronic level populations of molecular species are bounded by two limiting distributions:

1. Boltzmann thermal equilibrium distribution, and
2. Dissociation equilibrium distribution.

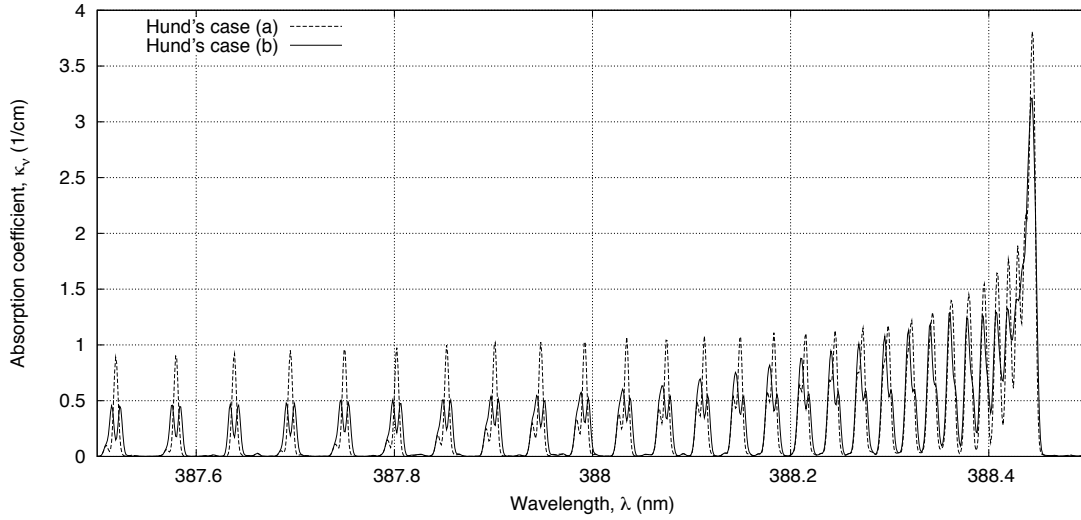


Figure 3.8: Absorption coefficient for the CN Violet 0-0 band modelled via Hund's case (a) and Hund's case (b).

Whereas the chemical equilibrium constraint for atomic species is via ionisation, the chemical equilibrium constraint for molecular species is via dissociation. This is due to the fact that the dissociation energy for a molecule is lower than the ionisation energy; thus a molecule will more readily dissociate before it ionises.

Where sufficient collisions have occurred to achieve thermal equilibrium conditions, the internal quantum states are populated according to the Boltzmann distribution. The number density of *electronic* level i is then:

$$N_i = N_{\text{diatom}} \frac{Q_{\text{el-}i}}{Q_{\text{int-diatom}}} = N_{\text{diatom}} \frac{Q_{\text{el-}i}}{\sum_e^{e_{\text{max}}} Q_{\text{el-}e}}, \quad (3.25)$$

where N_{diatom} is the total species population, $Q_{\text{el-}i}$ is the electronic partition function of level i and $Q_{\text{int-diatom}}$ is the total internal partition function. The electronic partition function for diatomic level i is:

$$Q_{\text{el-}i} = g_i \exp\left(-\frac{T_i}{k_B T_{\text{el}}}\right) \sum_v^{v_{\text{max}}} \exp\left(-\frac{G_v}{k_B T_{\text{vib}}}\right) \frac{1}{\sigma} \sum_J^{J_{\text{max}}} (2J+1) \exp\left(-\frac{F_J}{k_B T_{\text{rot}}}\right), \quad (3.26)$$

where g_i and T_i are the electronic degeneracy and energy, G_v is the energy of vibrational state with quantum number v , F_J is the energy of rotational state with quantum number J and $2J+1$ is the rotational state degeneracy⁵. The electronic degeneracy g_i is the product of the orbital and spin multiplicity of the state:

$$g_i = (2 - \delta_{0,\Lambda_i}) (2S_i + 1) \quad (3.27)$$

where $\delta_{0,\Lambda}$ is the Kronecker Delta function which is unity when $\Lambda = 0$ and zero otherwise and $2S+1$ is the spin multiplicity. The homonuclear factor σ in Equation 3.26 accounts for the symmetry of molecules with alike nuclei, and is equal to 2 for homonuclear molecules and 1 for heteronuclear molecules. To good accuracy the summation over the rotational states can be approximated by the following expression derived by Golden [20]:

⁵The degeneracy of vibrational levels does not appear as it is always unity.

$$Q_{i,v-\text{rot}} = \frac{1}{\sigma} \sum_J^{J_{\text{max}}} (2J+1) \exp\left(-\frac{F_J}{k_B T_{\text{rot}}}\right) \approx \frac{1}{\sigma} \left(\frac{k_B T_{\text{rot}}}{B_e - (v+1/2)\alpha_e}\right), \quad (3.28)$$

where α_e and B_e are spectroscopic constants of the electronic level.

As the characteristic time for chemical reactions is typically much shorter than that for thermal energy exchange, dissociation equilibrium provides another constraint on the population distribution. For a diatomic species comprised of atoms X and Y the dissociation equilibrium relation is found from the principal of detailed balancing:

$$\frac{N_{\text{diatom}}}{N_X N_Y} = \frac{Q_{\text{diatom}}}{Q_X Q_Y} \exp\left(\frac{D_{\text{diatom}}}{k_B T_{\text{tr}}}\right), \quad (3.29)$$

where N and Q denote the total population and total partition function of the indicated species and D_{diatom} is the average dissociation potential of the molecule⁶. The dissociation equilibrium population of electronic level i is found by substituting the Boltzmann relation in Equation 3.25 into Equation 3.29:

$$N_i = N_X N_Y \frac{Q_{\text{diatom}}}{Q_X Q_Y} \exp\left(\frac{D_i}{k_B T_{\text{tr}}}\right) \frac{Q_{\text{el-}i}}{Q_{\text{int-diatom}}} \quad (3.30)$$

where D_i is the dissociation potential taken from electronic level i .

To model the level populations in nonequilibrium, the rate of all transitions affecting the level must be considered. In the present work only *electronic* nonequilibrium is considered, where the electronic levels populations are solved via the collisional-radiative framework to be described in Section 3.2.

Irrespective of the electronic level population distribution, the rotational and vibrational populations are modelled via Boltzmann distributions governed by the respective modal temperatures, T_{rot} and T_{vib} . For a rovibronic level with quantum numbers e , v and J , the Boltzmann population in terms of an arbitrary electronic level population $N_{\text{el-}e}$ is:

$$N_{e,v,J} = N_{\text{el-}e} \frac{Q_{e,v,J} L_{e,J}}{Q_{\text{el-}e} \sigma} \quad (3.31)$$

where $L_{e,J}$ is the line alternation factor due to nuclear spin and $Q_{e,v,J}$ is the rovibronic partition function. $L_{e,J}$ is set to unity for heteronuclear molecules and is a function of the wave function symmetry for homonuclear molecules. Laux [21] gives the line alternation factor for integer nuclear spin (I) as:

$$L_{e,J} = \begin{cases} \frac{I+1}{2I+1} & \text{for } P_{ef} \times P_{gu} \times (-1)^{J^*} = 1 \\ \frac{I}{2I+1} & \text{for } P_{ef} \times P_{gu} \times (-1)^{J^*} = -1 \end{cases} \quad (3.32)$$

and for half integer nuclear spin as:

$$L_{e,J} = \begin{cases} \frac{I}{2I+1} & \text{for } P_{ef} \times P_{gu} \times (-1)^{J^*} = 1 \\ \frac{I+1}{2I+1} & \text{for } P_{ef} \times P_{gu} \times (-1)^{J^*} = -1 \end{cases} \quad (3.33)$$

where P_{ef} is 1 for e parity and -1 for f parity, P_{gu} is 1 for gerade and -1 for ungerade and $J^* = J$ for integer J and $J^* = J - \frac{1}{2}$ for half integer J . The rovibronic state is of e parity if $(-1)^{J^*} \times \text{rotational level parity} > 0$ and -1 otherwise. The rotational level parity for Σ

⁶It is assumed dissociation is governed by the translational temperature T_{tr} , thus the term $\exp(D_{\text{diatom}}/k_B T_{\text{tr}})$.

states is inferred from Figure 114 in the text of Huber and Herzberg [8]. As Λ type doubling is not considered in the present work, the line alternation factors for non- Σ states do not need to be considered. Figure 3.9 compares the intensity spectra of the N_2^+ First Negative 0-0 band head calculated with and without $L_{e,J}$. The spectra calculated via the Specair code of Laux [21,22] is also shown for reference. Apart from slight discrepancies in the calculated line-widths, good agreement with Specair is observed. The alternation of $L_{e,J}$ between $2/3$ $1/3$ for adjacent lines is successfully achieved by the present model.

The rovibronic partition function is:

$$Q_{e,v,J} = g_e \exp\left(-\frac{T_e}{k_B T_{el}}\right) \exp\left(-\frac{G_v}{k_B T_{vib}}\right) \frac{1}{\sigma} (2J+1) \exp\left(-\frac{F_J}{k_B T_{rot}}\right) \quad (3.34)$$

Substituting Equations 3.34, 3.26 and 3.28 into Equation 3.31 yields a simplified expression for $N_{e,v,J}$ that is amenable to numerical implementation:

$$N_{e,v,J} = N_{el-e} \frac{\exp\left(-\frac{G_v}{k_B T_{vib}}\right) (2J+1) \exp\left(-\frac{F_J}{k_B T_{rot}}\right) L_{e,J}}{\sum_v^{\nu_{max}} \exp\left(-\frac{G_v}{k_B T_{vib}}\right) \frac{k_B T_{rot}}{B_e - (v+1/2)\alpha_e} \sigma} \quad (3.35)$$

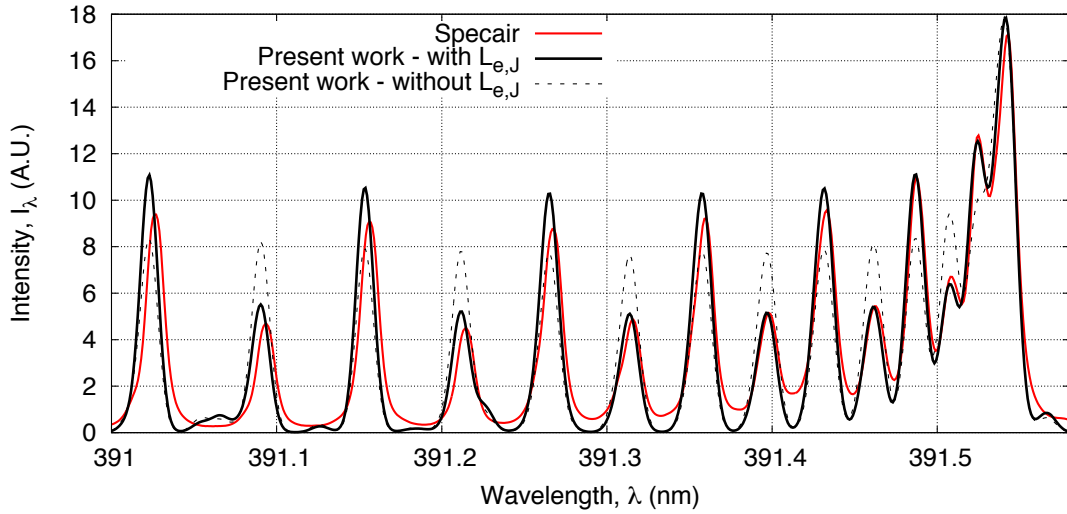


Figure 3.9: Comparison of intensity spectra for the N_2^+ First Negative 0-0 band head calculated with and without $L_{e,J}$.

Maximum vibrational and rotational quantum numbers

When calculating the electronic partition functions in Equation 3.26, the summation over the vibrational and rotational levels must be truncated at ν_{max} and J_{max} respectively. The strategy for determining these parameters described by Babou *et al.* [23] is adopted. The maximum vibrational quantum number ν_{max} is the last that has energy within the dissociation limit referenced from the minimum of the levels potential curve:

$$G_{\nu_{max}} \leq D \quad \text{and} \quad G_{\nu_{max}+1} > D \quad (3.36)$$

For some electronic states the vibrational energy begins to drop before the dissociation limit is reached⁷; the maximum vibrational quantum number in these cases are taken as the last level within the turning point:

$$\frac{\partial G_{v_{\max}}}{\partial v} \geq 0 \quad \text{and} \quad \frac{\partial G_{v_{\max}+1}}{\partial v} \leq 0 \quad (3.37)$$

For each permitted vibrational level $v \leq v_{\max}$ a maximum rotational quantum number J_{\max} must be determined. This is achieved by considering the last rotational level that remains within the potential energy curve:

$$G_v + F_{J_{\max}} \leq V_{J_{\max}}(r_{\max}) \quad \text{and} \quad G_v + F_{J_{\max}+1} > V_{J_{\max}+1}(r_{\max}) \quad (3.38)$$

The potential energy curve is the sum of the Morse and centrifugal potentials:

$$V_J(r) = D \left[1 - \exp\left(-2\beta \frac{r - r_e}{r_e}\right) \right]^2 + B_e \left(\frac{r_e}{r}\right)^2 J(J+1) \quad (3.39)$$

where r_e is the location of potential minimum and β is:

$$\beta = \frac{\omega_e}{4\sqrt{B_e D}}. \quad (3.40)$$

r_{\max} is the location of the potential maximum after the potential minimum, and is therefore found when:

$$\frac{\partial V_J(r_{\max})}{\partial r} = 0 \quad \text{and} \quad \partial^2 \frac{V_J(r_{\max})}{\partial r^2} < 0 \quad (3.41)$$

Rovibronic energies

The energy of a rovibronic level is comprised of electronic T_e , vibrational G_v and rotational F_J contributions. The unperturbed electronic term energies of diatomic species are available directly from the literature (*e.g.* Reference [24]). In contrast, the vibrational G_v and rotational F_J energies are calculated from expressions derived via quantum mechanics. The energy of vibrational level v is calculated by the Dunham expansion which accounts for rigid rotation and anharmonic oscillations:

$$G_v = \omega_e \left(v + \frac{1}{2}\right) - \omega_e x_e \left(v + \frac{1}{2}\right)^2 + \omega_e y_e \left(v + \frac{1}{2}\right)^3 + \omega_e z_e \left(v + \frac{1}{2}\right)^4 + \dots \quad (3.42)$$

where ω_e , $\omega_e x_e$, $\omega_e y_e$ and $\omega_e z_e$ are the Dunham coefficients⁸. The $\omega_e \left(v + \frac{1}{2}\right)$ term represents the contribution from purely harmonic vibration, whilst the higher order terms represent anharmonic corrections. Whilst the anharmonic corrections are neglected for the thermodynamic model, they must be retained for the spectral radiation model in order to produce a high fidelity spectra.

The appropriate expression for the rotational energy is dependent on the coupling case the transition is being modelled by. For Hund's case (a) the fine molecular structure is not considered, and the rotational energy is only a function of the rotational quantum number $J = N$ only:

⁷It should be noted this is not a physical phenomena, but rather an error due to extrapolation of spectroscopic data by the Dunham expansion

⁸Although the Dunham expansion is an infinite series, typically only the first 3 or 4 coefficients are available in the literature [8].

$$F_J = B_v J(J+1) - D_v J^2(J+1)^2, \quad (3.43)$$

where,

$$B_v = B_e(v + \frac{1}{2}) - \alpha_e(v + \frac{1}{2})^2 + \dots, \quad (3.44)$$

$$D_v = D_e(v + \frac{1}{2}) + \beta_e(v + \frac{1}{2})^2 + \dots, \quad (3.45)$$

and B_e and D_e are coupling constants for the electronic state which are also tabulated in the literature.

For a doublet state belonging to Hund's case (b) ($\Sigma^2 - \Sigma^2$ transition) separate expressions are required for the two spin split states:

$$F_{K=J-1/2} = B_v K(K+1) - D_v K^2(K+1)^2 + \frac{1}{2}\gamma K, \quad (3.46)$$

$$F_{K=J+1/2} = B_v K(K+1) - D_v K^2(K+1)^2 - \frac{1}{2}\gamma(K+1), \quad (3.47)$$

where γ is the spin splitting constant for the vibrational band. In the present work the γ values for the CN Violet transition are those presented by Prasad and Bernath [25]. The energies of the two spin split components for doublet states belonging to the intermediate (a)-(b) are:

$$F_{K=J-1/2} = B_v \left[K(K+1) - \Lambda^2 + \frac{Y(4-Y)}{8(K+1)}\Lambda^2 \right] - D_v \left(K + \frac{1}{2} \right)^4, \quad (3.48)$$

$$F_{K=J+1/2} = B_v \left[K(K+1) - \Lambda^2 + \frac{Y(4-Y)}{8K}\Lambda^2 \right] - D_v \left(K + \frac{1}{2} \right)^4, \quad (3.49)$$

where $Y = A/B_v$. For triplet states belonging to the intermediate (a)-(b) case, the energies of the three spin split components are:

$$F_{J=K+1} = B_v \left[J(J+1) - \sqrt{Z_1} - 2Z_2 \right] - D_v \left(J - \frac{1}{2} \right)^4 \quad (3.50)$$

$$F_{J=K} = B_v \left[J(J+1) + 4Z_2 \right] - D_v \left(J + \frac{1}{2} \right)^4 \quad (3.51)$$

$$F_{J=K-1} = B_v \left[J(J+1) + \sqrt{Z_1} - 4Z_2 \right] - D_v \left(J + \frac{3}{2} \right)^4 \quad (3.52)$$

where Z_1 and Z_2 are calculated as:

$$Z_1 = \Lambda^2 Y(Y-4) + \frac{4}{3} + 4J(J+1) \quad (3.53)$$

$$Z_2 = \frac{1}{3Z_1} \left[\Lambda^2 Y(Y-1) - \frac{4}{9} - 2J(J+1) \right] \quad (3.54)$$

Radiative transition probabilities

The radiative transition probability A_{ul} given in Equations 3.19 and 3.20 is calculated as:

$$A_{ul} = \frac{64\pi^4 \nu_{ul}^3}{3hc^3} \frac{(a_0 e)^2 (R_e^{v_u v_l})^2}{(2 - \delta_{0, \Lambda_u}) (2S + 1) 2J_u + 1} S_{J_l}^{J_u} \quad (3.55)$$

where ν_{ul} is the transition frequency in Hz, $(a_0 e)^2 (R_e^{v_u v_l})^2$ is the square of the electronic-vibrational transition moment expressed in statcoulombs and $S_{J_l}^{J_u}$ is the Hönl–London factor for the rotational transition. The electronic-vibrational transition moments proposed by Chauveau *et al.* [26] and Babou *et al.* [27] have been implemented in the present work. These two datasets were selected as they represent the most recent set of transition moments calculated with up-to-date electronic transition moment functions and a consistent treatment of the potential energy function (an RKR potential was used for all species). These diatomic systems and the respective references are summarised in Table 3.4. Note that the additional systems considered by Hyun [3] that are not covered in References [26,27] have also been included.

Table 3.4: Diatomic systems considered in the present work.

Diatomic Species	System name	Transition designation	Included bands (0 : $v_{u,\max}$; 0 : $v_{l,\max}$)	$R_e^{v_u v_l}$ Reference	
CO	Infrared	$X^1\Sigma^+ - X^1\Sigma^+$	(0:50; 0:50)	[27]	
	Fourth-Positive	$A^1\Pi - X^1\Sigma^+$	(0:23; 0:50)	[27]	
	BX (Hopfield–Birge)		$B^1\Sigma^+ - X^1\Sigma^+$	(0:2; 0:50)	[27]
		CX	$C^1\Sigma^+ - X^1\Sigma^+$	(0:9; 0:9)	[3]
	EX	$E^1\Pi - X^1\Sigma^+$	(0:5; 0:5)	[3]	
	FX	$F^1\Sigma^+ - X^1\Sigma^+$	(0:1; 0:0)	[3]	
	GX	$G^1\Pi - X^1\Sigma^+$	(0:2; 0:0)	[3]	
	Third-Positive	$b^3\Sigma^+ - a^3\Pi$	(0:2; 0:18)	[27]	
CO ⁺	Comet–tail	$A^2\Pi_i - X^2\Sigma^+$	(0:33; 0:31)	[27]	
	Baldet–Johnson	$B^2\Sigma^+ - A^2\Pi_i$	(0:33; 0:50)	[27]	
	First Negative	$B^2\Sigma^+ - X^2\Sigma^+$	(0:22; 0:35)	[27]	
CN	Red	$A^2\Pi_i - X^2\Sigma^+$	(0:38; 0:34)	[27]	
	Violet	$B^2\Sigma^+ - X^2\Sigma^+$	(0:25; 0:36)	[27]	
	LeBlanc	$B^2\Sigma^+ - A^2\Pi_i$	(0:25; 0:38)	[27]	
C ₂	Phillips	$A^1\Pi_u - X^1\Sigma_g^+$	(0:35; 0:21)	[27]	
	Mulliken	$D^1\Sigma_u^+ - X^1\Sigma_g^+$	(0:22; 0:21)	[27]	
	Delandres–D’Azambuja		$C^1\Pi_g - A^1\Pi_u$	(0:9; 0:32)	[27]
		Ballik–Ramsay	$b^3\Sigma_g^- - a^3\Pi_u$	(0:41; 0:39)	[27]
	Swan	$d^3\Pi_g - a^3\Pi_u$	(0:18; 0:33)	[27]	
	Fox–Herzberg	$e^3\Pi_g - a^3\Pi_u$	(0:15; 0:35)	[27]	
	Freymark	$E_1\Sigma_g^+ - A^1\Pi_u$	(0:6; 0:4)	[3]	
N ₂	First-Positive	$B^3\Pi_g - A^3\Sigma_u^+$	(0:21; 0:16)	[26]	
	Second-Positive	$C^3\Pi_u - B^3\Pi_g$	(0:4; 0:21)	[26]	
	Birge–Hopfield 1	$b^1\Pi_u - X^1\Sigma_g^+$	(0:19; 0:15)	[26]	

table continued on next page...

table continued from previous page...

Diatomic Species	System name	Transition designation	Included bands (0 : $v_{u,\max}$; 0 : $v_{l,\max}$)	$R_e^{v_u v_l}$	Reference
	Birge–Hopfield 2	$b^1\Sigma_u^+ - X^1\Sigma_g^+$	(0:28; 0:15)		[26]
	Carroll–Yoshino	$c_4^1\Sigma_u^+ - X^1\Sigma_g^+$	(0:8; 0:15)		[26]
	Worley–Jenkins	$c_3^1\Pi_u - X^1\Sigma_g^+$	(0:4; 0:15)		[26]
	Worley	$o_3^1\Pi_u - X^1\Sigma_g^+$	(0:4; 0:15)		[26]
N_2^+	Meinel	$A^2\Pi_u - X^2\Sigma_g^+$	(0:27; 0:21)		[26]
	First–Negative	$B^2\Sigma_u^+ - X^2\Sigma_g^+$	(0:8; 0:21)		[26]
	Second–Negative	$C^2\Sigma_u^+ - X^2\Sigma_g^+$	(0:6; 0:21)		[26]
NO	γ	$A^2\Sigma^+ - X^2\Pi_r$	(0:8; 0:22)		[26]
	β	$B^2\Pi_r - X^2\Pi_r$	(0:37; 0:22)		[26]
	δ	$C^2\Pi_r - X^2\Pi_r$	(0:9; 0:22)		[26]
	ϵ	$D^2\Sigma^+ - X^2\Pi_r$	(0:5; 0:22)		[26]
	γ'	$E^2\Sigma^+ - X^2\Pi_r$	(0:4; 0:22)		[26]
	β'	$B'^2\Delta - X^2\Pi_r$	(0:6; 0:22)		[26]
	11,000 Å	$D^2\Sigma^+ - A^2\Sigma^+$	(0:5; 0:8)		[26]
	Infrared	$X^2\Pi_r - X^2\Pi_r$	(0:22; 0:22)		[26]
O_2	Schumann–Runge	$B^3\Sigma_u^- - X^3\Sigma_g^-$	(0:19; 0:21)		[26]

The Hönl–London factor describes the strength of the rotational lines. The sum of all Hönl–London factors ending in a given lower rotational state must equate to the total degeneracy of the level:

$$\sum_{J_u} S_{J_l}^{J_u} = (2 - \delta_{0, \Lambda_u + \Lambda_l}) (2S_l + 1) (2J_l + 1) \quad (3.56)$$

The selection of the Hönl–London factors therefore depends on the transition type under consideration. The Hönl–London factors for all transitions belonging to Hund’s case (a) are shown in Table 3.5.

Table 3.5: Hönl–London factors for Hund’s case (a).

Branch	$S_{J_l}^{J_u}$ for $\Lambda_u = \Lambda_l = 0$	$S_{J_l}^{J_u}$ for $\Delta\Lambda = 0$	$S_{J_l}^{J_u}$ for $\Delta\Lambda = \pm 1$
P ($\Delta J = +1$)	$J_u + 1$	$\frac{(J_u + 1 + \Lambda_u)(J_u + 1 - \Lambda_u)}{J_u + 1}$	$\frac{(J_u + 1 \mp \Lambda_u)(J_u + 2 \mp \Lambda_u)}{2(J_u + 1)}$
Q ($\Delta J = 0$)	0	$\frac{(2J_u + 1)\Lambda_u^2}{J_u(J_u + 1)}$	$\frac{(J_u \pm \Lambda_u)(J_u + 1 \mp \Lambda_u)(2J_u + 1)}{2J_u(J_u + 1)}$
R ($\Delta J = -1$)	J_u	$\frac{(J_u + \Lambda_u)(J_u - \Lambda_u)}{J_u}$	$\frac{(J_u \pm \Lambda_u)(J_u - 1 \pm \Lambda_u)}{2J_u}$

For ${}^2\Sigma - {}^2\Sigma$ transitions belonging to Hund’s case (b), the Hönl–London factors presented by Mulliken [28] are implemented, Table 3.6.

Table 3.6: Hönl–London factors for ${}^2\Sigma - {}^2\Sigma$ transitions belonging to Hund’s case (b).

Branch	$S_{J_l}^{J_u}$
R ($\Delta J = +1$)	$\frac{(J_l + 1)^2 - \frac{1}{4}}{J_l + 1}$
Q ($\Delta J = 0$)	$\frac{(2J_l + 1)}{4J_l(J_l + 1)}$
P ($\Delta J = -1$)	$\frac{J_l^2 - \frac{1}{4}}{J_l}$

For the parallel doublet transitions modelled by the intermediate (a)-(b) case, the Hönl–London factors used by Arnold *et al.* [29] in the RAD/EQUIL code are imple-

mented, Table 3.7. In this table the upper sign corresponds with upper listed branch and U is defined as:

$$U = [Y^2 - 4Y + (2J + 1)] , \quad (3.57)$$

where $Y = A/B_v$.

Table 3.7: Hönl–London factors for ${}^2\Pi\text{--}{}^2\Sigma$ transitions belonging to Hund’s intermediate (a)-(b) case.

Branch		$S_{J_l}^{J_u}$
${}^2\Pi \Rightarrow {}^2\Sigma$	${}^2\Sigma \Rightarrow {}^2\Pi$	
P_2	R_2	$\frac{(2J_u+1)^2 \pm (2J_u+1)U_u(4J_u^2+4J_u+1-2Y_u)}{16(J_u+1)}$
$O P_{12}$	$S R_{21}$	
$Q P_{21}$	$Q R_{12}$	$\frac{(2J_u+1)^2 \mp (2J_u+1)U_u(4J_u^2+4J_u-7+2Y_u)}{16(J_u+1)}$
P_1	R_1	
Q_2	Q_2	$\frac{(2J_u+1)[(4J_u^2+4J_u-1) \pm U_u(8J_u^3+12J_u^2-2J_u+1-2Y_u)]}{16J_u(J_u+1)}$
$P Q_{12}$	$R Q_{21}$	
$R Q_{21}$	$P Q_{12}$	$\frac{(2J_u+1)[(4J_u^2+4J_u-1) \mp U_u(8J_u^3+12J_u^2-2J_u-7+2Y_u)]}{16J_u(J_u+1)}$
Q_1	Q_1	
R_2	P_2	$\frac{(2J_u+1)^2 \pm (2J_u+1)U_u(4J_u^2+4J_u-7+2Y_u)}{16J_u}$
$Q R_{12}$	$Q P_{21}$	
$S R_{21}$	$O P_{12}$	$\frac{(2J_u+1)^2 \mp (2J_u+1)U_u(4J_u^2+4J_u+1-2Y_u)}{16J_u}$
R_1	P_1	

Spectral distribution function

The spectral distribution function for diatomic lines is the same as that described for monatomic lines in Section 3.1.1, however resonance broadening is not considered. The trends observed for the monatomic linewidths in Figures 3.3 and 3.4 are therefore also applicable to the diatomic linewidths.

Figures 3.5a and 3.5b compare the sensitivity of diatomic bound-bound emissive power density and intensity for a 10 cm slab of atmospheric pressure equilibrium air to the diatomic line cut-off limit. The line cut-off limit $\Delta\nu_{\text{limit}}$ has been normalised by the voigt HWHM γ_V , and the emissive power density and intensity are normalised by the respective values at $\Delta\nu_{\text{limit}} = 1,000\gamma_V$. Both the emissive power density and intensity are reasonably well described with $\Delta\nu_{\text{limit}}/\gamma_V \geq 10$, although significant improvement is achieved with $\Delta\nu_{\text{limit}}/\gamma_V \geq 100$. In the present work the diatomic line cut-off limit is set to $\Delta\nu_{\text{limit}} = 10\gamma_V$ as a compromise between accuracy and efficiency.

Comparison with SPRADIAN07

As was done for monatomic bound-bound transitions, comparisons with the SPRADIAN07 code [3] have been made in order to verify the calculation of diatomic bound-bound spectral coefficients. For this purpose, the electronic-vibrational transition moments presented by Hyun [3] are used so that both codes are using the same fundamental data. Also, the line-alternation factor $L_{e,j}$ for homonuclear molecules is omitted and the CN Violet system is modelled via Hund’s case (a) for consistency with SPRADIAN07. The test case consists of a 10 cm slab of gas with temperature $T = 10,000$ K and pressure $p = 1$ atm. The number density of each radiator is $1 \times 10^{16} \text{ cm}^{-3}$, the electron number density is also $1 \times 10^{16} \text{ cm}^{-3}$ and total number density is $2.2 \times 10^{17} \text{ cm}^{-3}$. For each radi-

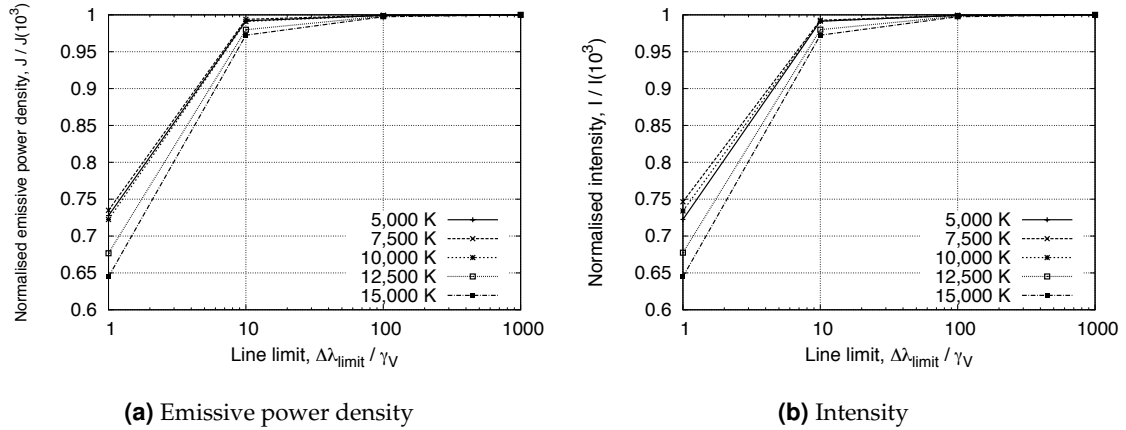


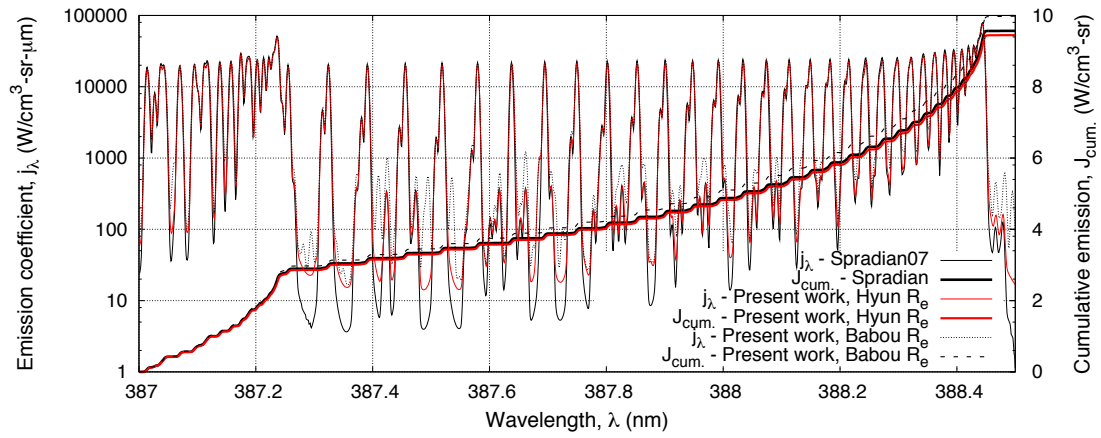
Figure 3.10: Sensitivity of diatomic bound-bound emissive power density and intensity for a 10 cm slab of equilibrium air to the diatomic line cut-off limit ($p = 1$ atm).

ator, the emissive power density J (W/cm^3) and intensity I (W/cm^2) is calculated in the spectral range $50 \leq \lambda \leq 2,000$ nm with 1,950,000 equidistant frequency intervals.

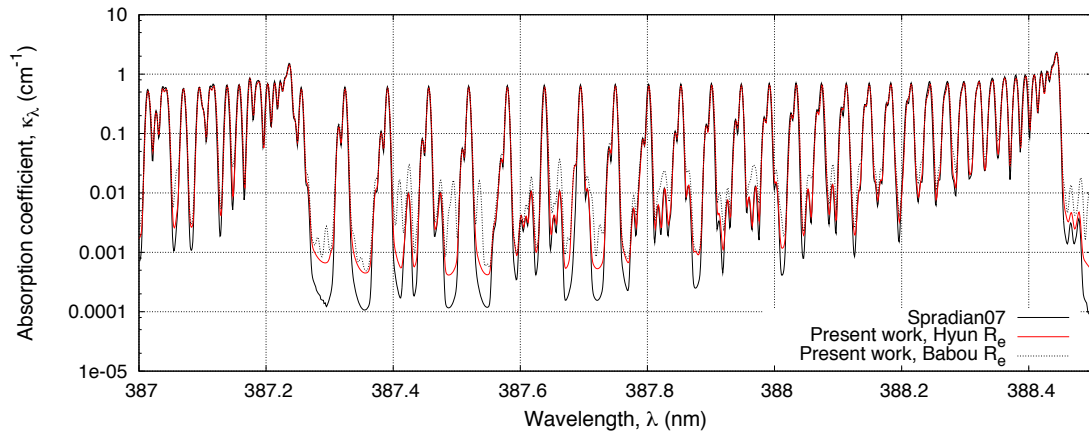
Table 3.8 summarises the comparison between the SPRADIAN07 code [3] and the present work for diatomic bound-bound transitions. The agreement for both emissive power density and intensity is within 5% for all the diatomic radiators considered, with key species such as C_2 , CN and N_2^+ agreeing within 2%. Figures 3.11a, 3.11b and 3.11c presents the emission coefficient, absorption coefficient and intensity spectra for the CN Violet 0-0 band-head in the spectral range $387 \leq \lambda \leq 388.5$ nm. Calculations using the vibration-electronic transition moments of both Hyun [3] and Babou *et al.* [27] are presented. The SPRADIAN07 data exhibits lower troughs between lines, indicating the line-widths are slightly smaller. While the cumulative emission for the Hyun R_e case shows only a 0.4% difference with SPRADIAN07, the differences in line shape between the two coefficient spectrums result in a slightly higher difference in cumulative intensity of 0.6%. Using the transition moments of Babou results in 10% higher intensity, and additional lines appear as a consequence of Babou considering more bands than Hyun. Overall, the agreement with SPRADIAN07 is very good and verifies the implementation of the equations describing bound-bound transitions in the present work.

Table 3.8: Comparison of diatomic bound-bound model from the present work with the SPRADIAN07 code [3].

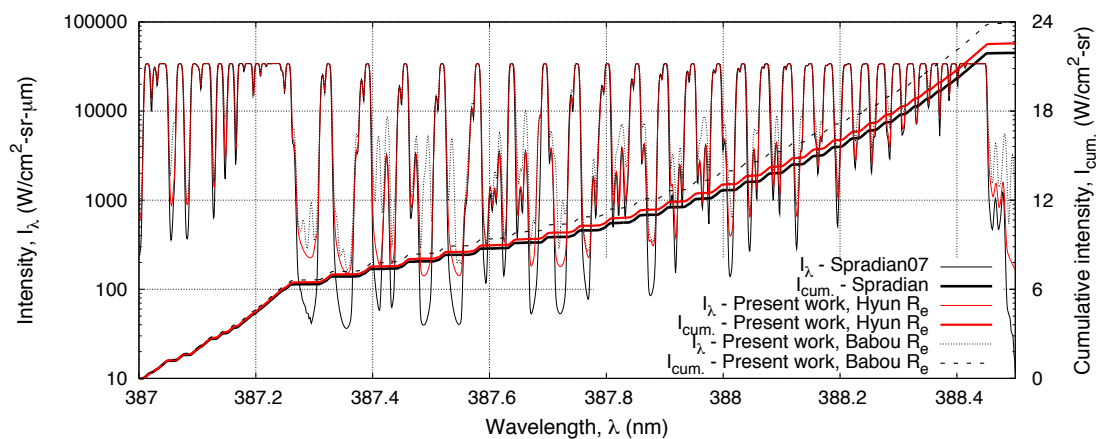
Species	Emissive power density, J (W/cm^3)			Intensity, I (W/cm^2)		
	SPRADIAN07	Present work	Difference (%)	SPRADIAN07	Present work	Difference (%)
C_2	1342	1328	-1.06	772.1	768.5	-0.46
CN	1079	1075	-0.42	470.6	473.6	0.62
CO	584.3	560.8	-4.03	181.2	177.1	-2.32
N_2	8.85	8.48	-4.08	5.61	5.42	-3.60
N_2^+	1235	1244	0.71	646.8	655.7	1.36
NO	112.0	109.0	-2.67	84.44	82.75	-2.04
O_2	78.79	82.31	4.47	61.12	62.51	2.22



(a) Emission



(b) Absorption



(c) Spectral and cumulative intensity

Figure 3.11: Comparison between SPRADIAN07 and the present work for the spectra of the CN Violet 0-0 band-head in the range $387 \leq \lambda \leq 388.5$ nm.

Comparison of transition moment data sets

As the transition moments of Chauveau *et al.* [26] and Babou *et al.* [27] and are being used in preference to the Hyun [3] data set, it is appropriate to compare the two. Table 3.9 presents a comparison of the diatomic species integrated emission and intensities using the Babou⁹ and Hyun transition moments. While C₂, CN, CO, N₂⁺ and O₂ show only minor deviations of 20% or less, NO and N₂ emission are increased by 74% and 787% when using the Babou transition moments.

Table 3.9: Comparison of integrated emission and intensity using the transition moments of Hyun [3] and of Babou *et al.* [27].

Species	Emissive power density, J (W/cm ³)			Intensity, I (W/cm ²)		
	Hyun R_e	Babou R_e	Difference (%)	Hyun R_e	Babou R_e	Difference (%)
C ₂	1328	1479	11.37	768.5	698.6	-10.0
CN	1075	1181	9.91	473.6	569.7	16.9
CO	560.8	473.6	-15.54	177.1	198.8	10.93
N ₂	8.48	75.25	787	5.42	27.3	80.2
N ₂ ⁺	1244	1204	-3.21	655.7	501.6	-30.7
NO	109.0	189.9	74.21	82.75	140.86	41.3
O ₂	82.31	99.66	21.08	62.51	76.12	17.9

The large difference for the N₂ molecule warrants further investigation, especially considering Hyun uses the well regarded data of Laux [30,31] for the N₂ transitions¹⁰. Table 3.10 presents a comparison of the integrated emission and intensities using the Chauveau *et al.* [26] and Hyun [3] transition moments for each system of the N₂ molecule. The well known First-Positive and Second-Positive systems, which radiate in the ultraviolet spectral region, are in good agreement, whilst the remaining systems exhibit significant differences. These systems with large differences — Birge-Hopfield 1, Birge-Hopfield 2, Carroll-Yoshino, Worley-Jenkins and Worley — are all vacuum ultraviolet systems. Due to the difficulty of performing emission spectroscopy in the VUV spectral region, there has been little experimental corroboration for theoretical calculation of the transition moments for these systems. Consequently there is a significant degree of uncertainty associated with the intensity of the N₂ VUV systems, and it is not uncommon for different sets of theoretical calculations to show substantial discrepancies. For example, Liebhart *et al.* calculated the electronic transition moments for N₂ VUV systems via an RKR reconstruction of the potential energy surface. Calculations of equilibrium air absorption spectra at 7,000 K deviated by up to an order of magnitude from the results obtained by Chauveau *et al.* [26], most notably for the strong band peaks at 95 nm. As the transition moments for the N₂ VUV systems used by Hyun [3] are from a yet to be published source, the Chauveau *et al.* [26] data is felt to be more appropriate for the present work.

3.1.3 Continuum transitions

In the present work the continuum transitions for atomic species and their ions are considered, whilst continuum transitions for molecular species are neglected. Furthermore, the models for atomic continuum transitions are based on approximate curve-fits and hydrogenic assumptions. Such a simplified treatment of continuum transitions is justified based on the:

⁹Here ‘Babou’ denotes the set of transition moments described in Table 3.4, which are mainly from Babou *et al.* [27].

¹⁰The calculations for the VUV system transition moments, however, are stated by Hyun [3] to be from Laux but are from a yet to be published source.

Table 3.10: Comparison of integrated emission and intensity using the transition moments from Hyun [3] and Chauveau *et al.* [26].

N ₂ systems	Emissive power density, J (W/cm ³)			Intensity, I (W/cm ²)		
	Hyun R_e	Chauveau R_e	Difference (%)	Hyun R_e	Chauveau R_e	Difference (%)
First-Positive	0.75	0.83	-9.33	0.60	0.66	-9.30
Second-Positive	1.18	1.18	-0.02	0.94	0.94	-0.02
Birge-Hopfield 1	57.09	3.88	1372	29.09	2.77	950
Birge-Hopfield 2	6.68	0.40	1566	4.47	0.32	1319
Carroll-Yoshino	5.91	1.65	257	1.86	0.52	259
Worley-Jenkins	2.07	0.17	1094	0.71	0.11	554
Worley	1.58	0.37	329	0.74	0.23	217

1. Low concentration of molecules and their ions for high-speed Earth and Mars entry,
2. Small contribution of continuum mechanisms to optically thin emission in CO₂-N₂ plasmas at temperature less than 15,000 K [27], and
3. Demonstrated efficacy of photoionisation curve-fits for N and O in high temperature air plasmas [32].

This rationale, however, is not valid for the cool boundary layer surrounding an aeroshell, as at low temperatures ($T \lesssim 6000$ K) the photoionisation and photodissociation continua of diatomic species can be significant [27, 33]. As a consequence, the omission of these mechanisms may lead to an underprediction of the radiative energy absorbed or emitted by the boundary layer. In addition, photodetachment processes for negative ions are estimated to be significant at temperatures up to temperatures of 12,000 K [33]. Nevertheless, the present approximate models capture the majority of the continuum transitions to a reasonable degree of accuracy.

Bound-free mechanisms

For atomic species, bound-free mechanisms refers to photoionisation and the inverse recombination process:



The spectral absorption coefficient due to photoionisation (PI) of electronic level i is:

$$\kappa_{\nu,i} = \sigma_{\nu,i} N_i \quad (3.59)$$

where $\sigma_{\nu,i}$ and N_i are the spectral photoionisation cross section and number density for level i . The spectral emission coefficient can then be derived by applying the microscopic reversibility principle [7]:

$$j_{\nu,i}^{PI} = N_{\text{ion}} N_e \frac{2h\nu^3}{c^2} \frac{g_i}{2Q_{\text{ion}}} \left(\frac{h^2}{2\pi m k_B T_e} \right)^{3/2} \sigma_{\nu,i} \exp \left[\frac{I - E_i - h\nu}{k_B T_e} \right] \quad (3.60)$$

Although accurate tabulations of spectral photoionisation cross sections are available via astrophysics databases such as TOPbase [34], the spectral resolution required to correctly implement them is excessive. Two approximate models for calculating the spectral photoionisation cross section are therefore implemented in the present work:

1. A Gaunt-factor corrected hydrogenic model, and

2. A step-model representation of the TOPbase tabulations [12]

Zeldovich and Raizer present the following expression for the hydrogenic spectral photoionisation cross section:

$$\sigma_{\nu,i} = \frac{64}{3\sqrt{3}} \frac{\pi^4 m Z^4 e^1 0}{\nu^3 c h^6 n_{\text{eff},i}} G_i \quad (3.61)$$

where G_i is the corrective Gaunt factor and $n_{\text{eff},i}$ is the effective shell number for level i :

$$n_{\text{eff},i} = \sqrt{\frac{I_{\text{H}}}{I - E_i}} \quad (3.62)$$

In the present work the following Gaunt factor proposed by Zeldovich and Raizer is implemented:

$$G_i = 1 - 0.173 \left(\frac{h\nu}{IZ^2} \right)^{1/3} \left[\frac{2}{n_{\text{eff},i}} \frac{IZ^2}{h\nu} - 1 \right] \quad (3.63)$$

As discussed by Johnston [12], although Equation 3.61 includes the hydrogenic approximations for the level degeneracy and ion partition function, replacing these parameters with their exact values (*e.g.* Reference [35]) does not improve the accuracy of the expression. Johnston postulates that this is because the Gaunt factor expressions may have been calculated for use with the original hydrogenic approximation.

For the first three levels of atomic nitrogen and oxygen, step model representations of the accurate TOPbase tabulations were constructed by Johnston [12]. The cross sections for the remaining levels are approximated by simple power function. For details of the step model the reader is referred to Reference [12] and [32]. In the present work the step model is preferred over the hydrogenic model for N and O.

Free-free mechanisms

Free-free or bremsstrahlung (literally meaning ‘braking radiation’ in German) radiation results from the acceleration of electrons due to the presence of an electric field. The bremsstrahlung absorption coefficient is presented by Zel’dovich and Raizer [7] as:

$$\kappa_{\nu} = \frac{4}{3} \sqrt{\frac{2\pi}{3m_e k_B T_e}} \left(\frac{Z^2 e^6}{h c m_e \nu^3} \right) N_{\text{ion}} N_e, \quad (3.64)$$

The spectral emission coefficient can then be derived via the principle of detailed balancing:

$$j_{\nu} = \frac{8}{3} \sqrt{\frac{2\pi}{3k_B T_e m_e}} \left(\frac{Z^2 e^6}{m_e c^3} \right) n_{\text{ion}} N_e \exp \left[-\frac{h\nu}{k_B T_e} \right]. \quad (3.65)$$

Generally speaking, bremsstrahlung radiation is most significant in the far-IR spectral region due to the negative exponential dependence on frequency.

3.1.4 Uncertainty of the radiation calculation

Many of the parameters required for the calculation of plasma radiation are highly uncertain. For example the transition probabilities proposed by Wiese [36], which the NIST data implemented in the present work is based on, have uncertainties ranging from $\pm 3\%$ to $\pm 75\%$. Kleb and Johnston [37] performed an uncertainty analysis of air radiation for

lunar return shock layers. Epistemic uncertainty was considered for atomic line oscillator strengths, atomic line Stark broadening widths, atomic photoionisation cross sections, negative ion photodetachment cross sections, molecular bands oscillator strengths and electron impact excitation rates. When direct numerical differentiation and Monte Carlo based methods were applied to a hypothetical lunar return peak heating condition at 10.3 km/s, the uncertainty in radiative heat-flux was found to be $\pm 30\%$. The largest contributors to this total uncertainty level were the atomic nitrogen oscillator strengths and Stark widths and the negative ion continuum. In the present work similar oscillator strengths¹¹ were considered, however a Stark widths are modelled via a less accurate method and the negative ion continuum. Therefore the $\pm 30\%$ uncertainty found in Reference [37] can be considered as a lower bound for the `photaura` radiation model.

3.2 Collisional-radiative modelling

Fundamental to the calculation of the spectral radiation coefficients is the determination of the electronic level populations. Under thermal equilibrium conditions, the level populations assume a Boltzmann distribution. When insufficient collisions have occurred for thermal equilibrium to be achieved, the rate equation for each level must be considered. The net population rate of level i (also referred to in the literature as the ‘master-equation’ [11]) is the difference between the rate of transitions moving electrons in and out of the level:

$$\frac{dN_i}{dt} = \left(\frac{dN_i}{dt}\right)_{\text{in}} - \left(\frac{dN_i}{dt}\right)_{\text{out}} \quad (3.66)$$

Although multidimensional simulations with the collisional-radiative equations fully coupled with the flowfield have recently been performed (*e.g.* Reference [38]), this is computationally prohibitive for the time accurate Navier–Stokes calculations performed in the present work. Rather, the quasi-steady-state (QSS) approximation proposed by Park [11] is applied to decouple the solution from the temporal evolution of the flowfield. The QSS approximation assumes the net population rate of level i to be much smaller than the individual incoming and outgoing rates:

$$\frac{dN_i}{dt} \ll \left(\frac{dN_i}{dt}\right)_{\text{in}}, \quad \text{and} \quad \frac{dN_i}{dt} \ll \left(\frac{dN_i}{dt}\right)_{\text{out}}.$$

The net population rate-of-change in Equation 3.66 can therefore be approximated as zero, and the QSS solution is found when the incoming and outgoing rates are balanced:

$$0 = \left(\frac{dN_i}{dt}\right)_{\text{in}} - \left(\frac{dN_i}{dt}\right)_{\text{out}} \quad (3.67)$$

As the population rates are functions of the immediate flow state only, the QSS solution is decoupled from the temporal evolution of the flowfield. However the QSS approximation is known not to be valid for the ground state [11]. The QSS approximation is therefore only applied to the excited levels, and the ground state population is solved by considering the number density balance for the species:

$$N_{i=1} = N_{\text{total}} - \sum_{i=2}^{N_{\text{levels}}} N_i \quad (3.68)$$

¹¹Implemented as transition probabilities in the present work.

3.2.1 Collisional-radiative mechanisms

The collisional-radiative mechanisms considered in the present work are:

1. Heavy particle impact excitation,
2. Electron impact excitation,
3. Heavy particle impact dissociation,
4. Electron impact dissociation,
5. Electron impact ionisation, and
6. Bound-bound radiative transitions.

Free-bound radiative transitions are omitted as the depopulation rates due to bound-bound radiative transitions are considerably more significant for the conditions of interest [12].

Heavy particle impact excitation

Heavy particle impact excitation reactions have the following form:



where M denotes an arbitrary heavy particle species, X denotes an arbitrary atomic species and i and j denote the lower (initial) and upper (final) electronic levels respectively. The net population rate of level i due to heavy particle impact excitation (HPIE) reactions is:

$$\left(\frac{dN_i}{dt} \right)_{\text{HPIE}} = \sum_{j \neq i} K_M(j, i) N_j N_M - \sum_{j \neq i} K_M(i, j) N_i N_M \quad (3.70)$$

In the present work, heavy particle impact excitation processes are only considered for diatomic species as the electron impact mechanisms dominate in the flow regime where atomic nonequilibrium is significant [11, 12]. The forward rate coefficients for heavy particle impact excitation are obtained from the literature in generalised Arrhenius form:

$$K_M(i, j) = C T_x^n \exp \left(-\frac{E_a}{k T_x} \right) \quad (3.71)$$

where the rate controlling temperature for the excitation process is the geometric average of the translational and vibrational temperatures:

$$T_x = \sqrt{T_{\text{trans}} T_{\text{vib}}} \quad (3.72)$$

The backward reaction rate coefficients are calculated via the principal of detailed balancing:

$$K_M(j, i) = \left[K_M(i, j) \frac{Q_i}{Q_j} \right]_{T_{\text{trans}}} \quad (3.73)$$

where Q_i and Q_j are the partition functions of the lower and upper electronic levels respectively. The rate controlling temperature for the de-excitation process is the translational temperature, and therefore $K_M(i, j)$, Q_i and Q_j in Equation 3.73 are evaluated at T_{trans} .

Electron impact excitation

Electron impact excitation reactions have the following form:



The net population rate of level i due to electron impact excitation (EIE) reactions is:

$$\left(\frac{dN_i}{dt} \right)_{\text{EIE}} = \sum_{j \neq i} K_e(j, i) N_j N_e - \sum_{j \neq i} K_e(i, j) N_i N_e \quad (3.75)$$

The forward rate coefficients are assumed to be governed by the free electron temperature T_e , and the backward reaction rate coefficient is calculated via the principal of detailed balancing:

$$K_e(j, i) = \left[K_e(i, j) \frac{Q_i}{Q_j} \right]_{T_e} \quad (3.76)$$

where the de-excitation process is assumed to be governed by the free electron temperature T_e .

Electron impact excitation processes for both diatomic and atomic species are considered in the present work. The forward rate coefficients for diatomic electron impact excitation are either obtained directly from the literature in generalised Arrhenius form, or calculated by integrating cross sections. In the present work, the following method proposed by Park [11] for calculating the diatomic electron impact excitation rate coefficient from the respective cross section is implemented:

$$K_e(i, j) = \frac{S \sum_{v_i} \sum_{v_j} q(v_i, v_j) \exp \left[-\frac{G_{v_i}}{kT_{\text{vib.}}} \right]}{\frac{kT_e}{B_{e,i}} Q_{\text{vib.},i}} \quad (3.77)$$

where $q(v_i, v_j)$ is the Frank-Condon factor for the vibronic transition between v_i and v_j , G_{v_i} is the vibrational energy of vibrational state v_i and $Q_{\text{vib.},i}$ is the vibrational partition function for electronic level i . The parameter S is defined as:

$$S = 5.47 \times 10^{-11} \sqrt{T_e} \exp \left(-\frac{T_j + G_{v,j} - T_i - G_{v,i}}{k_B T_e} \right) \quad (3.78)$$

where I is:

$$I = \left(\frac{B_{e,j} - B_{e,i}}{k_B T_e} + \frac{B_{e,i}}{k_B T_{\text{rot.}}} \right)^{-3/2} [C\Gamma(1.5)BD\Gamma(2.5)] \quad (3.79)$$

with:

$$C = \frac{T_j + G_{v,j} - T_i - G_{v,i}}{k_B T_e} B + A, \quad (3.80)$$

$$D = \frac{B_{e,j} - B_{e,i}}{k_B T_e} \left(\frac{B_{e,j} - B_{e,i}}{k_B T_e} + \frac{B_{e,i}}{k_B T_e} \right)^{-1}. \quad (3.81)$$

The parameters A and B are defined as:

$$A = \int_0^\infty \left[\frac{\sigma(\xi)}{\pi a_0^2} \right] \exp(-\xi) \xi d\xi \quad (3.82)$$

$$B = \int_0^\infty \left[\frac{\sigma(\xi)}{\pi a_0^2} \right] \exp(-\xi) d\xi \quad (3.83)$$

where a_0 is the first Bohr radius and σ is the electron impact excitation cross section with $\xi = x - 1$ and $x = E/E^*$ where E is the electron energy and E^* is the threshold energy (ΔT_e).

Although experimental and theoretical electron impact excitation cross sections for atomic species are available for some transitions from low lying states, for the majority of transitions we must rely on semi-empirical models. Here we will briefly describe the approximate electron impact excitation models considered for atoms, whilst a detailed description of the rate coefficient models selected for each species will be presented later.

Numerous empirical electron impact excitation models for atomic species were investigated in the comprehensive studies of Johnston [12] and Panesi [38]. As a baseline model, Panesi [38] implemented electron impact excitation reaction rate coefficients obtained by analytical integration of the Drawin [39] cross sections over a Maxwell-Boltzmann velocity distribution:

$$K_e(i, j) = \begin{cases} \sqrt{\frac{8k_B T_e}{\pi m_e}} 4\pi a_0^2 \alpha \left(\frac{I_H}{k_B T_e} \right)^2 I_1(a) & \text{for an optically allowed transition} \\ \sqrt{\frac{8k_B T_e}{\pi m_e}} 4\pi a_0^2 \alpha \left(\frac{E_j - E_i}{k_B T_e} \right)^2 I_2(a) & \text{for an optically forbidden transition} \end{cases} \quad (3.84)$$

where m_e is the mass of an electron, a_0 is the first Bohr radius, $\alpha = 0.05$, I_H is the ionisation energy of the hydrogen atom from the ground state and $I_1(a)$ and $I_2(a)$ are calculated as:

$$I_1(a) = 0.63255 a^{-1.6454} e^{-a}, \quad \text{where} \quad a = \frac{E_j - E_i}{k_B T_e} \quad (3.85)$$

$$I_2(a) = 0.23933 a^{-1.4933} e^{-a}, \quad \text{where} \quad a = \frac{E_{\text{ionise},i} - E_i}{k_B T_e} \quad (3.86)$$

The hydrogenic model of Gryzinski [40] was implemented by both Johnston [12] and Panesi [38]. Similarly as for the Drawin model, the Gryzinski model is semi-empirical and universally applicable to all transition types. The electron impact excitation reaction rate coefficients are calculated by integrating the Gryzinski cross section σ_{ij} over a Maxwellian velocity distribution:

$$K_e(i, j) = \frac{8\pi}{\sqrt{m}} \left(\frac{1}{2\pi m k_B T_e} \right)^{1.5} \left[\int_{\Delta E_{i,j}}^\infty \sigma_{i,j}(E) \exp\left(-\frac{E}{k_B T_e}\right) E dE + \int_{\Delta E_{i,j+1}}^\infty \sigma_{i,j+1}(E) \exp\left(-\frac{E}{k_B T_e}\right) E dE \right] \quad (3.87)$$

where $\sigma_{ij}(E)$ is given by Eq. 3.88 for $\Delta E_{ij} + E_{\text{ionise}} - E_i \leq E$ and by Eq. 3.89 for $\Delta E_{ij} + E_{\text{ionise}} - E_i \geq E$:

$$\begin{aligned}\sigma_{i,j}(E) &= \frac{4.2484 \times 10^{-6}}{\Delta E_{ij}^2} \left(\frac{E}{E_{\text{ionise}} - E_i + E} \right)^{1.5} \\ &\times \left\{ \frac{2}{3} \left[\frac{E_{\text{ionise}} - E_i}{E} + \frac{\Delta E_{i,j}}{E} \left(1 - \frac{E_{\text{ionise}} - E_i}{E} \right) - \left(\frac{\Delta E_{i,j}}{E} \right)^2 \right] \right\}\end{aligned}\quad (3.88)$$

$$\begin{aligned}\sigma_{i,j}(E) &= \frac{4.2484 \times 10^{-6}}{\Delta E_{ij}^2} \left(\frac{E}{E_{\text{ionise}} - E_i + E} \right)^{1.5} \\ &\times \left\{ \frac{2}{3} \left[\frac{E_{\text{ionise}} - E_i}{E} + \frac{\Delta E_{i,j}}{E} \left(1 - \frac{E_{\text{ionise}} - E_i}{E} \right) - \left(\frac{\Delta E_{i,j}}{E} \right)^2 \right] \right. \\ &\times \left. \left[\left(1 + \frac{\Delta E_{ij}}{E_{\text{ionise}} - E_i} \right) \left(1 - \frac{\Delta E_{i,j}}{E} \right) \right]^{0.5} \right\}\end{aligned}\quad (3.89)$$

In the present implementation the cross section integration is performed by a change of variables to shift the limits to [-1,1] and applying 10 point Gaussian quadrature.

Johnston [12] and Panesi [38] also considered the empirical models proposed by Allen [41], Van Regmorter [42] and Park [11,43]. Both authors found the Park [11,43] models to give substantially larger rates than other more accurate models for electron impact excitation of N and O, and were therefore not implemented. Although Johnston [12] preferred the models of Allen [41] and Van Regmorter [42] over that of Gryzinski [40], Panesi [38] found the Drawin model described in Equation 3.84 to be in good agreement with these models. Therefore in the present work only the Gryzinski [40] and Drawin [39] empirical models will be considered.

Heavy particle impact dissociation

Heavy particle impact dissociation reactions have the following form:



where A and B are the constituent atoms of the diatomic molecule AB. The net population rate of level i due to heavy particle impact dissociation (HPID) reactions is:

$$\left(\frac{dN_i}{dt} \right)_{\text{HPID}} = K_M(d, i) N_A N_B N_M - K_M(i, d) N_i N_M \quad (3.91)$$

where d denotes the dissociated state. The forward rate coefficients for diatomic heavy particle impact dissociation are obtained from the literature in generalised Arrhenius form, where the rate controlling temperature for the excitation process is the geometric average of the translational and vibrational temperatures:

$$T_x = \sqrt{T_{\text{trans}} T_{\text{vib}}} \quad (3.92)$$

The heavy particle impact recombination rate coefficient is related to the heavy particle impact dissociation rate coefficient via the principal of detailed balancing:

$$K_M(d, i) = \left[K_M(i, d) \frac{Q_i}{Q_A Q_B} \right]_{T_{\text{trans}}} \quad (3.93)$$

where the total partition functions Q must include the formation energy contribution $\exp(-h_f/k_B T_{\text{trans}})$ to account for the dissociation potential of the lower state, and the rate controlling temperature for the de-excitation process is the translational temperature T_{trans} .

Electron impact dissociation

Similarly as for heavy particle impact dissociation, electron impact dissociation reactions have the following form:



where A and B are the constituent atoms of the molecule AB. The net population rate of level i due to electron impact dissociation (EID) reactions is:

$$\left(\frac{dN_i}{dt} \right)_{\text{EID}} = K_e(d, i) N_A N_B N_e - K_e(i, d) N_i N_e \quad (3.95)$$

where c denotes the ionised state. The forward rate coefficients for diatomic electron impact dissociation are obtained from the literature in generalised Arrhenius form, where the rate controlling temperature for the dissociation process is the free electron temperature:

$$T_x = T_e \quad (3.96)$$

The electron impact recombination rate coefficient is calculated via the principal of detailed balancing:

$$K_e(d, i) = \left[K_e(i, d) \frac{Q_i}{Q_A Q_B} \right]_{T_e} \quad (3.97)$$

where the total partition functions Q must include the formation energy contribution $\exp(-h_f/k_B T_{\text{trans}})$ to account for the dissociation potential of the lower state, and the rate controlling temperature for the recombination process is assumed to be the free electron temperature T_e .

Electron impact ionisation

Electron impact ionisation reactions have the following form:



where X^+ is the ionised species. The net population rate of level i due to electron impact excitation (EII) reactions is:

$$\left(\frac{dN_i}{dt} \right)_{\text{EII}} = K_e(c, i) N_{\text{ion}} N_e N_e - K_e(i, c) N_i N_e \quad (3.99)$$

where c denotes the ionised states. The ionisation process is assumed to be governed by the free electron temperature T_e , and the electron impact recombination rate coefficient is related to the electron impact ionisation rate coefficient via the principal of detailed balancing:

$$K_e(c, i) = \left[K_e(i, c) \frac{Q_i}{Q_{\text{ion}} Q_e} \right]_{T_e} \quad (3.100)$$

where the total partition functions Q must include the formation energy contribution $\exp(-h_f/k_B T_{\text{trans}})$ to account for the ionisation potential of the lower state, and the rate controlling temperature for the recombination process is assumed to be the free electron temperature T_e .

Electron impact ionisation processes are only considered for atomic species in the present work, as the dissociation and excitation processes are much more significant for diatomic species. Similarly as for electron impact excitation, electron impact ionisation rate coefficients based on experimental measurements or theoretical calculations are preferred in the present work. The ionisation cross sections, however, are typically only provided for low lying states and we must rely on empirical models for the remainder. Two empirical models for electron impact ionisation are considered, both based on the hydrogenic cross sections of Drawin [39]. The model implemented by Johnston [12] gives the electron impact ionisation reaction rate coefficients as:

$$K_e(i, c) = 1.46 \times 10^{-10} \sqrt{T_e} \left(\frac{I_H}{I - E_i} \right)^2 \zeta y \psi_1(y) \quad (3.101)$$

where the number of equivalent electrons ζ is 3 for ground electronic states and 1 for excited states and y is the reduced energy of the incoming electrons:

$$y = \frac{I - E_i}{k_B T_e}, \quad (3.102)$$

and the function ψ_1 is:

$$\psi_1(y) = \frac{\exp(-y)}{1 + y} \left\{ \frac{1}{20 + y} + \ln \left[1.25 \left(1 + \frac{1}{y} \right) \right] \right\} \quad (3.103)$$

Panesi [38, 44] implemented another model for electron impact ionisation that is also based on the Drawin cross sections. For this model the ionisation rate coefficients are calculated by the optically allowed expression presented in Equation 3.84 with $\alpha = 1$ and $a = (I - E_i)/(k_B T_e)$.

Bound-bound radiative transitions

Bound-bound radiative transitions have the following form:

$$X_j \rightleftharpoons X_i + h\nu \quad (3.104)$$

where the emitted photon energy is equivalent to the energy difference between the two levels:

$$h\nu = E_j - E_i \quad (3.105)$$

The net population rate of level i due to bound-bound radiative transitions (BBRT) is:

$$\left(\frac{dN_i}{dt} \right)_{\text{BBRT}} = \sum_{j>i} \Lambda_{j,i} A(j, i) N_j - \sum_{j<i} \Lambda_{i,j} A(i, j) N_i \quad (3.106)$$

where $A(i, j)$ is the spontaneous transition probability from level i to level j and $\Lambda_{i,j}$ is the associated escape factor. For atomic radiators, the total transition probability between

two nonequilibrium electronic levels is calculated by averaging the degeneracy weighted transition probabilities for all transitions between the two levels:

$$A(i, j) = \frac{\sum_{i'} \sum_{j'} g'_i A(i', j')}{\sum_{i'} \sum_{j'} g'_i} \quad (3.107)$$

where i' and j' denote sub-levels belonging to grouped-levels i and j respectively. For diatomic radiators, the total transition probability between two nonequilibrium electronic levels can be calculated either from radiative lifetimes $\gamma_{i,j}$ where available:

$$A(i, j) = \frac{1}{\gamma_{i,j}}, \quad (3.108)$$

or by averaging over the weighted vibrational transition probabilities [3]:

$$A(i, j) = \frac{\sum_{v_i} Q_{\text{vib.},v_i} \sum_{v_j} A_{\text{vib.}}(v_i, v_j)}{\sum_{v_i} Q_{\text{vib.},v_i}} \quad (3.109)$$

where $Q_{\text{vib.},v_i}$ is the vibrational partition function for vibrational level v_i and $A_{\text{vib.}}(v_i, v_j)$ is the vibrational transition probability corresponding to the v_i, v_j band. The inverse process of radiative absorption is accounted for by the escape factor Λ that is the ratio of re-absorbed to emitted radiative energy¹²:

$$\Lambda = \frac{E_{\text{abs.}}}{E_{\text{em.}}} \quad (3.110)$$

As radiative re-absorption is determined by solving the radiation transport equations, the introduction of an escape factor here implies that the collisional-radiative and radiative transport equations should be solved in a coupled manner. Unfortunately such a procedure is computationally prohibitive, and therefore the escape factor is approximated in the present work — specifically, solutions with optically thick and optically thin transitions are presented, providing approximate lower and upper bounds for the solution space.

3.2.2 Master equation formulation and solution

The implemented master equation for an electronic level i is:

$$\begin{aligned} \frac{\partial N_i}{\partial t} = & \overbrace{\sum_{j \neq i} K_e(j, i) N_j N_e - \sum_{j \neq i} K_e(i, j) N_i N_e}^{\text{electron impact excitation}} \\ & + \overbrace{\sum_{j \neq i} K_M(j, i) N_j N_M - \sum_{j \neq i} K_M(i, j) N_i N_M}^{\text{heavy particle impact excitation}} \\ & + \overbrace{K_e(c, i) N_+ N_e^2 - K_e(i, c) N_i N_e}^{\text{electron impact ionisation}} + \overbrace{K_e(d, i) N_{XY} N_e - K_e(i, d) N_i N_e}^{\text{electron impact dissociation}} \\ & + \overbrace{\sum_{j > i} \Lambda_{j,i} A(j, i) N_j - \sum_{j < i} \Lambda_{i,j} A(i, j) N_i}^{\text{radiative excitation}}, \end{aligned} \quad (3.111)$$

¹²An escape factor of one represents an optically thin transition where no re-absorption occurs, while an escape factor of zero represents an optically thick transition where complete re-absorption occurs.

As quenching or exchange reactions amongst different species are not considered in the present work, a linear QSS system can be formulated independently for each nonequilibrium radiator. For a radiator with N_{levels}^* nonequilibrium electronic levels¹³, the QSS system is formed by considering the master equations (see Equation 3.111) for non-ground states in the QSS limit (see Equation 3.67) with closure provided by the total population of the radiator (see Equation 3.68). The resulting system can be expressed in matrix form as:

$$\mathbf{M}\vec{x} = \vec{b} \quad (3.112)$$

where \mathbf{M} is a square matrix of dimension N_{levels}^* and \vec{x} and \vec{b} are vectors of dimension N_{levels}^* . The elements of \mathbf{M} are:

$$\begin{aligned} M(i=1, j) &= 1 + \sum_{j^*}^{N_{\text{eqs}}} f_{\text{eq}, j^*} \\ M(i \neq 1, j=i) &= - \left[\sum_{k \neq i} (K_M(i, k)N_M + K_e(i, k)N_e) + K_e(i, c)N_e \right. \\ &\quad \left. + K_M(i, d)N_M + K_e(i, d)N_e + \sum_{k < i} (\Lambda_{i, k}A(i, k)) \right] \\ M(i \neq 1, j \neq i) &= \begin{cases} K_e(j, i)N_e + K_M(j, i)N_M & \text{for } j < i \\ K_e(j, i)N_e + K_M(j, i)N_M + \Lambda_{j, i}A(j, i) & \text{for } j > i \end{cases} \end{aligned} \quad (3.113)$$

and the elements of \vec{b} are:

$$\begin{aligned} \vec{b}(i=1) &= N_s \\ \vec{b}(i \neq 1) &= -K_e(i, c)N_e - K_M(i, d)N_M - K_e(i, d)N_e - \Lambda_{c, i}A(c, i)N_+N_e \end{aligned} \quad (3.114)$$

where N_s is the total species number density from the CFD solver. The elements of \vec{x} are the nonequilibrium level number densities:

$$\vec{x}(i) = N_i \quad (3.115)$$

The matrix line corresponding to the ground state ($i = 1$) is the population summation from Equation 3.68, while the remaining matrix lines are each of the master equations from Equation 3.111. As mentioned previously, the nonequilibrium levels may be a sub-set of the electronic levels for the species. A level not considered by the collisional-radiative model with index j^* can be equilibrated with a nonequilibrium level j by the Boltzmann equation:

$$N_{j^*} = N_j \frac{Q_{\text{int}, j^*}}{Q_{\text{int}, j}} = N_j f_{\text{eq}, j^*} \quad (3.116)$$

¹³As will be discussed, the number of nonequilibrium electronic levels N_{levels}^* is not necessarily the total number of electronic levels N_{levels}

where Q_{int,j^*} and $Q_{\text{int},j}$ are the internal partition functions of the equilibrated and nonequilibrium levels respectively, and the ratio of these two partition functions is defined as the Boltzmann equilibrium factor f_{eq,j^*} . To correctly account for the level equilibration in the collisional-radiative model, the $M(i = 1, j)$ matrix elements must include the sum of all Boltzmann equilibrium factors for nonequilibrium level j :

$$\sum_{j^*}^{N_{\text{eqs}}} f_{\text{eq},j^*} \quad (3.117)$$

Equation 3.112 can then be easily solved via direct matrix inversion:

$$\vec{x} = M^{-1}\vec{b} \quad (3.118)$$

where M^{-1} is calculated via Gaussian elimination in the current implementation of the model. At low temperatures, however, there are insufficient collisions for the QSS condition (Equation 3.67) to remain valid, and the results of Equation 3.118 cannot be used. Therefore in the present work the electronic levels of nonequilibrium radiators are assumed to be in Boltzmann distributions for free electron temperatures of 2,000 K and under.

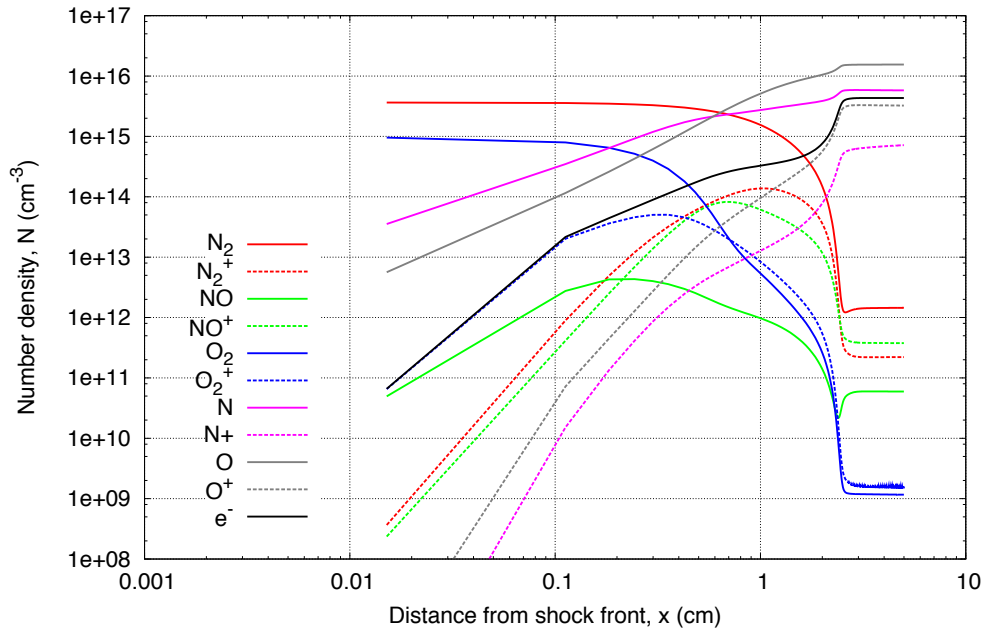
3.2.3 Collisional-radiative model for N₂–O₂ mixtures

Before deciding upon an appropriate collisional-radiative model for N₂–O₂ mixtures, it is instructive to consider a typical Earth re-entry shock layer. Figures 3.12a and 3.12b present post-shock species number density and radiative emission profiles respectively for the Fire II $t = 1634$ s condition. For this analysis, the electronic states of the radiators are assumed to be populated by Boltzmann distributions. Immediately behind the shock, O₂ rapidly dissociates and quickly forms a large population of O atoms, while N₂ dissociation proceeds at a slightly slower rate, leading to significant N₂ and N₂⁺ radiation up to 2 cm behind the shock. The radiative emission of N₂ and N₂⁺, however, is quickly exceeded by the lines of N and O as the heavily dissociated and partially ionised equilibrium state is approached. As bound-bound transitions of NO, O₂, N⁺ and O⁺ only make minor contributions to the radiative emission, it is sufficient to consider the electronic levels of these species as being populated by Boltzmann distributions. Conversely the radiative emission from N₂, N₂⁺, N and O bound-bound transitions are significant, and the electronic levels of these species should be calculated via collision-radiative modelling. Furthermore, as the free electron number density is almost the same order of magnitude as that of the heavy particles, reactions due to heavy particle impact can be omitted.

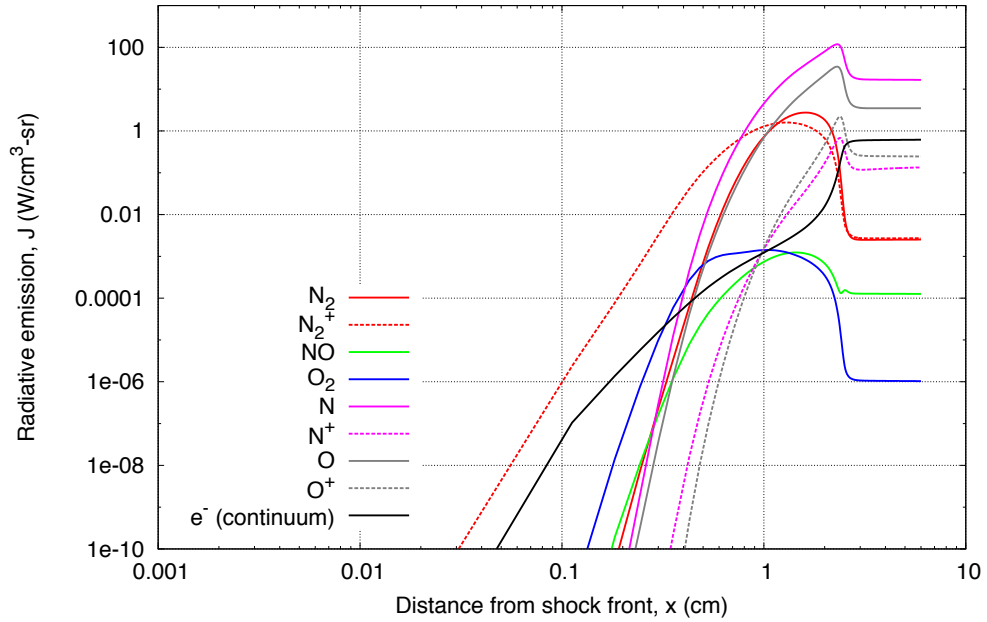
(a) Atomic species: N and O

The collisional processes considered for the atomic species N and O are electron impact excitation and ionisation, and the radiative processes considered are bound-bound optically allowed transitions. Table 3.11 summarises the implemented rate coefficients for each of these mechanisms. Where more than one model are presented for a mechanism, they are listed in order of preference (*e.g.* for the electron impact excitation of N, the rates of Frost *et al.* [45] are preferred with the remaining transitions described by the semi-empirical model of Gryzinski [40]).

For the radiative transitions, the transition probabilities $A(i, j)$ for the nonequilibrium levels are calculated using Equation 3.107 where the individual line transition probabilities are obtained from the NIST Atomic Species Database [10] (see Table 3.2). For the electron impact transitions, the rate coefficients $K_e(i, j)$ and $K_e(i, c)$ are either calculated



(a) Number density profile



(b) Radiative emission profile

Figure 3.12: Post-shock species number density and radiative emission profiles along the stagnation streamline of the Fire II $t = 1634$ s condition ($p_\infty = 2$ Pa, $T_\infty = 195$ K, $u_\infty = 11,360$ m/s).

using semi-empirical models or obtained directly from the literature in the form of curve-fits.

Although all electron impact excitation and ionisation rates for N and O are able to be calculated by the previously-described semi-empirical models, the hydrogenic assumptions of these models are not appropriate for transition originating from the inner core of electronic levels [11]. Rates derived from experimental measurements of theoretical calculations are therefore preferred for transitions originating from the ground and low lying metastable states. Fortunately, experimental measurements and quantum mechanical calculations of these transitions are much simpler and more readily available than for the high lying levels.

Table 3.11: Summary of the collisional-radiative mechanisms implemented for N and O.

Species	Electronic levels	CR mechanisms	Models
N	All	Electron impact excitation	(a) Frost <i>et al.</i> [45] (b) Gryzinski [40]
		Electron impact ionisation	(a) Soon and Kunc [46] (b) Drawin (Reference [38])
		Radiative decay	NIST Atomic Spectra Database [10]
O	All	Electron impact excitation	(a) Zatsarinny and Tayal [47] (b) Gryzinski [40]
		Electron impact ionisation	(a) Soon and Kunc [46] (b) Drawin (Reference [38])
		Radiative decay	NIST Atomic Spectra Database [10]

(a) Electron impact excitation of N

Frost *et al.* [45] performed R-matrix calculations for N and N⁺ electron impact excitation transitions from the first 3 energy levels to all levels with principle quantum number n less than 3. Panesi [48] demonstrated improved agreement with the EAST shock tube data when implementing this model for N. The rate coefficient is given as a function of the effective collision strength $\gamma_{i,j}$:

$$K_e(i, j) = 2\sqrt{\pi}\alpha c a_0^2 \sqrt{\frac{E_H}{kT_e}} \frac{\gamma_{i,j}(T_e)}{g_i} \exp\left(-\frac{\Delta E_{i,j}}{kT_e}\right), \quad (3.119)$$

where α is the fine structure constant and the effective collision strength $\gamma_{i,j}$ has been curve fitted against the tabulated values provided by Frost in the range $0.5 \leq T_e \leq 12.0$ eV.

Bultel *et al.* [49] presented electron impact excitation and ionisation rates for the ground and metastable states of N and O. These rates were presented as generalised Arrhenius curve-fits in the temperature range $2,000 \leq T \leq 10,000$ K and were implemented in the collisional-radiative model of Panesi [38]. The excitation rates from the ground state of nitrogen are based on the R-matrix calculations of Berrington [50].

Figure 3.12 compares the electron impact excitation rate coefficient for a selection of optically allowed and optically forbidden transitions of atomic nitrogen for which Frost *et al.* [45] present rate coefficients. The indices of the initial i and final j electronic levels are given in the y-axis label in the form $K_e(i, j)$. For the ground to first and second excited level transitions, Figures 3.13a and 3.13b respectively, the R-matrix based rates of Bultel and Frost are bounded by the semi-empirical Gryzinski and Drawin models. The Bultel rates are up to an order of magnitude lower than the Frost rates for the

$2s^2 2p^3 \ ^4S^\circ \Rightarrow 2s^2 2p^3 \ ^2D^\circ$ optically forbidden transition. For these transitions the theoretical calculations of Frost *et al.* [45] are preferred as they are more recent than those of Berrington [50].

The Frost, Drawin and Gryzinski models exhibit qualitative agreement for the remaining transitions shown, Figures 3.13c to 3.13h. Quantitatively, it is encouraging to observe that the data of Frost is bounded by the Gryzinski and Drawin models for almost all transitions, although there is no trend as to which forms the upper or lower bound. For the 1-5 and 3-15 transitions in Figures 3.13c and 3.13f, for example, the Gryzinski data shows exceptional agreement with the calculations of Frost. In contrast, for the 1-20 and 3-20 transitions in Figures 3.13c and 3.13h, the Gryzinski model substantially underestimates the data of Frost while the Drawin model shows good agreement. Furthermore the semi-empirical models differ by up to two orders of magnitude for some transitions. In the present work the theoretical calculations of Frost *et al.* [45] are preferred where available, with the remaining transitions described by the Gryzinski [40] model. The decision to implement the Gryzinski model in preference to the Drawin model is based on the findings of Panesi [38, 48], where the Gryzinski model gave improved agreement with the air shock tube spectroscopy experiments performed in the EAST facility.

(b) Electron impact ionisation of N

Johnston [12] implemented the ionisation rate coefficients proposed by Kunc and Soon [51] for the ground and first two excited levels of atomic nitrogen. The rate coefficients are calculated as:

$$K_e(i, c) = 1.0 \times 10^{-8} \left[\frac{I_H}{I - E_i} \right] \frac{Q_i}{2l_i + 1} \exp(-\beta) G_i(\beta) \quad (3.120)$$

where,

$$G_i(\beta) = \sqrt{\frac{\beta}{\beta + 1} \frac{A}{\beta + \chi}}, \quad (3.121)$$

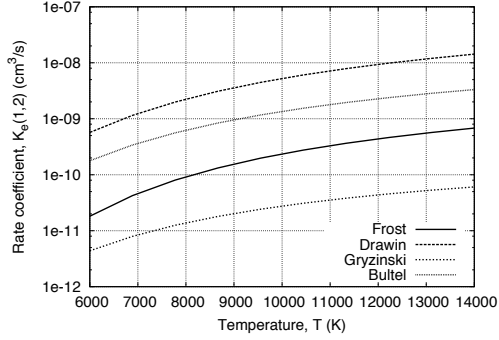
and,

$$\beta = \frac{I - E_i}{kT_e}. \quad (3.122)$$

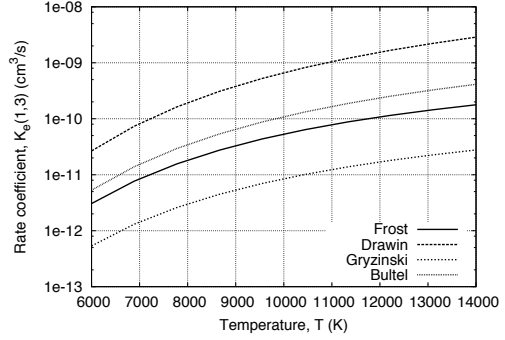
The parameter I_H is the ionisation energy of the hydrogen energy (Rydberg energy), l_i is the angular momentum quantum number of the level i , A and χ are fitting constants for the species and G_i is level dependent angular factor. For atomic nitrogen A is equal to 27.71, χ is equal to 5.58 and Q_i is equal to 3 for the ground state and 3/2 for the first and second excited states. This expression is a curve-fit based on the rate coefficient derived from experimentally measured electron impact ionisation cross sections for the ground state of N.

Panesi [38] implemented the ionisation rate coefficients presented by Bultel *et al.* [49], which were obtained from the compilation of Tawara and Kato [52] and the combined binary-encounter Bethe (BEB) and scaled plane-wave Born (PWB) calculations of Kim and Desclaux [53]. The rates for the ground and first two excited states of N were presented as generalised Arrhenius curve-fits in the temperature range $2,000 \leq T_e \leq 10,000$ K.

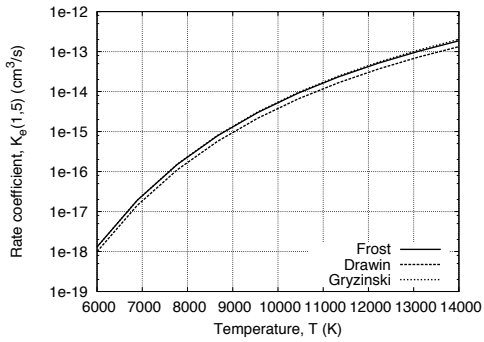
Figure 3.13 compares the electron impact ionisation rate coefficient for various transitions of N. The experimentally-fitted Kunc and Soon rates and the theoretically-fitted



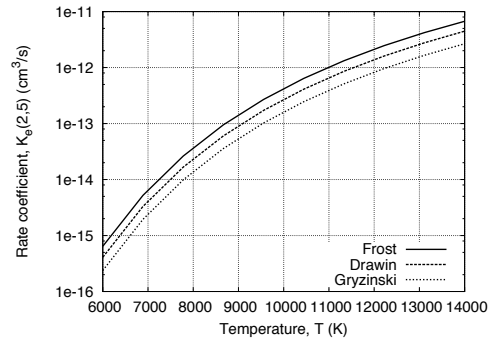
(a) $2s^2 2p^3 4S^{\circ} \Rightarrow 2s^2 2p^3 2D^{\circ}$ (forbidden)



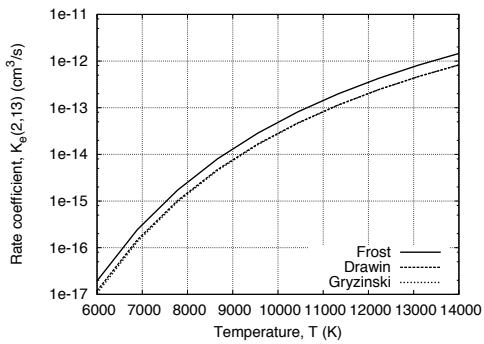
(b) $2s^2 2p^3 4S^{\circ} \Rightarrow 2s^2 2p^3 2P^{\circ}$ (forbidden)



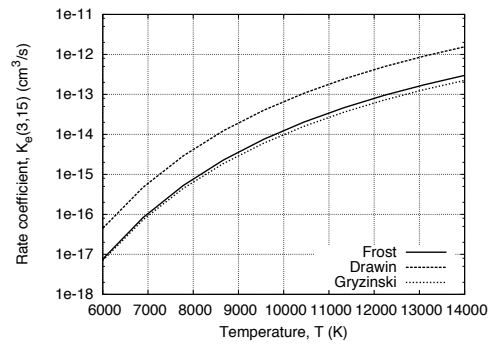
(c) $2s^2 2p^3 4S^{\circ} \Rightarrow 2s^2 2p^2(^3P)3s^4P$ (forbidden)



(d) $2s^2 2p^3 2D^{\circ} \Rightarrow 2s^2 2p^2(^3P)3s^4P$ (allowed)

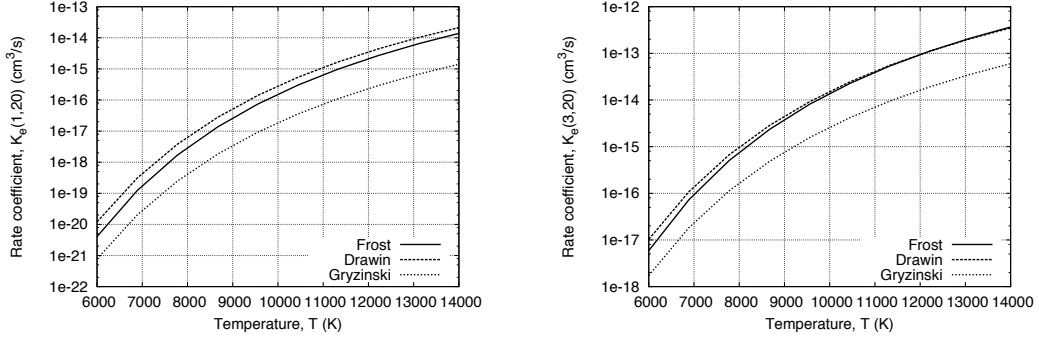


(e) $2s^2 2p^3 2D^{\circ} \Rightarrow 2s^2 2p^2(^1D)3s^2D$ (allowed)



(f) $2s^2 2p^3 2P^{\circ} \Rightarrow 2s^2 2p^2(^3P)4s^2P$ (allowed)

Figure 3.13: Comparison of electron impact excitation rate coefficients for atomic nitrogen.



(g) $2s^2 2p^3 \ ^4S^o \Rightarrow 2s^2 2p^2 (^3P) 3d^4 D$ (forbidden) (h) $2s^2 2p^3 \ ^2P^o \Rightarrow 2s^2 2p^2 (^3P) 3d^4 D$ (forbidden)

Figure 3.12: (Continued) Comparison of electron impact excitation rate coefficients for various transitions of N.

Bultel rates are in good agreement for the ground and metastable states, Figures 3.13a to 3.13c. The Drawin models implemented by Johnston [12] and Panesi [38] bound the Kunc and Soon rates and Bultel rates, with Panesi's implementation being in closer agreement. Although the difference between the two Drawin models decreases as ionisation from higher levels and low electron temperatures is considered (see Figure 3.13d), Johnston's implementation is between approximately 2 and 100 times smaller for all levels. Therefore in the present work the experimentally-fitted electron impact ionisation coefficients of Soon and Kunc [46] are preferred for the first three levels, while the Drawin model implemented by Panesi [12] is used for the remaining levels.

(a) Electron impact excitation of O

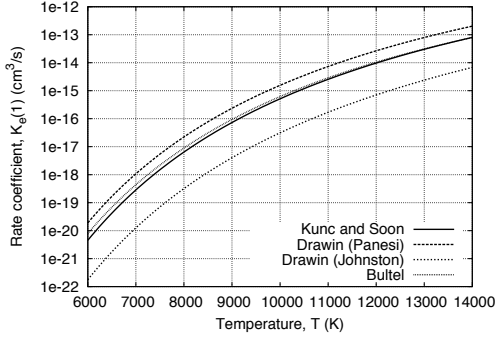
Zatsarinny and Tayal [47] calculated electron impact excitation rates for the ground and first two excited states of atomic oxygen using a *B*-spline *R*-matrix approach. The forward rate coefficients for the Zatsarinny and Tayal model are calculated as:

$$K_e(i, j) = \frac{8.629 \times 10^{-6}}{g_i \sqrt{T_e}} \gamma_{ij}(T_e) \exp\left(\frac{-\Delta E_{ij}}{kT_e}\right) \quad (3.123)$$

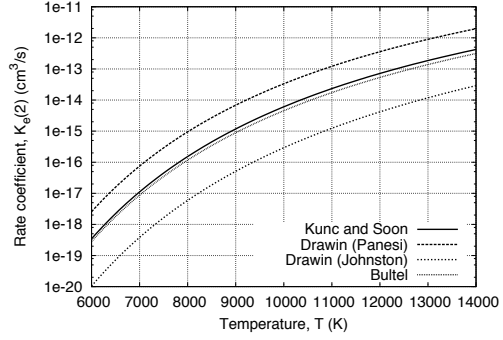
where the dimensionless effective collision strength γ_{ij} is tabulated as a function of the free electron temperature in Reference [47]. The Zatsarinny and Tayal [47] data was the preferred source of accurate atomic oxygen electron impact excitation rates in the collisional-radiative model proposed by Johnston [12].

Panesi [38] implemented the electron impact excitation rate coefficients presented by Bultel *et al.* [49] which are based on the literature survey of Itikawa and Ichimura [54]. Electron impact excitation rate coefficients for transitions from the ground state to the two metastable states of atomic oxygen were presented as generalised Arrhenius curve-fits in the temperature range $2,000 \leq T_e \leq 10,000$ K.

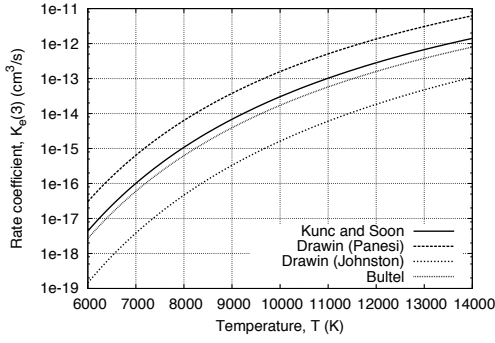
Figure 3.18 compares the electron impact excitation rate coefficient for various transitions of atomic oxygen for which Zatsarinny and Tayal [47] present data. For the transitions to the metastable states, Figures 3.14a and 3.14b, the Drawin and Gryzinski models substantially overestimated the theoretical calculations of Zatsarinny and Tayal. Such a discrepancy is to be expected for these inner core transitions due to the hydrogenic as-



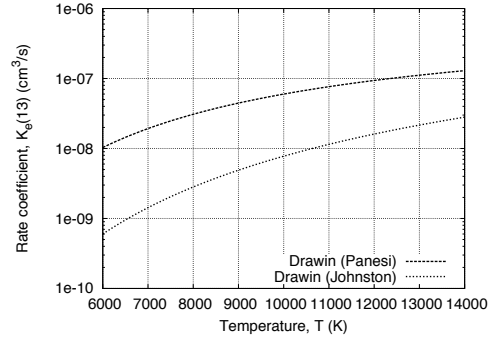
(a) $N 2s^2 2p^3 4S^{\circ} \Rightarrow N^+$



(b) $N 2s^2 2p^3 2D^{\circ} \Rightarrow N^+$



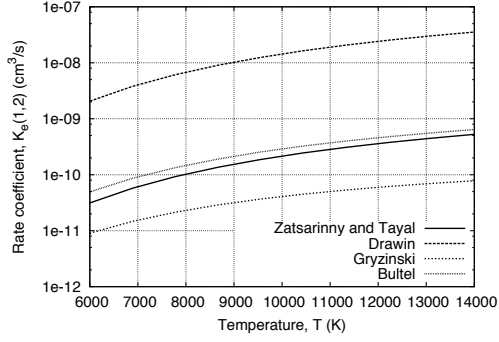
(c) $N 2s^2 2p^3 2P^{\circ} \Rightarrow N^+$



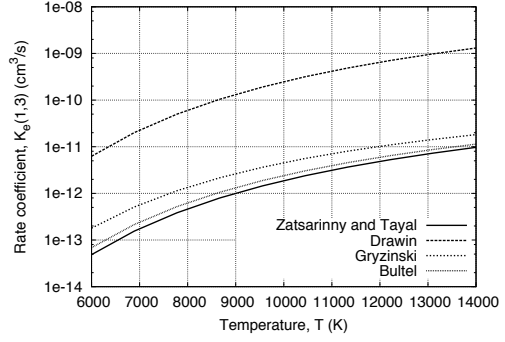
(d) $N 2s^2 2p^2 (1D) 3s^2 D \Rightarrow N^+$

Figure 3.13: Comparison of electron impact ionisation rate coefficients for atomic nitrogen.

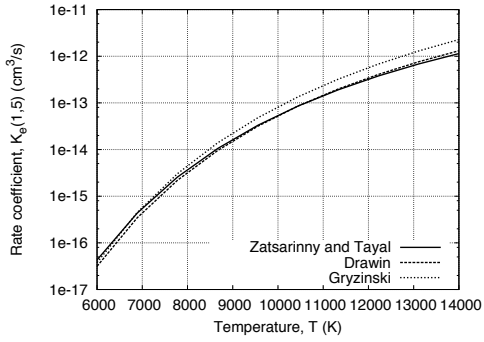
assumptions of the Drawin and Gryzinski models. For the transitions where Bultel rates are available, good agreement is observed with the theoretical calculations of Zatsarinny and Tayal. Although the *B*-spline *R*-matrix calculations of Zatsarinny and Tayal are bounded by the semi-empirical models for most transitions, for some forbidden transitions such as 2–5 and 3–7, Figures 3.14d and 3.14f, the semi-empirical models underestimate the theoretical rates by at least an order of magnitude for the temperature range considered. For some transitions such as 1-16 and 1-17, Figures 3.14g and 3.14h, the Gryzinski model closely matches the Zatsarinny and Tayal calculations, whilst for others such as 1-20, Figure 3.14i, the Drawin model shows exceptional agreement. The Drawin and Gryzinski models show considerable variability in their relative magnitudes, being within a factor of 2 of each other for some transitions and in excess of 10^4 for some transitions to high lying states (e.g. Figures 3.14i and 3.14j). As for atomic nitrogen, we must again bear in mind that Panesi [38, 48] found improved agreement with experiment when using the Gryzinski model. Therefore in the present work the Zatsarinny and Tayal [47] rate coefficients are preferred where available, while the Gryzinski [40] model is applied to the remaining transitions.



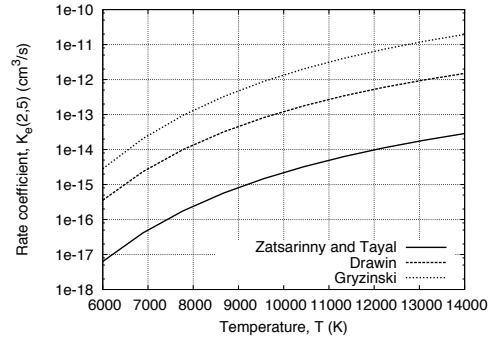
(a) $2s^2 2p^4 \ ^3P \Rightarrow 2s^2 2p^4 \ ^1D$ (Forbidden)



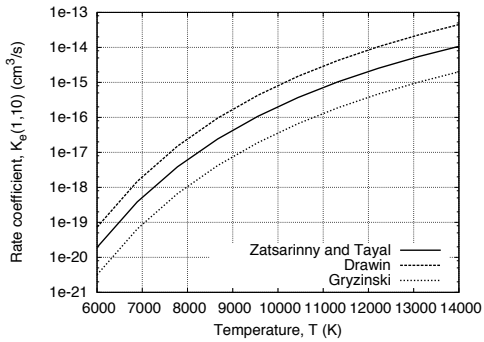
(b) $2s^2 2p^4 \ ^3P \Rightarrow 2s^2 2p^4 \ ^1S$ (Forbidden)



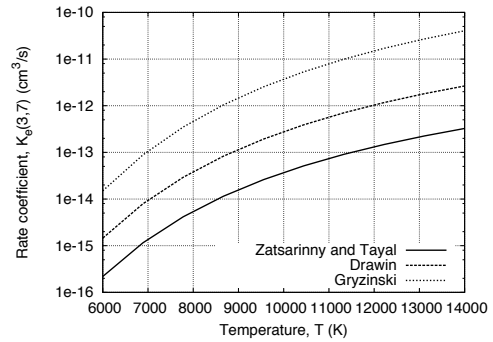
(c) $2s^2 2p^4 \ ^3P \Rightarrow 2s^2 2p^3(^4S^o) 3p^5 p$ (Allowed)



(d) $2s^2 2p^4 \ ^1D \Rightarrow 2s^2 2p^3(^4S^o) 3p^5 p$ (Forbidden)

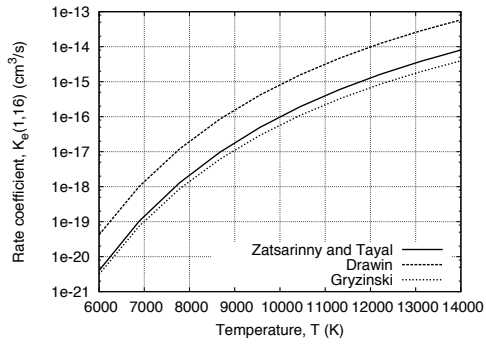


(e) $2s^2 2p^4 \ ^3P \Rightarrow 2s^2 2p^3(^4S^o) 3d^5 D^o$ (Forbidden)

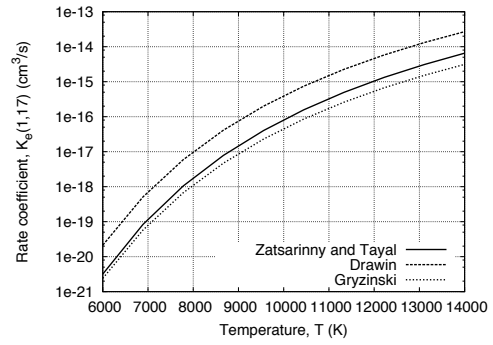


(f) $2s^2 2p^4 \ ^1S \Rightarrow 2s^2 2p^3(^4S^o) 3p^3 S$ (Forbidden)

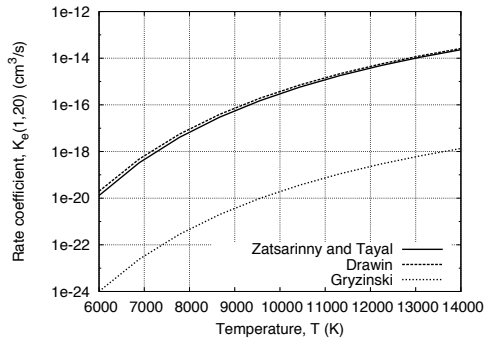
Figure 3.14: Comparison of electron impact excitation rate coefficients for atomic oxygen.



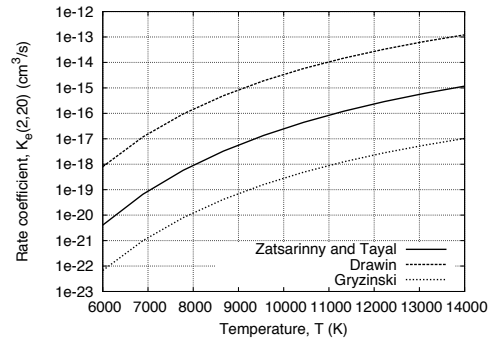
(g) $2s^22p^4\ ^3P \Rightarrow 2s^22p^3(^4S^{\circ})5s\ ^3S^{\circ}$ (Allowed)



(h) $2s^22p^4\ ^3P \Rightarrow 2s^22p^3(^2D^{\circ})3s\ ^1D^{\circ}$ (Forbidden)



(i) $2s^22p^4\ ^3P \Rightarrow 2s^22p^3(^4S^{\circ})4f\ ^5F$ (Forbidden)



(j) $2s^22p^4\ ^1D \Rightarrow 2s^22p^3(^4S^{\circ})4f\ ^5F$ (Forbidden)

Figure 3.13: (Continued) Comparison of electron impact excitation rate coefficients for atomic oxygen.

Electron impact ionisation of O

Johnston [12] implemented the ionisation rate coefficients proposed by Soon and Kunc [46] for the ground and first two excited levels of atomic oxygen. The rate coefficients are calculated as described in Equations 3.120 to 3.122, where $A = 30.52$ and $\chi = 4.0$ for levels $i=1, 2$ and 3 of O. This expression is a curve-fit based on the rate coefficient derived from experimentally measured electron impact ionisation cross sections for the ground state of O.

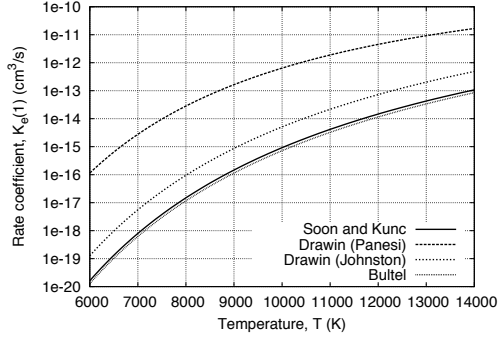
Panesi [38] implemented the ionisation rate coefficients presented by Bultel *et al.* [49], which were obtained from the compilation of Tawara and Kato [52] and the combined binary-encounter Bethe (BEB) and scaled plane-wave Born (PWB) calculations of Kim and Desclaux [53]. The rates for the ground and first two excited states of O were presented as generalised Arrhenius curve-fits in the temperature range $2,000 \leq T_e \leq 10,000$ K.

Figure 3.14 compares the electron impact ionisation rate coefficient for various transitions of O. The experimentally-fitted Soon and Kunc models and theoretically-fitted Bultel models give excellent agreement for ionisation from the ground state, but exhibit up to an order of magnitude deviation for the $2s^22p^4\ ^1S$ multiplet (see Figure 3.14c). The two Drawin models bound the experimentally fitted model of Kunc and Soon [51], with Panesi's implementation being in closer agreement, especially for ionisation of the $2s^22p^4\ ^1S$ multiplet. Similarly as for atomic nitrogen, Johnston's implementation of the Drawin cross sections is between approximately 2 and 100 times smaller than Panesi's implementation for all levels. Therefore in the present work the electron impact ionisation coefficients of Kunc and Soon [51] are preferred for the first three levels, while the Drawin model implemented by Panesi [12] is used for the remaining levels.

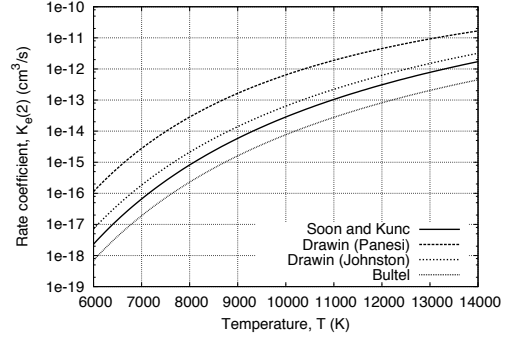
(b) Diatomic species: N_2 and N_2^+

Johnston [12] presented collisional-radiative models for N_2 and N_2^+ compiled from both theoretically calculated and experimentally measured rate-coefficients in the literature. The majority of the collisional rate coefficients are based on the theoretical calculations of Teulet *et al.* [55], however other more accurate data was preferred where available. Since Johnston formulated this model, a set of collisional rate coefficients for the diatomic species CN, CO, N_2 , N_2^+ , O_2 and NO have been proposed by Park [56, 57]. The rate coefficients are based on experimentally measured cross sections where available, and theoretically estimated otherwise.

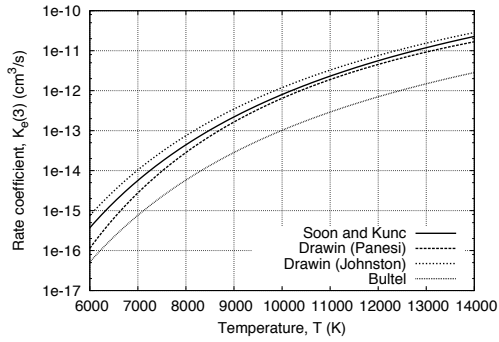
The most critical reactions for N_2 and N_2^+ at Earth re-entry conditions are those populating the upper states of radiative transitions via electron impact excitation. Figure 3.15 compares the electron impact excitation rates populating the $B^3\Pi_g$ state of N_2 (upper state for the First Positive band system), and Figure 3.16 compares the electron impact excitation rates populating the $B^2\Sigma_u^+$ state of N_2^+ (upper state for the First Negative band system). While the two models agree to within a factor of 4 for the $N_2(X^1\Sigma_g^+) + e^- \rightleftharpoons N_2(B^3\Pi_g) + e^-$ transition, Figure 3.15a, the Park rates are substantially higher than those of Johnston for the other transitions. For the $N_2(A^2\Pi_u) + e^- \rightleftharpoons N_2(B^2\Sigma_u^+) + e^-$ transition in Figure 3.16b, for example, the Park rate is almost four orders of magnitude greater than the Johnston rate. Although not shown here, the estimated rate of Teulet *et al.* [55] for the $N_2^+(X^2\Sigma_g^+) + e^- \rightleftharpoons N_2^+(B^2\Sigma_u^+) + e^-$ transition is quite similar to the Park rate. In Johnston's [12] survey of the literature, this rate of Teulet was found to overestimate those from more accurate theoretical calculations. Furthermore, in Reference [58] the rates of Teulet were required to be reduced by factors of 10 and 70 for N_2 and N_2^+ respectively in order to achieve agreement with experiment. In the present work therefore



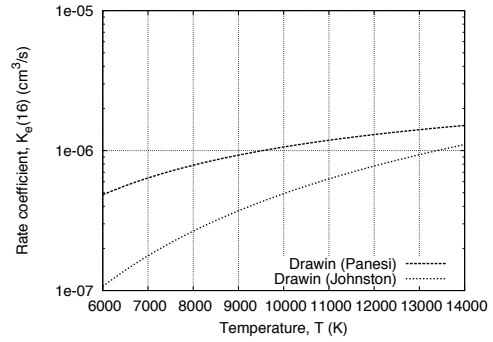
(a) $O 2s^2 2p^4 3P \Rightarrow O^+$



(b) $O 2s^2 2p^4 1D \Rightarrow O^+$



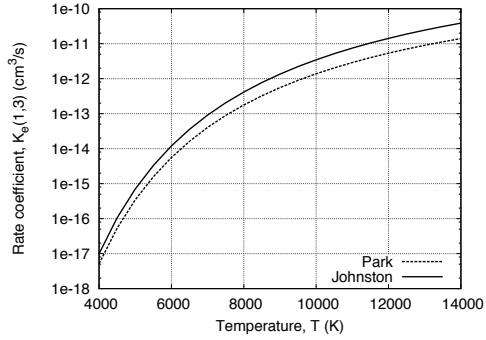
(c) $O 2s^2 2p^4 1S \Rightarrow O^+$



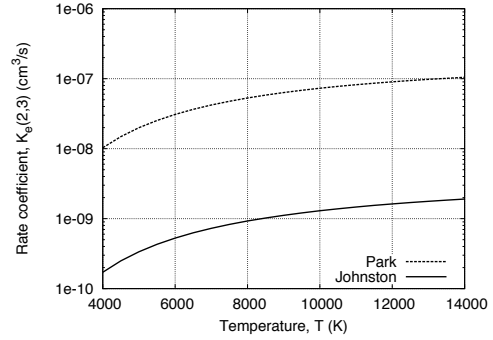
(d) $O 2s^2 2p^3 (4S^o) 5s 3S^o \Rightarrow O^+$

Figure 3.14: Comparison of electron impact ionisation rate coefficients for atomic oxygen.

we choose to adopt the Johnston [12] model. The collisional-radiative models for N_2 and N_2^+ are summarised in § A.4 and A.5 respectively. The nonequilibrium levels considered for N_2 are $X^1\Sigma_g^+$, $A^3\Sigma_u^+$, $B^3\Pi_g$ and $C^3\Pi_u$, while those for N_2^+ are $X^2\Sigma_g^+$, $A^2\Pi_u$, $B^2\Sigma_u^+$ and $C^2\Sigma_u^+$.

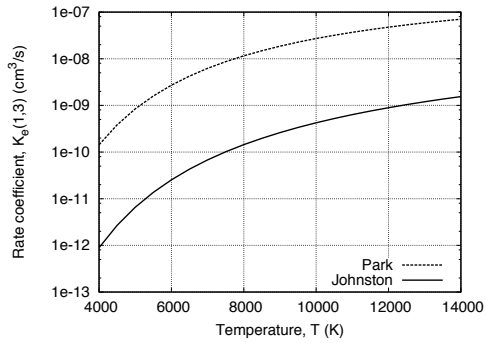


(a) $X^1\Sigma_g^+ \Rightarrow B^3\Pi_g$

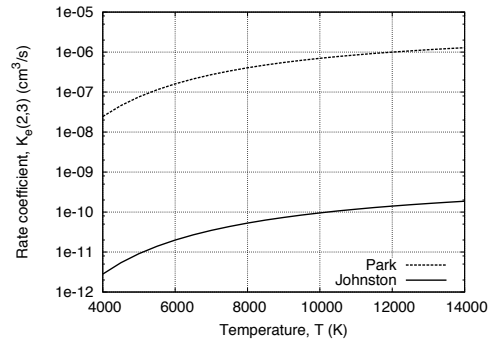


(b) $A^3\Sigma_u^+ \Rightarrow B^3\Pi_g$

Figure 3.15: Comparison of electron impact excitation rate coefficients for transitions to the $B^3\Pi_g$ state of N_2 .



(a) $X^2\Sigma_g^+ \Rightarrow B^2\Sigma_u^+$

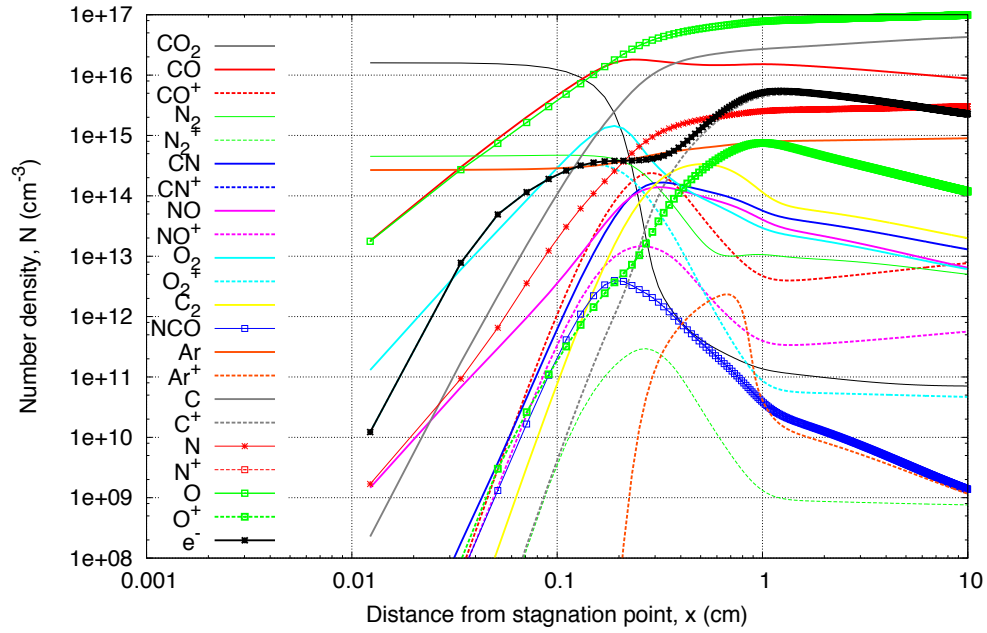


(b) $A^2\Pi_u \Rightarrow B^2\Sigma_u^+$

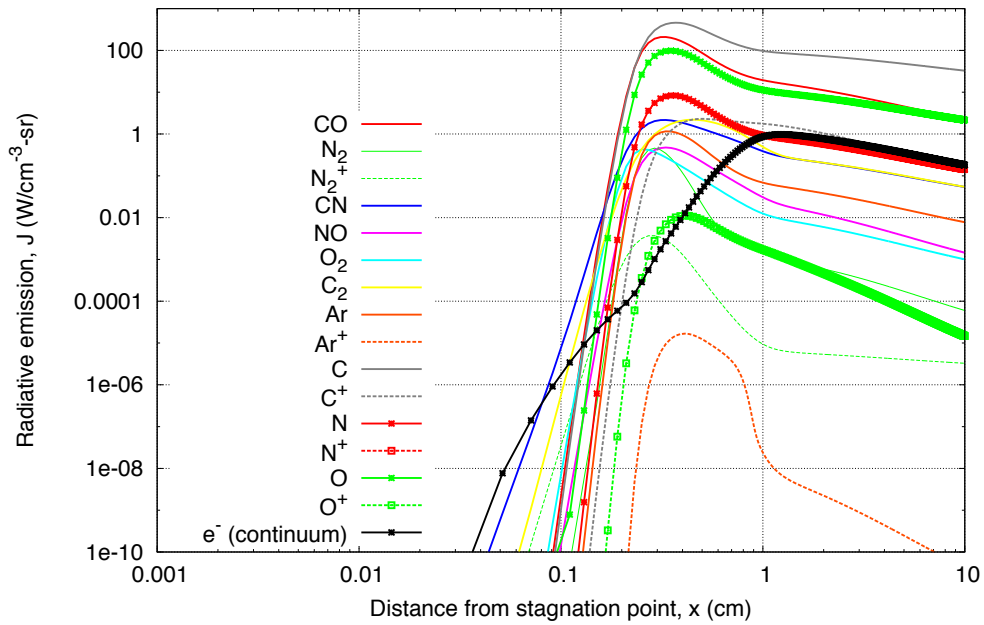
Figure 3.16: Comparison of electron impact excitation rate coefficients for transitions to the $B^2\Sigma_u^+$ state of N_2^+ .

3.2.4 Collisional-radiative model for CO_2 – N_2 – Ar mixtures

Figures 3.17a and 3.17b present post-shock species number density and radiative emission profiles respectively for a hypothetical high-speed Mars aerocapture condition from the trajectory study of Braun *et al.* [59]. The condition corresponds to that predicted for an aerocapture vehicle with nose radius 10 m at 44.9 km altitude that entered the Martian atmosphere at 9.79 km/s. This point is just prior to peak heating and is characterised by very strong thermochemical nonequilibrium. Behind the shock CO_2 and N_2 quickly dissociate, forming an initial pool of CO and N_2 molecules and C , N and O atoms that allow exchange reactions to begin. At 1 mm behind the shock reactions have only just begun to occur and CN and C_2 systems and continuum transitions dominate the radiative emission. By 3 mm behind the shock dissociation is essentially complete and peak emission occurs, with atomic C lines, the VUV CO band systems and atomic O lines contributing 99% of the radiative emission. At peak emission the next strongest radiators are N , CN , C_2 and Ar , however their emission strength is on average two orders of magnitude



(a) Number density profile



(b) Radiative emission profile

Figure 3.17: Post-shock species number density and radiative emission profiles along the stagnation streamline for a hypothetical Mars aerocapture entry condition ($p_\infty = 6.2$ Pa, $T_\infty = 161$ K, $u_\infty = 9,440$ m/s).

less than CO and C. As equilibrium is approached the total radiation drops by an order of magnitude, C, CO and O (in that order) continue to dominate the radiative emission while the relative strength of the continuum transitions increases due to the growing free electron population. The free electron mole-fraction between peak-emission and chemical equilibrium is in the order of 10^{-2} .

Although C, CO and O are by far the strongest radiators when Boltzmann level populations is assumed, the nonequilibrium emission is likely to be significantly less due to radiative depletion of the high lying states. Therefore in the present work we chose to apply the collisional-radiative model to all the significant radiators – C, CO, O, C₂, N, Ar and CN. Although the free electron number density is lower than for Earth re-entry, electron impact collisional processes should still dominate over heavy particle collisions due to their high efficiency. Thus in the present work only electron impact collisional processes will be considered.

Atomic species: Ar, C, N and O

Table 3.12 summarises the implemented rate coefficients for the collisional-radiative mechanisms considered for atomic species in CO₂-N₂-Ar mixtures. Similarly as for N₂-O₂ mixtures, the collisional mechanisms considered for atoms are electron impact excitation and ionisation, and the radiative processes considered are bound-bound optically allowed transitions. As the Drawin [39] cross sections have been shown to be appropriate for calculating non-Boltzmann emission from atomic argon [60,61], the Drawin model proposed by Panesi [38] is applied to calculate electron impact excitation and ionisation of Ar. For atomic carbon, accurate rate-coefficients are obtained from Suno and Kato [62] for electron impact excitation and ionisation of the ground and metastable states. The remaining transitions described by the semi-empirical model of Gryzinski [40]. The collisional-radiative models selected for N and O in N₂-O₂ mixtures are retained here for application to CO₂-N₂-Ar mixtures. All radiative transition probabilities $A(i, j)$ are calculated using Equation 3.107 where the individual line transition probabilities are obtained from the NIST Atomic Species Database [10] (see Table 3.2).

(a) Electron impact excitation of C

Suno and Kato [62] compiled an extensive electron-impact ionisation, excitation and charge exchange cross section database for carbon atoms and ions. Although the database extends to electron temperatures of 1 keV as it is designed nuclear fusion applications, the presented electron impact excitation data for C is shown to agree well with the experimental data of Duneath *et al.* [63] and others obtained in the 1 eV electron temperature range of present interest. The data are presented as curve fits for the collision strength Ω_{ij} from which the rate coefficient can be calculated as:

$$K_e(i, j) = \frac{8.010 \times 10^{-8}}{\omega_i \sqrt{T_e}} y \int_1^\infty \Omega_{ij} e^{-yX} dX \quad (3.124)$$

where T_e is in eV, $y = \Delta E_{ij}/T_e$ and $X = E_e/\Delta E_{ij}$.

Figure 3.17 compares the electron impact excitation rate coefficient for various transitions of C. For all transitions except excitation of the ground to first excited state (see Figure 3.18a), the Suno and Kato rates are between 2 and 1000 times larger than the rates predicted by the approximate models. This is in contrast to the electron impact excitation rates of N and O, where the accurate rates for low lying levels was bounded by the approximate models. It is therefore possible that the rates of Suno and Kato are not suitable for the temperature range considered. In the absence of better electron impact excitation

Table 3.12: Summary of the collisional-radiative mechanisms implemented for Ar, C, N and O in CO₂-N₂-Ar mixtures.

Species	Electronic levels	CR mechanisms	Models
Ar	All	Electron impact excitation	Drawin (Reference [38])
		Electron impact ionisation	Drawin (Reference [38])
C	All	Radiative decay	NIST Atomic Spectra Database [10]
		Electron impact excitation	(a) Suno and Kato [62] (b) Gryzinski [40]
		Electron impact ionisation	(a) Suno and Kato [62] (b) Drawin (Reference [38])
		Radiative decay	NIST Atomic Spectra Database [10]
N	All	Electron impact excitation	(a) Frost <i>et al.</i> [45] (b) Gryzinski [40]
		Electron impact ionisation	(a) Soon and Kunc [46] (b) Drawin (Reference [38])
		Radiative decay	NIST Atomic Spectra Database [10]
O	All	Electron impact excitation	(a) Zatsarinny and Tayal [47] (b) Gryzinski [40]
		Electron impact ionisation	(a) Soon and Kunc [46] (b) Drawin (Reference [38])
		Radiative decay	NIST Atomic Spectra Database [10]
		Radiative decay	NIST Atomic Spectra Database [10]

data, however, the Suno and Kato rates are preferred for transitions from low lying states while the Gryzinski model is applied to the remaining transitions.

(b) Electron impact ionisation of C

Suno and Kato [62] present the electron impact ionisation cross section for the ground state of atomic carbon in the following form:

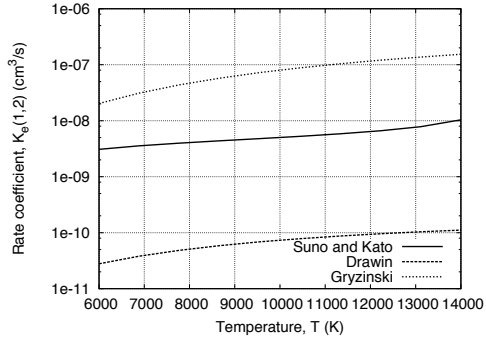
$$\sigma = \frac{1 \times 10^{-13}}{IE} \left\{ A_1 \ln(E/I) + \sum_{j=2}^{N_A} A_j \left(1 - \frac{I}{E}\right)^{j-1} \right\} \quad (3.125)$$

where I and E are the ionisation and electron energy in eV and A_j are fitting coefficients. The fitting coefficients have been selected to match the experimental measurements of Brook *et al.* [64]. Unfortunately Brook considered electron energies in the range $7 \leq E \leq 1000$ eV, which is outside the $E \approx 1$ eV range of present interest. In the absence of electron impact ionisation cross sections for the metastable state of atomic carbon, in the present work Equation 3.125 is applied where I is taken as the level-specific ionisation potential. The rate coefficient is then calculated as:

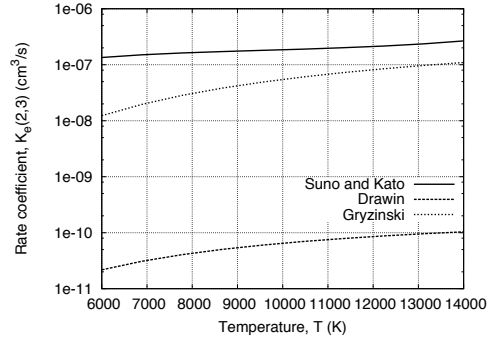
$$K_e(i, c) = \frac{8.010 \times 10^{-8}}{g_i \sqrt{T_e}} y \int_1^\infty \Omega_i e^{-yX} dX \quad (3.126)$$

where T_e is in eV, $y = I/T_e$, $X = E_e/I$ and collision strength Ω_i for level i is obtained from the cross section in Equation 3.125 via:

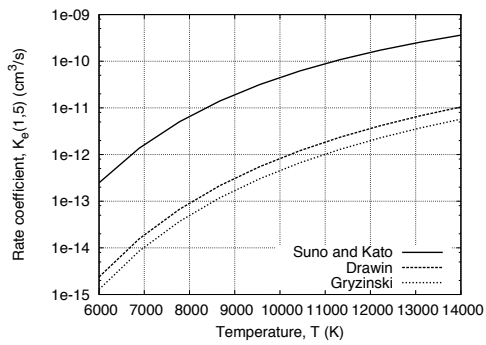
$$\Omega_i = \sigma_i \frac{g_i E}{1.1969 \times 10^{-15}} \quad (3.127)$$



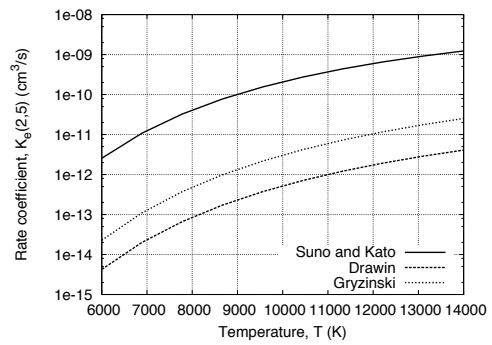
(a) $2s^22p^2P^o \Rightarrow 2s2p^2^4P$ (Forbidden)



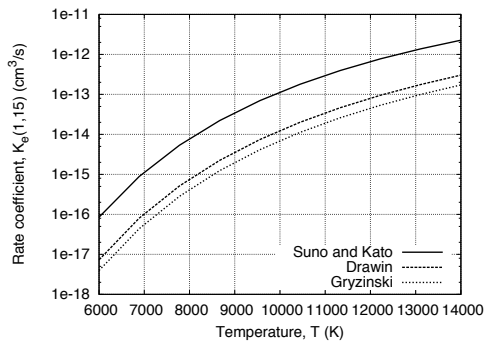
(b) $2s2p^2^4P \Rightarrow 2s2p^2^2D$ (Forbidden)



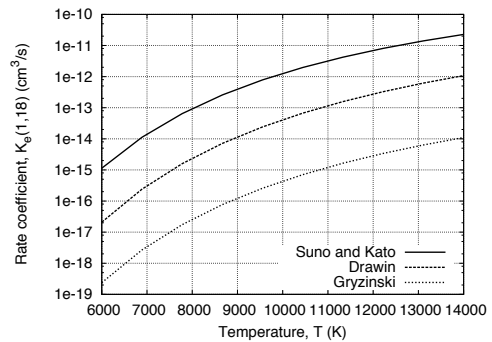
(c) $2s^22p^2P^o \Rightarrow 2s2p^2^2P$ (Allowed)



(d) $2s2p^2^4P \Rightarrow 2s2p^2^2P$ (Forbidden)

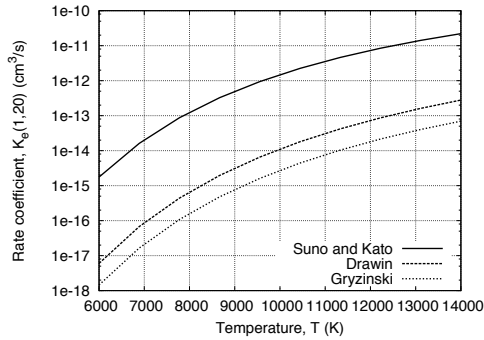


(e) $2s^22p^2P^o \Rightarrow 2p^3^2P^o$ (Forbidden)

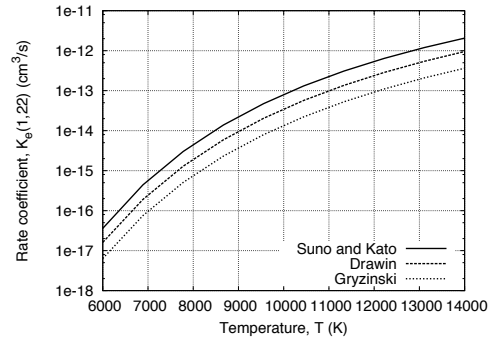


(f) $2s^22p^2P^o \Rightarrow 2s^25p^2P^o$ (Allowed)

Figure 3.18: Comparison of electron impact excitation rate coefficients for atomic carbon.



(g) $2s^22p^2P^\circ \Rightarrow 2s^25d^2D$ (Forbidden)



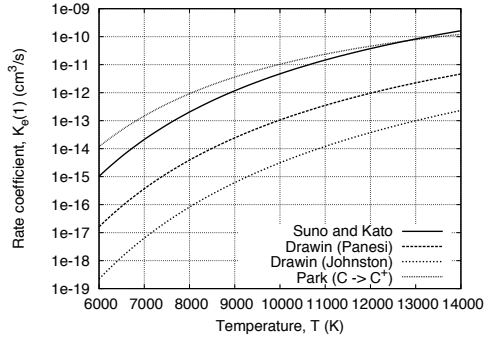
(h) $2s^22p^2P^\circ \Rightarrow 2s^25g^2G$ (Allowed)

Figure 3.17: (Continued) Comparison of electron impact excitation rate coefficients for atomic carbon.

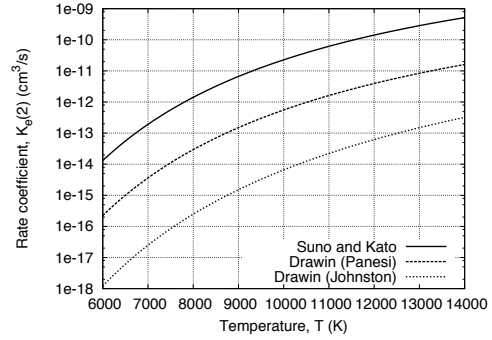
Figure 3.18 compares the electron impact ionisation rate coefficient for various transitions of C. The bulk atomic carbon ionisation rate proposed by Park [65] is overlaid in Figure 3.18a for comparison with the ground state rates. The Suno and Kato rates are approximately three orders of magnitude greater than the semi-empirical models, and is similar to the Park rate for the ground state. Similarly as for atomic nitrogen and oxygen, Johnston's implementation of the Drawin cross sections is between approximately 2 and 100 times smaller than Panesi's implementation for all levels. Therefore in the present work the electron impact ionisation coefficients of Suno and Kato [62] are preferred for the first three levels, while the Drawin model implemented by Panesi [12] is used for the remaining levels.

Diatomic species: C₂, CN and CO

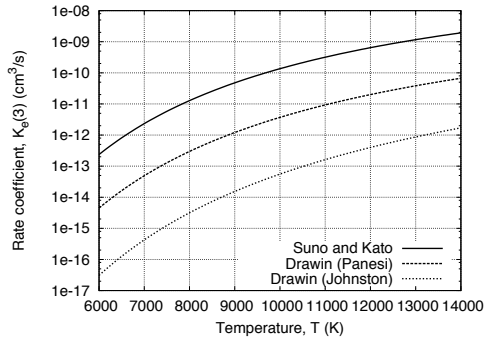
Zalugin [66] proposed a simple collisional-radiative model for C₂, CN and CO. This model was 'tuned' to match the intensity profiles measured for a 3.45 km/s CO₂-N₂ shock tube condition, and was applied with limited success to the 8.5 km/s CO₂-N₂ EAST shock tube condition in Reference [67]. Since this model was formulated, however, a set of collisional-rate coefficients and cross sections for the diatomic species CN, CO, N₂, N₂⁺, O₂ and NO have been proposed by Park [56,57]. Figure 3.19 compares the rates that are given by both Zalugin [66] and Park [56,57]. The Zalugin model predicts faster collisional excitation of CN and slower collisional excitation of CO. As the Park model is more comprehensive and based on experimental and theoretical cross sections where possible, the Park collisional excitation rates for CO and CN are tentatively accepted for inclusion in the nominal collisional-radiative models for these species. Due to numerical difficulties encountered when implementing the heavy particle impact rates given in Reference [57], however, these were omitted for CO and CN. Considering just the electron impact processes should be sufficient for the conditions of interest due to the substantial levels of ionisation. It should be emphasised that the heavy particle impact processes would need to be considered for the collisional-radiative model to be valid for less energetic conditions where ionisation levels are low. The system radiative transition probabilities are calculated from the electronic-vibration transition moments via Equation 3.109. For C₂ we resort to the Zalugin [66] model as this species was not considered



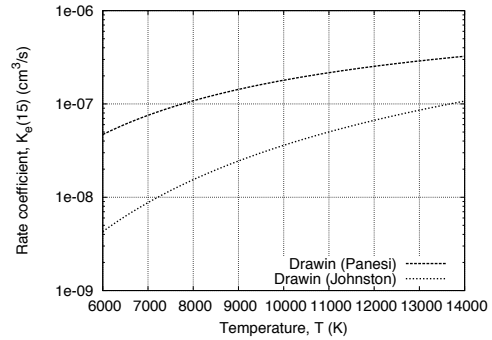
(a) $C 2s^2 2p^2 P^o \Rightarrow C^+$



(b) $C 2s 2p^2 4 P \Rightarrow C^+$



(c) $C 2s 2p^2 2 D \Rightarrow C^+$



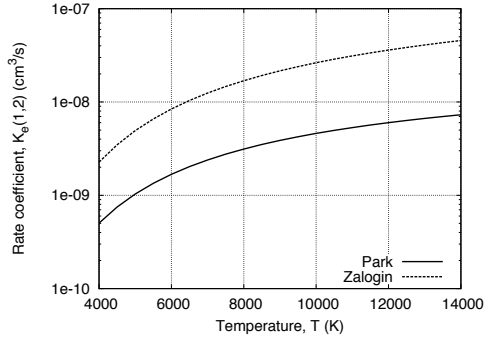
(d) $C 2p^3 2 P^o \Rightarrow C^+$

Figure 3.18: Comparison of electron impact ionisation rate coefficients for atomic carbon.

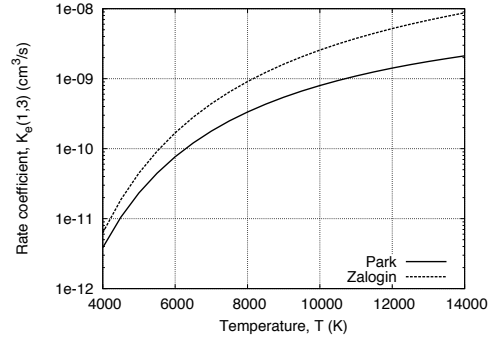
by Park in References [56,57]. The nominal collisional-radiative models for C_2 , CN and CO are summarised in § A.1 to A.3 respectively.

3.2.5 A note on the selection of data sources

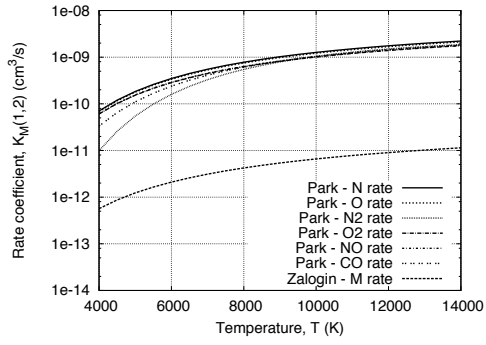
In formulating the collisional-radiative model the sometimes contradictory issues of data source consistency and model accuracy arise. On one hand, it is desirable to formulate a model that is both internally consistent (uses the same data source for all processes) and externally consistent (uses the same data sources as other models, such as the spectral radiation model). On the other hand, it is desirable to formulate a model that best reproduces the physical phenomena. The goal of the present work is to develop effective engineering tools, and therefore the latter approach has been preferred to give the calculations the best chance of reproducing experimental data. As an example, consider the electron impact excitation for the nitrogen atom. An internally consistent model would have to use an empirical model such as that of Drawin [39] for all transitions, however this model has been shown to give inaccurate results when applied to the excitation of low lying electronic states [38,48]. Although it breaks the internal consistency of the model, implementing the computational chemistry rates of Frost *et al.* [45] gives much improved comparisons with experiment [38,48]. This principal of selecting the most effective data, rather than the most consistent data, has been applied throughout



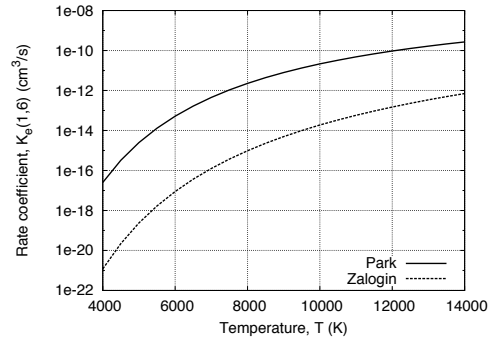
(a) $\text{CN } X^2\Sigma^+ + e^- \Rightarrow \text{CN } A^2\Pi + e^-$



(b) $\text{CN } X^2\Sigma^+ + e^- \Rightarrow \text{CN } B^2\Sigma^+ + e^-$



(c) $\text{CN } X^2\Sigma^+ + \text{M} \Rightarrow \text{CN } A^2\Pi + \text{M}$



(d) $\text{CO } X^1\Sigma^+ + e^- \Rightarrow \text{CO } A^1\Pi + e^-$

Figure 3.19: Comparison of excitation rate coefficients obtained from Park [56, 57] and Zalogin [66].

the collisional-radiative model.

4

The equilibrium air model

The equilibrium air spectral model makes the gray-gas approximation to define a mean absorption coefficient. The empirical correlation of Olstad [68] is implemented:

$$\kappa_P = 7.94 \left(\frac{\rho}{\rho_0} \right)^{1.10} \left(\frac{T}{10^4} \right)^{6.95} \quad (4.1)$$

where the numerical factor at the front has been changed so that κ_P has units of m^{-1} instead of ft^{-1} and ρ_0 is sea-level density, taken to be 1.225 kg/m^3 . The total emissivity is then calculated from the Stefan-Boltzmann equation:

$$j = \frac{\kappa_P \sigma T^4}{\pi} \quad (4.2)$$

5

Other spectral models

Fluid Gravity's Parade code and KAIST's Spradian07 code can be used to compute the spectral coefficients within the `cfcd3` framework if the user has access to these codes. For documentation on the Parade and Spradian07 codes see References [2] and [3], respectively.

Part II

User's Guide

6

Installation

The core of the radiation module is written in C++, and there are a number of tools written in Python. The input data is provided to the C++ binaries via Lua files. Assuming the Compressible-Flow CFD repository is located in `~/cfcfd3-hg`, the radiation module can be compiled and installed via:

```
$ cd ~/cfcfd3-hg/lib/radiation/build
$ make install
```

The installation of the radiation library can then be tested via an automated script:

```
$ cd ~/cfcfd3-hg/lib/radiation/build
$ make test
```

If the user has the SPRADIAN07 code available and wishes to use it as the spectral model, this needs to be requested at compile time:

```
$ cd ~/cfcfd3-hg/lib/radiation/build
$ make WITH_SPRADIAN=1 install
```

The SPRADIAN07 source code and input-data files need to be located in `~/cfcfd3-hg/extern/spradian07`:

```
$ ls ~/cfcfd3-hg/extern/spradian07/
atom.dat      diatom.dat   radipac6.f90 triatom.dat
```

If the Parade code is available to the user, no additional compilation arguments are need to use it as the spectral model. The `parade` binary, however, needs to be in the system path and the Parade template file needs to include the location of the Parade data files (the Parade template file can be automatically generated from the python input script - see `lib/radiation/test/parade-air-EQ-radiators.py` for an example).

To test all the spectral models the following make command can be run:

```
$ cd ~/cfcfd3-hg/lib/radiation/build
$ make WITH_SPRADIAN=1 WITH_PARADE=1 test
```

Creating Input Files

The input data to create a new spectral radiation model is provided in a Lua file. This file contains all the parameters and spectroscopic data to perform a spectral radiation calculation. To simplify the creation of this file for the user, the Lua file can be created via the `radmodel.py` Python tool that interfaces with a library of spectroscopic data. Instructions on how to use this tool can be obtained from the command line via:

```
$ radmodel.py --help
Usage: radmodel.py -i rad_desc.py|--input-script=rad_desc.py
                    -L LUA_output.lua|--LUA-file=LUA_output.lua

Options:
  -h, --help            show this help message and exit
  -i INFILE, --input-script=INFILE
                        input Python script for radiation description
  -L LUAFILE, --LUA-file=LUAFILE
                        output configuration file for 'librad' C++ module in
                        LUA format
```

where `rad-model.py` is the name of the user created Python script describing the spectral model and `rad-model.lua` is the desired name of the resulting Lua file. § 7.1 describes how to construct a basic radiation input file that allows spectra to be calculated via the spectral model described in § 3, while the process of creating a more advanced input file for modelling monatomic plasmas to a higher degree of accuracy is described in § 7.2. Examples of input files for interfacing with the PARADE and SPRADIAN07 codes via the `cfcfd3/lib/radiation` framework are provided in textsection 7.3 and § 7.4.

7.1 Creating an input file for the baseline modelling

A sample user-created Python script for creating a basic radiation model for 11 species, 2 temperature air is shown below:

Listing 7.1: Example spectral model input file, `air-radiators.py`

```
# Filename: air-radiators.py
# Author: Daniel F. Potter
# Date: 11th of March 2013
# Usage: radmodel.py -i air-radiators.py -L rad-model.lua
# Define a simple radiation model for two temperature, 11 species air

# 1. Select the spectral model
```

```

gdata.spectral_model = "photaura"

# 2. Define the spectral grid
gdata.lambda_min = 70.0
gdata.lambda_max = 1200.0
gdata.spectral_points = 113000

# 3. Request and define the radiating species
species = [ "N2", "N2_plus", "NO", "NO_plus", "O2", "O2_plus",
            "N", "N_plus", "O", "O_plus", "e_minus" ]
radiators = [ "N2", "N2_plus", "NO", "O2", "N", "N_plus",
              "O", "O_plus", "e_minus" ]
for rad_name in radiators:
    rad = gdata.request_radiator(rad_name)
    rad.default_data()
    rad.isp = species.index(rad_name)
    rad.iTe = 1
    if rad.type == "diatomic_radiator":
        rad.iTv = 1

# radmodel.py does the rest!

```

In '1.', the desired spectral model is selected:

```

# 1. Select the spectral model
gdata.spectral_model = "photaura"

```

Here the in-house photaura model has been selected; other available models are spradian and parade. See § 7.3 and § 7.4 for examples of input files for these models. The spectral grid is then defined in '2.':

```

# 2. Define the spectral grid
gdata.lambda_min = 70.0
gdata.lambda_max = 1200.0
gdata.spectral_points = 113000

```

Here the range $70 \leq \lambda \leq 1200$ nm has been requested, discretised by 113000 equidistant points in frequency space. Finally, in '3.' the desired radiating species are requested and defined:

```

# 3. Request and define the radiating species
species = [ "N2", "N2_plus", "NO", "NO_plus", "O2", "O2_plus",
            "N", "N_plus", "O", "O_plus", "e_minus" ]
radiators = [ "N2", "N2_plus", "NO", "O2", "N", "N_plus",
              "O", "O_plus", "e_minus" ]
for rad_name in radiators:
    rad = gdata.request_radiator(rad_name)
    rad.default_data()
    rad.isp = species.index(rad_name)
    rad.iTe = 1
    if rad.type == "diatomic_radiator":
        rad.iTv = 1

```

A radiator is requested from the library with the `gdata.request_radiator()` function call. If the radiator is not present in the library, `radmodel.py` will fail here with an error message indicating what species was not available. The `rad.default_data()` function requests the nominal electronic level and transition probability set from the library – for most calculations this will be suitable. Other data that is set here is the species

index in the `rad.iSp` field, the electronic temperature index in the `rad.iTe` field and the vibrational temperature index in the `rad.iTv` field.

7.2 Creating an input file for the advanced modelling of monatomic plasma

The baseline modelling makes a number of assumptions regarding atomic lines and photoionisation that enable rapid calculations but reduces the accuracy of the resulting spectra. Specifically, these assumptions are the use of approximate Stark broadening widths and hydrogenic photoionisation cross-sections. The input file shown below demonstrates how to create an input file that allows accurate Stark broadening [69] and photoionisation cross-sections [34] to be used:

Listing 7.2: Example advanced spectral model input file, `argon-radiators-NIST-TB.py`

```
# Filename: argon-radiators-NIST-TB.py
# Author: Daniel F. Potter
# Date: 20th of October 2013
# Usage: radmodel.py -i argon-radiators-NIST-TB.py -L rad-model.lua
#
# This file demonstrates the creation of an advanced radiation model using individual
# lines and levels from NIST ASD, Griem Stark widths, OPBase photoionisation
# cross-sections and adaptive spectral gridding.

# 1. Select the spectral model
gdata.spectral_model = "photaura"

# 2. Define the spectral grid
gdata.lambda_min = 1.0e7 / 150000.0          # 150000 cm-1
gdata.lambda_max = 1.0e7 / 1000.0           # 1000 cm-1
gdata.spectral_points = 100
gdata.adaptive_spectral_grid = True

# 3. Request and define the radiating species
params = {
"species"           : [ 'Ar', 'Ar_plus', 'e_minus' ],
"radiators"         : [ 'Ar', 'Ar_plus', 'e_minus' ],
"QSS_radiators"     : [ 'Ar' ],
"no_emission_radiators" : [ 'Ar_plus' ],
"iTe"               : 1,
"atomic_level_source" : "NIST_ASD",
"atomic_line_source"  : "NIST_ASD",
"atomic_PICS_source"  : "TOPBase"
}
declare_radiators( params, gdata )

# 4. Customise some of the parameters
# line adaptation
gdata.radiators[params["radiators"].index('Ar')].line_set.npoints = 50
gdata.radiators[params["radiators"].index('Ar_plus')].line_set.npoints = 10
# QSS lower temperature limit
gdata.radiators[params["radiators"].index('Ar')].QSS_model.T_lower = 0.0
```

In '1.', the desired spectral model is selected:

```
# 1. Select the spectral model
gdata.spectral_model = "photaura"
```

Here the in-house photaura model has again been selected. The selection of this model is required to enable the advanced modelling of monatomic plasmas. The spectral grid is then defined in '2.':

```
# 2. Define the spectral grid
gdata.lambda_min = 1.0e7 / 150000.0      # 150000 cm-1
gdata.lambda_max = 1.0e7 / 1000.0        # 1000 cm-1
gdata.spectral_points = 100
gdata.adaptive_spectral_grid = True
```

Here the spectral range $1000 \leq \eta \leq 150000 \text{ cm}^{-1}$ has been requested, with an adaptive spectral grid. The photaura model will determine the appropriate spectral grid to accurately capture the atomic lines and bound-free continuum, which is then superimposed on top of a uniformly spaced grid with `gdata.spectral_points`. To ensure that the free-free continuum is accurately captured, it is recommended `gdata.spectral_points` be set to a small but non-zero value; 100 uniform points has been found to be appropriate for cases where the free-free continuum contributes less than 1% of the total emissive power.

In '3.' the desired radiating species are requested and defined in a python dictionary called `params`:

```
# 3. Request and define the radiating species
params = {
"species"           : [ 'Ar', 'Ar_plus', 'e_minus' ],
"radiators"         : [ 'Ar', 'Ar_plus', 'e_minus' ],
"QSS_radiators"     : [ 'Ar' ],
"no_emission_radiators" : [ 'Ar_plus' ],
"iTe"               : 1,
"atomic_level_source" : "NIST_ASD",
"atomic_line_source"  : "NIST_ASD",
"atomic_PICS_source"  : "TOPBase"
}
declare_radiators( params, gdata )
```

The "species" field is the list, in order, of the species present in the gas or CFD flow-field under consideration. The "radiators" field is the list of radiating species to be considered. Electrons need to be included if the bound-free and free-free continua wish to be modelled. Species that are to have their electronic populations determined via the QSS approach are listed in the "QSS_radiators" field. The "no_emission_radiators" field can be used to 'turn off' the radiation of a radiator. This is useful when the bound-bound radiation from an ion is very weak, but the radiator needs to be declared to enable the photoionisation of its neutral counterpart to be modelled. The "iTe" field species the index of the electron temperature used in the flowfield solution. The remaining three fields all the user to control what datasources and models are used for the atomic lines, levels and photoionisation (PICS) cross-section. The "atomic_level_source" and "atomic_line_source" fields need to be the same, and selected from one of the following:

NIST_ASD Line and level data from the NIST Atomic Species Database [10]

TOPBase Line and level data from the Opacity Project database [34]

The "atomic_pics_source" field defines the photoionisation modelling and can be one of the following:

hydrogenic Hydrogenic approximation of the photoionisation cross-sections (see § 3.1.3)

TOPBase Theoretical photoionisation cross-sections from the Opacity Project database [34]

The hydrogenic model can overpredict the bound-free emission, especially where high ($\geq 10\%$) levels of ionisation are present. In these cases it is recommended the TOPbase model is used. Currently the input data required for this advanced treatment of monatomic radiation has only been collected for Ar-I and Ar-II (see `cfcfd3/lib/radiation/data`). If the user wishes to use the advanced modelling for other species, the existing files can be used as a guide and the raw data can be downloaded from the following websites:

NIST ASD <http://www.nist.gov/pml/data/asd.cfm>

TOPBase <http://cdsweb.u-strasbg.fr/topbase/topbase.html>

Griem Stark broadening <http://griem.obspm.fr/>

Finally, the radiator parameters can be adjusted from the default values as has been demonstrated in '4.':

```
# 4. Customise some of the parameters
# line adaptation
gdata.radiators[params["radiators"].index('Ar')].line_set.npoints = 50
gdata.radiators[params["radiators"].index('Ar_plus')].line_set.npoints = 10
# QSS lower temperature limit
gdata.radiators[params["radiators"].index('Ar')].QSS_model.T_lower = 0.0
```

The interested user should consult the declarations of the python classes in `cfcfd3/lib/radiation/tools/rl_defs.py`.

7.3 Examples input file interfacing with PARADE

Below is an example python input file for modelling an argon plasma with the PARADE code inside the `cfcfd3/lib/radiation` c++ framework:

Listing 7.3: Example PARADE spectral model input file, `argon-radiators-parade.py`

```
# Filename: argon-radiators-parade.py
# Author: Daniel F. Potter
# Date: 20th of October 2013
# Usage: radmodel.py -i argon-radiators-parade.py -L rad-model.lua
#
# This file demonstrates the creation of a radiation input file for modelling an argon
# plasma with the parade code from FluidGravity.

# 1. Select the spectral model
gdata.spectral_model = "parade"

# 2. Define the spectral grid
gdata.lambda_min = 1.0e7 / 150000.0
gdata.lambda_max = 1.0e7 / 1000.0
gdata.spectral_points = int ( ( 1.0e7 / gdata.lambda_min - 1.0e7 / gdata.lambda_max ) * 1.0 )

# 3. Request and define the radiating species
radiators = [ "Ar", "Ar_plus", "e_minus" ]
species = [ "Ar", "Ar_plus", "e_minus" ]
for rad_name in radiators:
```

```

rad = gdata.request_radiator(rad_name)
rad.isp = species.index(rad_name)
rad.iTe = 0

```

```

# 3. Specify the location of the Parade input data files and create template files
parade_data_path = "/Users/dpotter/myproject/DLR/work/programs/parade31/Data"
gdata.create_parade_template_files(parade_data_path)

```

7.4 Examples input file interfacing with SPRADIAN07

Currently the user must manually create the .lua input file when using the 'spradian' spectral model. Below is an example Lua input file for modelling an argon plasma with the SPRADIAN07 code inside the cfcfd3/lib/radiation c++ framework:

Listing 7.4: Example SPRADIAN07 spectral model input file, argon-radiators-spradian.lua

```

-- Filename: argon-radiators-parade.py
-- Author: Daniel F. Potter
-- Date: 20th of October 2013
-- Usage: ready to use as the input file for creating the radiation model
-- e.g. in python:
-- import radpy
-- rsm = create_radiation_spectral_model( "argon-radiators-spradian.lua" )

spectral_data = {
    spectral_model = 'spradian',
    iT = 0,
    iTr = 0,
    iTv = 0,
    iTe = 0,
    radiators = { 'Ar', 'Ar_plus', 'e_minus' },
    lambda_min = 66.666667,
    lambda_max = 10000.000000,
    spectral_points = 14900,
    spectral_blocks = 1,
    adaptive_spectral_grid = false,
}

Ar = {}
Ar.isp = 0
Ar.type = 'atomic_radiator'
Ar.mol_weight = 3.994800e-02

Ar_plus = {}
Ar_plus.isp = 0
Ar_plus.type = 'atomic_radiator'
Ar_plus.mol_weight = 3.994745e-02

e_minus = {}
e_minus.isp = 0
e_minus.type = 'atomic_radiator'
e_minus.mol_weight = 5.485799e-07

```


8

Examples

- 8.1 VKI minitorch
- 8.2 EAST CO₂-N₂ shock tube
- 8.3 Hayabusa
- 8.4 Rutowski hemisphere

References

- [1] P.A. Jacobs, R.J. Gollan, T. Marty, V. Wheatley, and D.F. Potter. The compressible-flow cfd project. Viewed 20th of October 2013, <http://mech.uq.edu.au/cfcfd/>, October 2013.
- [2] M. Winter, B. Pfeiffer, M. Fertig, M. Auweter-Kurtz, and A.J. Smith. Recent status of the Plasma Radiation Database PARADE. *37th AIAA Thermophysics Conference, Portland, Oregon, (AIAA 2004-2585)*, 28 June - 1 July 2004.
- [3] S.-Y. Hyun. *Radiation code SPRADIAN07 and its applications*. PhD thesis, KAIST, Daejeon, Korea, 2009.
- [4] M F Modest. *Radiative Heat Transfer*. Academic Press, London, 2nd edition, 2003.
- [5] D. Potter. *Modelling of radiating shock layers for atmospheric entry at earth and mars*. PhD thesis, School of Mechanical and Mining Engineering, University of Queensland, Australia, 2011.
- [6] R. Sobbia, P. Leyland, Y. Babou, D. Potter, and L. Maraffa. Plasma radiation for atmospheric entry at titan: Emission spectroscopy measurements and numerical re-building. *Journal of Plasma Physics*, 79(5):525–534, October 2013.
- [7] Y.B. Zel'dovich and Y.P. Raizer. *Physics of shock waves and high-temperature hydrodynamic phenomena*. New York: Academic Press, 1967.
- [8] K.P. Huber and G. Herzberg. *Molecular Spectra and Molecule Structure IV. Constants of Diatomic Molecules*. Van Nostrand Reinhold, New York, 1950.
- [9] I. Kovács. *Rotational structure in the Spectra of Diatomic molecules*. Adam Hilger Ltd., London, 1969.
- [10] Yu. Ralchenko, A.E. Kramida, J. Reader, and the NIST ASD Team. NIST Atomic Spectra Database (version 3.1.5). [Online]. Available: <http://physics.nist.gov/asd3> [2010, March 22]. National Institute of Standards and Technology, Gaithersburg, MD., 2008.
- [11] C. S. Park. *Nonequilibrium Hypersonic Aerothermodynamics*. John Wiley and Sons, 1990.
- [12] C. Johnston. *Nonequilibrium shock-layer radiative heating for Earth and Titan entry*. PhD thesis, Virginia Polytechnic Institute and State University, Blacksburg, Virginia, November 2006.
- [13] E.E. Whiting. An empirical approximation to the Voigt profile. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 8(6):1374 – 1384, 1968.

- [14] J.J. Olivero and R.L. Longbothum. Empirical fits to the Voigt line width: a brief review. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 17(2):233–236, 1977.
- [15] W. E. Nicolet. Advanced methods for calculating radiation transport in ablation-product contaminated boundary layers. Contractor Report CR-1656, NASA, Washington, D.C., September 1970.
- [16] G. Traving. *Plasma diagnostics*, chapter Interpretation of Line Broadening and Line Shift. AIP Press, 1995.
- [17] H.R. Griem. *Plasma Spectroscopy*. McGraw Hill, 1964.
- [18] W.A. Page, D.L. Compton, W.J. Borucki, and D.L. Ciffone. Radiative transfer in inviscid non-adiabatic stagnation region shock layers. AIAA Paper 68-784, 1968.
- [19] A.P. Thorne. *Spectrophysics*. Chapman and Hall, 1974.
- [20] S.A. Golden. Approximate spectral absorption coefficients of electronic transitions in diatomic molecules. *Journal of Quantitative Radiation Transfer*, 7:225–250, 1967.
- [21] C.O. Laux. Radiation and nonequilibrium collisional-radiative models. VKI special course on physico-chemical models for high enthalpy and plasma flows modeling, Rhode-St-Genése, Belgium, June 4-7 2002.
- [22] C.O. Laux, T.G. Spence, C.H. Kruger, and R.N. Zare. Optical diagnostics of atmospheric pressure air plasmas. *Plasma Sources Science and Technology*, 12:125–138, 2003.
- [23] Y. Babou, P. Rivière, M.-Y. Perrin, and A. Soufiani. High-temperature and nonequilibrium partition function and thermodynamic data of diatomic molecules. *International Journal of Thermophysics*, 30:416–438, 2009.
- [24] K.P. Huber and G. Herzberg. Constants of diatomic molecules. In P.J. Linstrom and W.G. Mallard, editors, *NIST Chemistry WebBook, NIST Standard Reference Database Number 69*. National Institute of Standards and Technology, Gaithersburg MD, 20899, <http://webbook.nist.gov>, (retrieved May 8, 2010).
- [25] C.V.V. Prasad and P.F. Bernath. Fourier transform jet-emission spectroscopy of the $A^2\pi_i-X^2\sigma^+$ transition of cn. *Journal of Molecular Spectroscopy*, 156(2):327–340, December 1992.
- [26] S. Chauveau, M.-Y. Perrin, P. Rivière, and A. Soufiani. Contributions of diatomic molecular electronic systems to heated air radiation. *Journal of quantitative spectroscopy and radiative transfer*, 72(4):503–530, February 2002.
- [27] Y. Babou, P. Rivière, M.-Y. Perrin, and A. Soufiani. Spectroscopic data for the prediction of radiative transfer in $\text{CO}_2\text{-N}_2$. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 110:89–108, 2009.
- [28] R.S. Mulliken. The interpretation of band spectra. Part IIc. Empirical band types. *Review of Modern Physics*, 3(89):89–115, 1931.
- [29] J. O. Arnold, E. E. Whiting, and G. C. Lyle. Line by line calculation of spectra from diatomic molecules and atoms assuming a Voigt line profile. *Journal of Quantitative Spectroscopy and Radiation Transfer*, 9:775–798, 1969.

- [30] C.O. Laux. *Optical diagnostics and radiative emission of air plasmas*. PhD thesis, Stanford University, 1993.
- [31] C.O. Laux and C.H. Knucinnut. Arrays of radiative transition probabilities for the N₂ First and Second Positive, NO Beta and Gamma, N₂⁺ First Negative and O₂ Schumann-Runge band systems. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 48(1):9–24, 1992.
- [32] C.O. Johnston, B.R. Hollis, and K. Sutton. Spectrum modeling for air shock-layer radiation at lunar-return conditions. *Journal of Spacecraft and Rockets*, 45(5):856–878, September-October 2008.
- [33] S. Chauveau, C. Deron, M.-Y. Perrin, P. Rivière, and A. Soufiani. Radiative transfer in LTE air plasmas for temperatures up to 15,000 K. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 77(2):113–130, March 2003.
- [34] W. Cunto, C. Mendoza, F. Ochsenbein, and C.J. Zeippen. TOPbase at the CDS. *Astronomy and Astrophysics*, 275(1):5–8, August 1993.
- [35] L.C. Hartung. *Nonequilibrium radiative heating prediction method for aeroassist flow-fields with coupling to flowfield solvers*. PhD thesis, Department of Mechanical and Aerospace Engineering, North Carolina State University, 1991.
- [36] W.L. Wiese. New reference data table for carbon, nitrogen and oxygen spectra. *Spectrochimica Acta, Part B: Atomic Spectroscopy*, 51(8):775 – 777, 1996.
- [37] B. Kleb and C.O. Johnston. Uncertainty analysis of air radiation for lunar return shock layers. AIAA Paper 2008-6388, Honolulu, Hawaii, August 2008.
- [38] M. Panesi. *Physical models for nonequilibrium plasma flow simulations at high speed re-entry conditions*. PhD thesis, von Karman Institute for Fluid Dynamics, Rhode-St-Genése, Belgium, 2009.
- [39] H.W. Drawin. *Plasma diagnostics*, chapter Collision and transport cross sections, pages 842–876. 1968.
- [40] M. Gryzinski. Classical theory of electronic and ionic inelastic collisions. *Physical review*, 115:374–383, 1959.
- [41] C.W. Allen. *Astrophysical Quantities*. London, 1962.
- [42] H. Van Regemorter. Rate of collisional excitation in stellar atmospheres. *Astrophysical Journal*, 136:906–915, 1962.
- [43] C.S. Park. Comparison of electron and electronic temperature in a recombining nozzle flow of ionized nitrogen - hydrogen mixture: Part 1. Theory. *Journal of Plasma Physics*, 9:187–215, 1973.
- [44] M. Panesi, T. Magin, A. Bourdon, A. Bultel, and O. Chazot. Analysis of the FIRE II flight experiment by means of a collisional radiative model. AIAA Paper 2008-1205, January 2008.
- [45] R.L. Frost, P. Awakowicz, H.P. Summers, and N.R. Badnell. Calculated cross sections and measured rate coefficients for electron-impact excitation of neutral and singly ionized nitrogen. *Journal of Applied Physics*, 84(6):2989 – 3003, 1998.

- [46] W.H. Soon and J.A. Kunc. Thermal nonequilibrium in partially ionized atomic oxygen. *Physical Review A*, 41(2):825–843, January 1990.
- [47] O. Zatsarinny and S.S. Tayal. Electron collisional excitation rates for O I using the B-spline R-matrix approach. *The Astrophysical Journal Supplement Series*, 148:575–582, October 2003.
- [48] M. Panesi, Y. Babou, and O. Chazot. Predictions of nonequilibrium radiation: analysis and comparison with EAST experiments. Aiaa 2008-3812, Seattle, Washington, June 2008.
- [49] A. Bultel, B.G. Chéron, A. Bourdon, O. Motapon, and I.F. Schneider. Collisional-radiative model in air for Earth re-entry problem. *Physics of Plasmas*, 13(4):11, April 2006.
- [50] K.A. Berrington, P.G. Burke, and W.D. Robb. The scattering of electrons by atomic nitrogen. *Journal of Physics B: Atomic and molecular physics*, 8(15):2500–2511, October 1975.
- [51] J.A. Kunc and W.H. Soon. Collisional-radiative nonequilibrium in partially ionized atomic nitrogen. *Physical Review A*, 40(10):5822–5842, 1989.
- [52] H. Tawara and M. Kato. Electron impact ionization data for atoms and ions - updated in 1998. *NIFS Data Series S*, 51:596p, 1999.
- [53] Y.-K. Kim and J.-P. Desclaux. Ionization of carbon, nitrogen and oxygen by electron impact. *Physical Review A*, 66(1):012708 1–12, July 2002.
- [54] Y. Itikawa and A. Ichimura. Cross sections for collisions of electrons and photons with atomic oxygen. *Journal of Physical and Chemical Reference Data*, 19(3):637–652, 1990.
- [55] P. Teulet, J.P. Sarrette, and A.M. Gomes. Calculation of electron-impact inelastic cross-section and rate coefficients for diatomic molecules. Application to air molecules. *Journal of Quantitative Spectroscopy and Radiative Transfer*, 62(5):549–569, July 1999.
- [56] C.S. Park. Rate parameters for electronic excitation of diatomic molecules II. Electron-impact processes. AIAA Paper 2008-1446, Reno, Nevada, 2008.
- [57] C.S. Park. Rate parameters for electronic excitation of diatomic molecules II. Heavy particle-impact processes. AIAA Paper 2008-1446, Reno, Nevada, 2008.
- [58] C.O. Johnston. A comparison of east shock-tube radiation measurements with a new air radiation model. AIAA Paper 2008-1245, Reno, Nevada, 2008.
- [59] R.D. Braun, R.W. Powell, and L.C. Hartung. Effect of interplanetary trajectory options on a manned Mars aerobrake configuration. Technical Paper 3019, NASA Langley Research Center, 1 September 1990.
- [60] J. Vlček. A collisional-radiative model applicable to argon discharges over a wide range of conditions. I: Formulation and basic data. *Journal of Physics D: Applied Physics*, 22(5):623–631, May 1989.

- [61] A. Bogaerts, R. Gijbels, and J. Vlcek. Collisional-radiative model for an argon glow discharge. *Journal of Applied Physics*, 84(1):121–136, July 1998.
- [62] H. Suno and T. Kato. Cross section database for carbon atoms and ions: Electron-impact ionization, excitation, and charge exchange in collisions with hydrogen atoms. *Atomic Data and Nuclear Tables*, 92(4):407–455, July 2006.
- [63] K.M. Dunseath, W.C. Fon, V.M. Burke, R.H.G. Reid, and C.J. Noble. Electron-impact excitation of the $n \leq 4$ levels of carbon. *Journal of Physics B: Atomic, Molecular and Optical Physics*, 30(2):277–287, 1997.
- [64] E. Brook, M.F.A. Harrison, and A.C.H. Smith. Measurements of electron impact ionisation cross sections of He, C, O and N atoms. *Journal of Physics B (Atomic and Molecular Physics)*, 11(17):3115–3132, September 1978.
- [65] C.S. Park, J.T. Howe, R.L. Jaffe, and G.V. Candler. Review of chemical-kinetic problems of future NASA missions, II: Mars entries. *Journal of Thermophysics and Heat Transfer*, 8(1):9 – 22, 1994.
- [66] G.N. Zalogin, P.V. Kozlov, L.A. Kuznetsova, S.A. Losev, V.N. Makarov, Yu.V. Romanenko, and S.T. Surzhikov. Radiation excited by shock waves in a CO₂-N₂-Ar mixture: Experiment and theory. *Technical Physics (Russia)*, 46(6):654 – 61, 2001.
- [67] D.F. Potter, R.J. Gollan, P.A. Jacobs, and P. Leyland. Numerical simulations and analysis of the 8.5 km/s CO₂-N₂ EAST shock tube condition. In H. Lacoste and L. Ouwehand, editors, *Proceedings of the 3rd International Workshop on Radiation of High Temperature Gases in Atmospheric Entry*, number ESA SP-667, Heraklion, Greece, 2008. European Space Agency.
- [68] W.B. Olstad. Stagnation-point solutions for an inviscid radiating shock layer. In *Proceedings of the 1965 Heat Transfer and Fluid Mechanics Institute*, pages 138–156. Stanford University Press, 1965.
- [69] H.R. Griem. *Spectral line broadening by plasmas*. Academic Press, New York, 1974.
- [70] W.M. Huo and H.T. Thuemmel. *Molecular physics and hypersonic flows*, chapter Electron-air molecule collisions in hypersonic flows, pages 115–138. Kluwer Academic Publishers, 1995.
- [71] J. Zobel, U. Mayer, K. Jung, and H. Ehrhardt. Absolute differential cross sections for electron-impact excitation of CO near threshold: I. The valence states of CO. *Journal of Physics B: Atomic, Molecular and Optical physics*, 29(4):813–838, February 1996.
- [72] L.A. Morgan and J. Tennyson. Electron impact excitation cross sections for CO. *Journal of Physics B: Atomic, Molecular and Optical Physics*, 26(15):2429–2441, August 1993.
- [73] R. Olszewski, P. Wolinski, and M. Zubek. Excitation of carbon monoxide by electron impact in the 8-17 eV energy range. *Chemical Physics Letters*, 297(5-6):537–542, December 1998.
- [74] G.G. Chernyi and S.A. Losev. Development of thermal protection system for interplanetary flight. ISTC Report 036-96, International Science and Technology Center, Moscow State University, August 1999.

- [75] M. Capitelli, C.M. Ferreira, B.F. Gordiets, and A.I. Osipov. *Plasma Kinetics in Atmospheric Gases*. Springer-Verlag, Berlin, 2000.
- [76] I.V. Kurochkin, L.S. Polak, A.V. Pustogarov, D.I. Slovetskii, and V.V. Ukolov. Thermochemical nonequilibrium in an arc plasma stabilized by nitrogen injected through the porous wall of a channel - Particle excitation mechanisms and vibrational nonequilibrium. *High temperature (English translation)*, 16(6):995–1004, May 1979.
- [77] F. Fresnet, G. Baravian, L. Magne, S. Pasquiers, C. Postel, V. Puech, and A. Rousseau. Influence of water on NO removal by pulsed discharge in N₂/H₂O/NO mixtures. *Plasma Sources Science and Technology*, 11(2):152+160, 2002.
- [78] S.V. Pancheshnyi, S.M. Starikovskaia, and A.Yu. Starikovskii. Collisional deactivation of N₂(C³π_u, v = 0, 1, 2, 3) states by N₂, O₂, H₂ and H₂O molecules. *Chemical Physics*, 262:349–357, October 2000.
- [79] V.A. Gorelov, M.K. Gladyshev, A.Y. Kireev, I.V. Yegorov, Y.A. Plastinin, and G.F. Karabadzhak. Experimental and numerical study of nonequilibrium ultraviolet NO and N emission in shock layer. *Journal of Thermophysics and Heat Transfer*, 12(2):172–179, 1998.
- [80] O. Nagy. Excitation cross-sections of N₂⁺ molecular ion by electron impact and the vibrational energy levels of the three target states. *Chemical Physics*, 286(1):109–114, January 2003.
- [81] R.C. Flagan and J.P. Appleton. Excitation mechanisms of the nitrogen First-Positive and First-Negative radiation at high temperature. *Journal of Chemical Physics*, 56(3):1163–1173, February 1972.

A

Diatomic collisional-radiative models

In this Appendix the collisional-radiative models for diatomic molecules implemented in this work are presented. The collisional rates are expressed in generalised Arrhenius form (see Equation 3.71), and the radiative transitions are expressed via the average radiative transition probability for the system. If the radiative transition rates are not provided in the literature, they are calculated from the electronic-vibration transition moments via Equation 3.109. The collisional-radiative models for C_2 , CN, CO, N_2 and N_2^+ are presented in § A.1 to A.5 respectively.

A.1 Collisional-radiative model for C₂

The implemented collisional-radiative model for C₂ is from Zalogin [66], and is presented in Table A.1.

Reaction	A (cm ³ /s)	n	E_a (K)	Source
<i>Electron impact excitation</i>				
$C_2(X^1\Sigma_g^+) + e^- \rightleftharpoons C_2(d^3\Pi_g) + e^-$	1.3×10^{-8}	0.00	28,807	Zalogin [66]
<i>Heavy particle impact excitation</i>				
$C_2(X^1\Sigma_g^+) + M \rightleftharpoons C_2(d^3\Pi_g) + M$	8.6×10^{-11}	0.00	28,807	Zalogin [66]
<i>Radiative transitions</i>				
Reaction	A (s ⁻¹)			Source
$C_2(X^1\Sigma_g^+) \Rightarrow C_2(d^3\Pi_g) + h\nu$	9.3×10^6			Zalogin [66]

Table A.1: Implemented collisional-radiative model for C₂.

A.2 Collisional-radiative model for CN

The implemented collisional-radiative model for CN uses the electron impact excitation cross sections and electron impact dissociation coefficients compiled by Park [56], and is presented in Table A.2. The electron impact excitation rate coefficients have been calculated from the cross sections given in Reference [56] via Equations 3.77 to 3.83, and then curve-fitted to the generalised Arrhenius expression. Heavy particle impact processes are omitted and the radiative transition probabilities have been calculated via Equation 3.109.

Reaction	A (cm ³ /s)	n	E_a (K)	Source
<i>Electron impact excitation</i>				
$\text{CN}(X^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(A^2\Pi) + e^-$	1.9×10^{-9}	0.24	13,302	Same as N ₂
$\text{CN}(X^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(B^2\Sigma^+) + e^-$	3.1×10^{-7}	-0.25	37,052	Same as N ₂
$\text{CN}(X^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(a^4\Sigma^+) + e^-$	1.6×10^{-14}	1.09	46,616	Same as N ₂
$\text{CN}(X^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(D^2\Pi) + e^-$	3.0×10^{-9}	0.16	78,393	Same as N ₂
$\text{CN}(A^2\Pi) + e^- \rightleftharpoons \text{CN}(B^2\Sigma^+) + e^-$	1.0×10^{-5}	-0.22	23,750	Huo <i>et al.</i> [70]
$\text{CN}(A^2\Pi) + e^- \rightleftharpoons \text{CN}(a^4\Sigma^+) + e^-$	5.9×10^{-8}	0.18	33,314	Huo <i>et al.</i> [70]
$\text{CN}(A^2\Pi) + e^- \rightleftharpoons \text{CN}(D^2\Pi) + e^-$	2.0×10^{-5}	-0.29	65,091	Huo <i>et al.</i> [70]
$\text{CN}(B^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(a^4\Sigma^+) + e^-$	5.2×10^{-12}	0.70	9,564	Huo <i>et al.</i> [70]
$\text{CN}(B^2\Sigma^+) + e^- \rightleftharpoons \text{CN}(D^2\Pi) + e^-$	4.6×10^{-8}	0.22	41,341	Huo <i>et al.</i> [70]
$\text{CN}(a^4\Sigma^+) + e^- \rightleftharpoons \text{CN}(D^2\Pi) + e^-$	6.9×10^{-5}	-0.40	31,777	Huo <i>et al.</i> [70]
<i>Electron impact dissociation</i>				
$\text{CN}(X^2\Sigma^+) + e^- \rightleftharpoons \text{C} + \text{N} + e^-$	1.3×10^{15}	0.43	88,966	Park [56]
$\text{CN}(A^2\Pi) + e^- \rightleftharpoons \text{C} + \text{N} + e^-$	5.9×10^{15}	0.46	75,564	Park [56]
$\text{CN}(B^2\Sigma^+) + e^- \rightleftharpoons \text{C} + \text{N} + e^-$	2.4×10^{15}	0.55	51,576	Park [56]
$\text{CN}(a^4\Sigma^+) + e^- \rightleftharpoons \text{C} + \text{N} + e^-$	1.6×10^{15}	0.60	41,890	Park [56]
$\text{CN}(D^2\Pi) + e^- \rightleftharpoons \text{C} + \text{N} + e^-$	6.3×10^{15}	0.92	9,964	Park [56]
<i>Radiative transitions</i>				
Reaction	A (s ⁻¹)	Source		
$\text{CN}(A^2\Pi) \Rightarrow \text{CN}(B^2\Sigma^+) + h\nu$	9.3×10^5	This work (Eq. 3.109)		
$\text{CN}(X^2\Sigma^+) \Rightarrow \text{CN}(A^2\Pi) + h\nu$	1.7×10^6	This work (Eq. 3.109)		
$\text{CN}(X^2\Sigma^+) \Rightarrow \text{CN}(B^2\Sigma^+) + h\nu$	2.0×10^8	This work (Eq. 3.109)		

Table A.2: Implemented collisional-radiative model for CN.

A.3 Collisional-radiative model for CO

The implemented collisional-radiative model for CO uses the electron impact excitation cross sections and electron impact dissociation coefficients compiled by Park [56], and is presented in Table A.3. The electron impact excitation rate coefficients have been calculated from the cross sections given in Reference [56] via Equations 3.77 to 3.83, and then curve-fitted to the generalised Arrhenius expression. Heavy particle impact processes are omitted and the radiative transition probabilities have been calculated via Equation 3.109.

Reaction	A (cm ³ /s)	n	E_a (K)	Source
<i>Electron impact excitation</i>				
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{CO}(a^2\Pi) + e^-$	3.9×10^{-6}	-0.34	70,049	Zobel <i>et al.</i> [71]
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{CO}(a'^3\Sigma^+) + e^-$	1.2×10^{-6}	-0.28	80,320	Zobel <i>et al.</i> [71]
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{CO}(d^3\Delta) + e^-$	3.5×10^{-7}	-0.27	87,938	Zobel <i>et al.</i> [71]
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{CO}(e^3\Sigma^-) + e^-$	5.3×10^{-7}	-0.22	92,412	Morgan <i>et al.</i> [72]
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{CO}(A^1\Pi) + e^-$	1.3×10^{-5}	-0.43	93,629	Olszewski <i>et al.</i> [73]
$\text{CO}(a^2\Pi) + e^- \rightleftharpoons \text{CO}(a'^3\Sigma^+) + e^-$	5.6×10^{-7}	-0.07	10,271	Huo <i>et al.</i> [70]
$\text{CO}(a^2\Pi) + e^- \rightleftharpoons \text{CO}(d^3\Delta) + e^-$	5.4×10^{-6}	-0.31	17,889	Huo <i>et al.</i> [70]
$\text{CO}(a^2\Pi) + e^- \rightleftharpoons \text{CO}(e^3\Sigma^-) + e^-$	9.4×10^{-6}	-0.36	22,364	Huo <i>et al.</i> [70]
$\text{CO}(a^2\Pi) + e^- \rightleftharpoons \text{CO}(A^1\Pi) + e^-$	2.1×10^{-6}	-0.25	23,580	Huo <i>et al.</i> [70]
$\text{CO}(a'^3\Sigma^+) + e^- \rightleftharpoons \text{CO}(d^3\Delta) + e^-$	8.8×10^{-7}	-0.14	7,618	Huo <i>et al.</i> [70]
$\text{CO}(a'^3\Sigma^+) + e^- \rightleftharpoons \text{CO}(e^3\Sigma^-) + e^-$	8.2×10^{-6}	-0.38	12,092	Huo <i>et al.</i> [70]
$\text{CO}(a'^3\Sigma^+) + e^- \rightleftharpoons \text{CO}(A^1\Pi) + e^-$	1.3×10^{-6}	-0.27	13,309	Huo <i>et al.</i> [70]
$\text{CO}(d^3\Delta) + e^- \rightleftharpoons \text{CO}(e^3\Sigma^-) + e^-$	1.0×10^{-7}	0.07	4,475	Huo <i>et al.</i> [70]
$\text{CO}(d^3\Delta) + e^- \rightleftharpoons \text{CO}(A^1\Pi) + e^-$	4.6×10^{-8}	0.09	5,691	Huo <i>et al.</i> [70]
$\text{CO}(e^3\Sigma^-) + e^- \rightleftharpoons \text{CO}(A^1\Pi) + e^-$	1.8×10^{-9}	0.36	1,216	Huo <i>et al.</i> [70]
<i>Electron impact dissociation</i>				
$\text{CO}(X^1\Sigma^+) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	2.1×10^{14}	0.37	129,271	Park [56]
$\text{CO}(a^2\Pi) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	3.1×10^{15}	0.52	58,742	Park [56]
$\text{CO}(a'^3\Sigma^+) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	3.3×10^{15}	0.56	48,352	Park [56]
$\text{CO}(d^3\Delta) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	4.2×10^{15}	0.61	40,635	Park [56]
$\text{CO}(e^3\Sigma^-) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	3.9×10^{15}	0.64	36,098	Park [56]
$\text{CO}(A^1\Pi) + e^- \rightleftharpoons \text{C} + \text{O} + e^-$	4.8×10^{15}	0.64	34,864	Park [56]
<i>Radiative transitions</i>				
Reaction	A (s ⁻¹)			Source
$\text{CO}(X^1\Sigma^+) \Rightarrow \text{CO}(A^1\Pi) + h\nu$	1.4×10^9			This work (Eq. 3.109)

Table A.3: Implemented collisional-radiative model for CO.

A.4 Collisional-radiative model for N₂

The implemented collisional-radiative model for N₂ is that compiled by Johnston [12], and is presented in Table A.4.

Reaction	A (cm ³ /s)	n	E_a (K)	Source
<i>Electron impact excitation</i>				
$N_2(X^1\Sigma_g^+) + e^- \rightleftharpoons N_2(A^3\Sigma_u^+) + e^-$	4.0×10^{-9}	0.1	71,610	Chernyi [74]
$N_2(X^1\Sigma_g^+) + e^- \rightleftharpoons N_2(B^3\Pi_g) + e^-$	4.6×10^{-8}	-0.1	85,740	Chernyi [74]
$N_2(X^1\Sigma_g^+) + e^- \rightleftharpoons N_2(C^3\Pi_u) + e^-$	3.8×10^{-9}	0.1	127,900	Capitelli [75]
$N_2(A^3\Sigma_u^+) + e^- \rightleftharpoons N_2(B^3\Pi_g) + e^-$	5.0×10^{-9}	0.0	13,495	Chernyi [74]
$N_2(B^3\Pi_g) + e^- \rightleftharpoons N_2(C^3\Pi_u) + e^-$	2.9×10^{-9}	0.28	46,655	Teulet [55]
<i>Electron impact dissociation</i>				
$N_2(X^1\Sigma_g^+) + e^- \rightleftharpoons N + N + e^-$	4.11×10^{-33}	6.16	113,263	Teulet [55]
$N_2(A^3\Sigma_u^+) + e^- \rightleftharpoons N + N + e^-$	6.61×10^{-20}	2.98	41,669	Teulet [55]
$N_2(B^3\Pi_g) + e^- \rightleftharpoons N + N + e^-$	4.50×10^{-23}	3.73	55,586	Teulet [55]
$N_2(C^3\Pi_u) + e^- \rightleftharpoons N + N + e^-$	5.14×10^{-21}	3.27	12,892	Teulet [55]
<i>Heavy particle impact excitation</i>				
$N_2(X^1\Sigma_g^+) + N_2 \rightleftharpoons N_2(A^3\Sigma_u^+) + N_2$	1.83×10^{-12}	-0.5	71,600	Kurochkin [76]
$N_2(A^3\Sigma_u^+) + N_2 \rightleftharpoons N_2(B^3\Pi_g) + N_2$	1.99×10^{-11}	0.0	13,495	Chernyi [74]
$N_2(B^3\Pi_g) + N_2 \rightleftharpoons N_2(C^3\Pi_u) + N_2$	8.47×10^{-11}	0.0	42,476	Fresnet [77]
Reaction	A (s ⁻¹)	Source		
$N_2(B^3\Pi_g) \Rightarrow N_2(A^3\Sigma_u^+) + h\nu$	1.4×10^5	Chernyi [74]		
$N_2(C^3\Pi_u) \Rightarrow N_2(B^3\Pi_g) + h\nu$	2.6×10^7	Pancheshnyi [78]		

Table A.4: Implemented collisional-radiative model for N₂.

A.5 Collisional-radiative model for N_2^+

The implemented collisional-radiative model for N_2^+ is that compiled by Johnston [12], and is presented in Table A.5.

Reaction	A (cm^3/s)	n	E_a (K)	Source
<i>Electron impact excitation</i>				
$N_2^+ (X^2\Sigma_g^+) + e^- \rightleftharpoons N_2^+ (A^2\Pi_u) + e^-$	7.1×10^{-11}	0.0	13,300	Gorelov [79]
$N_2^+ (X^2\Sigma_g^+) + e^- \rightleftharpoons N_2^+ (B^2\Sigma_u^+) + e^-$	2.0×10^{-11}	0.73	36,649	Nagy [80]
$N_2^+ (X^2\Sigma_g^+) + e^- \rightleftharpoons N_2^+ (C^2\Sigma_u^+) + e^-$	6.6×10^{-9}	0.41	85,038	Teulet [55]
$N_2^+ (A^2\Pi_u) + e^- \rightleftharpoons N_2^+ (B^2\Sigma_u^+) + e^-$	1.0×10^{-9}	0.0	23,500	Gorelov [79]
$N_2^+ (A^2\Pi_u) + e^- \rightleftharpoons N_2^+ (C^2\Sigma_u^+) + e^-$	1.3×10^{-7}	0.11	78,403	Teulet [55]
$N_2^+ (B^2\Sigma_u^+) + e^- \rightleftharpoons N_2^+ (C^2\Sigma_u^+) + e^-$	3.9×10^{-9}	0.34	49,622	Teulet [55]
<i>Electron impact dissociation</i>				
$N_2^+ (X^2\Sigma_g^+) + e^- \rightleftharpoons N^+ + N + e^-$	8.02×10^{-31}	5.54	101,117	Teulet [55]
$N_2^+ (A^2\Pi_u) + e^- \rightleftharpoons N^+ + N + e^-$	8.27×10^{-26}	4.38	88,142	Teulet [55]
$N_2^+ (B^2\Sigma_u^+) + e^- \rightleftharpoons N^+ + N + e^-$	2.58×10^{-32}	5.81	64,328	Teulet [55]
$N_2^+ (C^2\Sigma_u^+) + e^- \rightleftharpoons N^+ + N + e^-$	1.31×10^{-28}	4.93	35,906	Teulet [55]
<i>Heavy particle impact excitation</i>				
$N_2^+ (X^2\Sigma_g^+) + N_2 \rightleftharpoons N_2^+ (A^2\Pi_u) + N_2$	3.8×10^{-2}	-2.33	12,978	Nagy [80]
$N_2^+ (X^2\Sigma_g^+) + N_2 \rightleftharpoons N_2^+ (B^2\Sigma_u^+) + N_2$	1.9×10^{-2}	-2.33	36,600	Flagan [81]
Reaction	A (s^{-1})	Source		
$N_2^+ (A^2\Pi_u) \Rightarrow N_2^+ (X^2\Sigma_g^+) + h\nu$	6.7×10^4	Chernyi [74]		
$N_2^+ (B^2\Sigma_u^+) \Rightarrow N_2^+ (X^2\Sigma_g^+) + h\nu$	1.5×10^7	Gorelov [79]		
$N_2^+ (C^2\Sigma_u^+) \Rightarrow N_2^+ (X^2\Sigma_g^+) + h\nu$	1.4×10^7	Chernyi [74]		

Table A.5: Implemented collisional-radiative model for N_2^+ .